DESY Particle and Astroparticle Seminar
20, 21. Feb. 2018

UHECRs: Exploring the Universe at the Highest Energies

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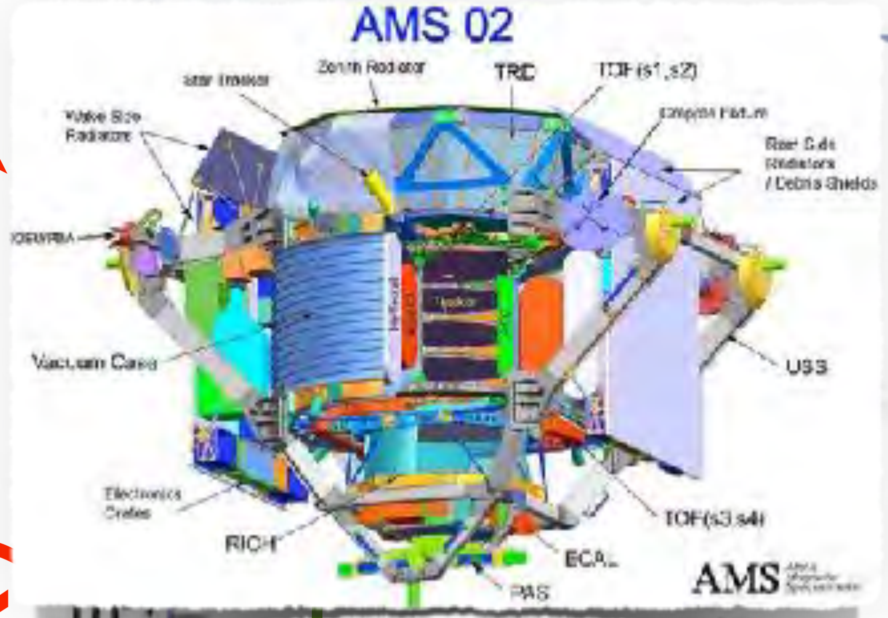
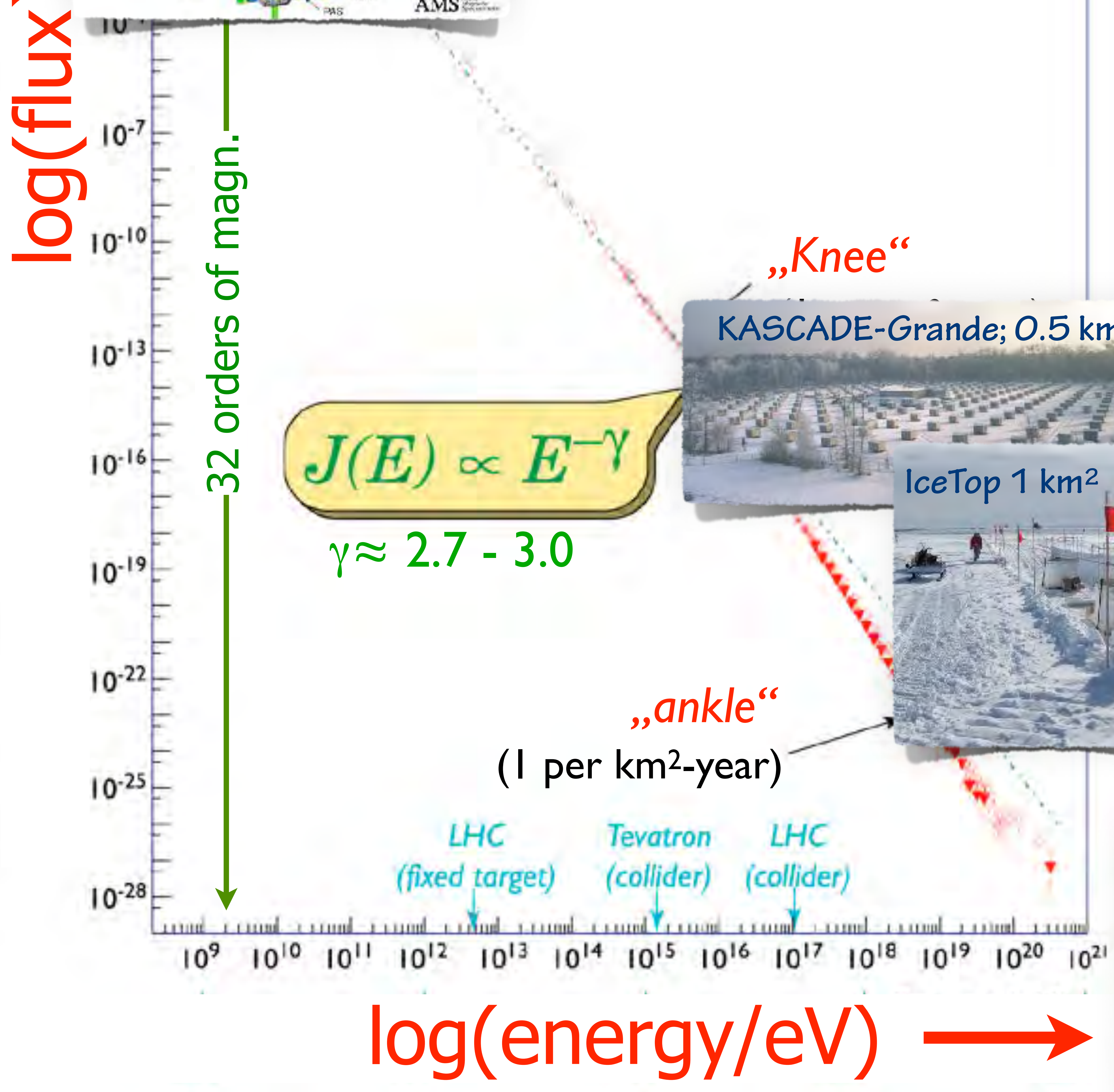
PIERRE AUGER OBSERVATORY

BERGISCHE UNIVERSITÄT WUPPERTAL

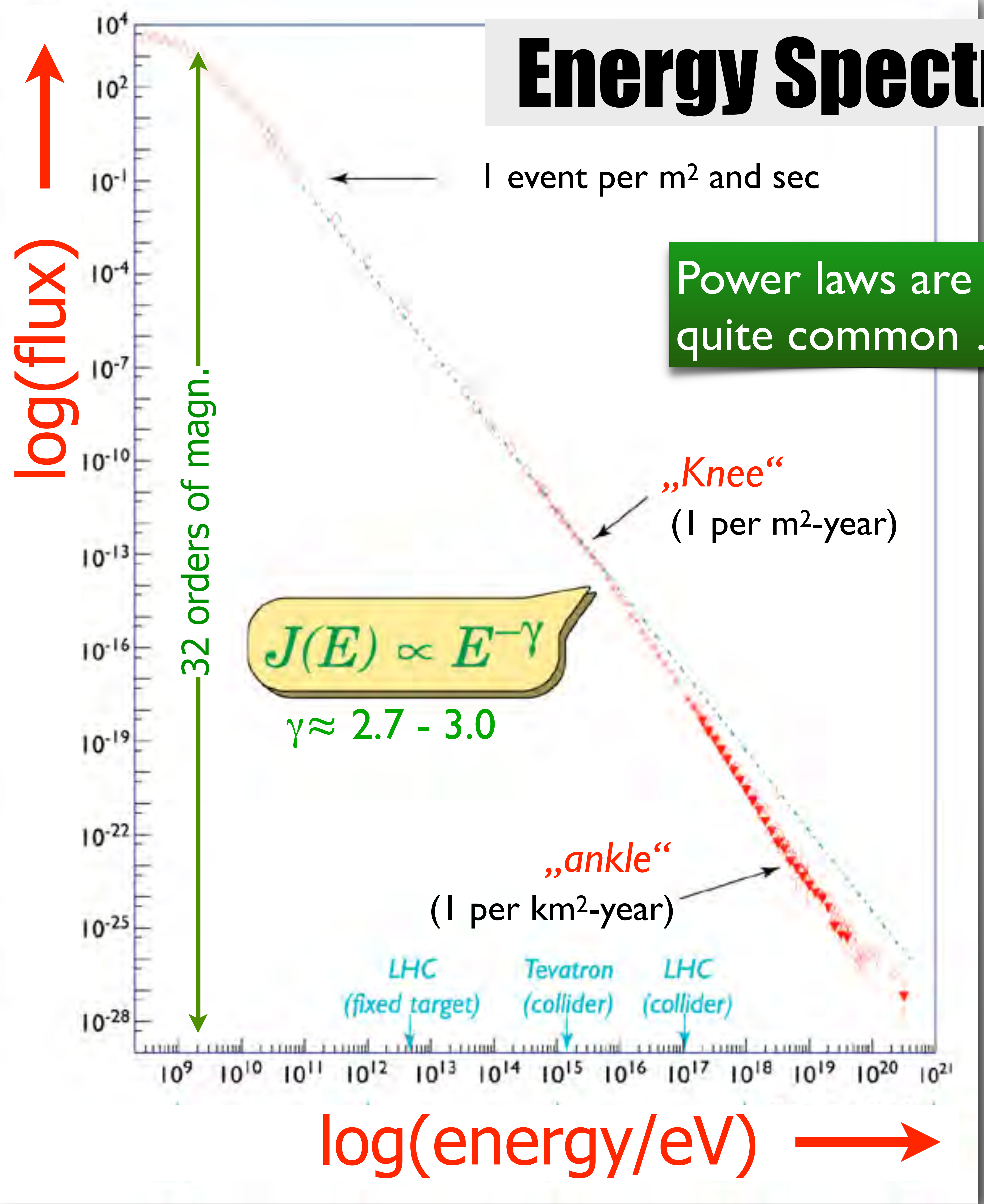
Bundesministerium für Bildung und Forschung

Energy Spectrum of Cosmic Rays

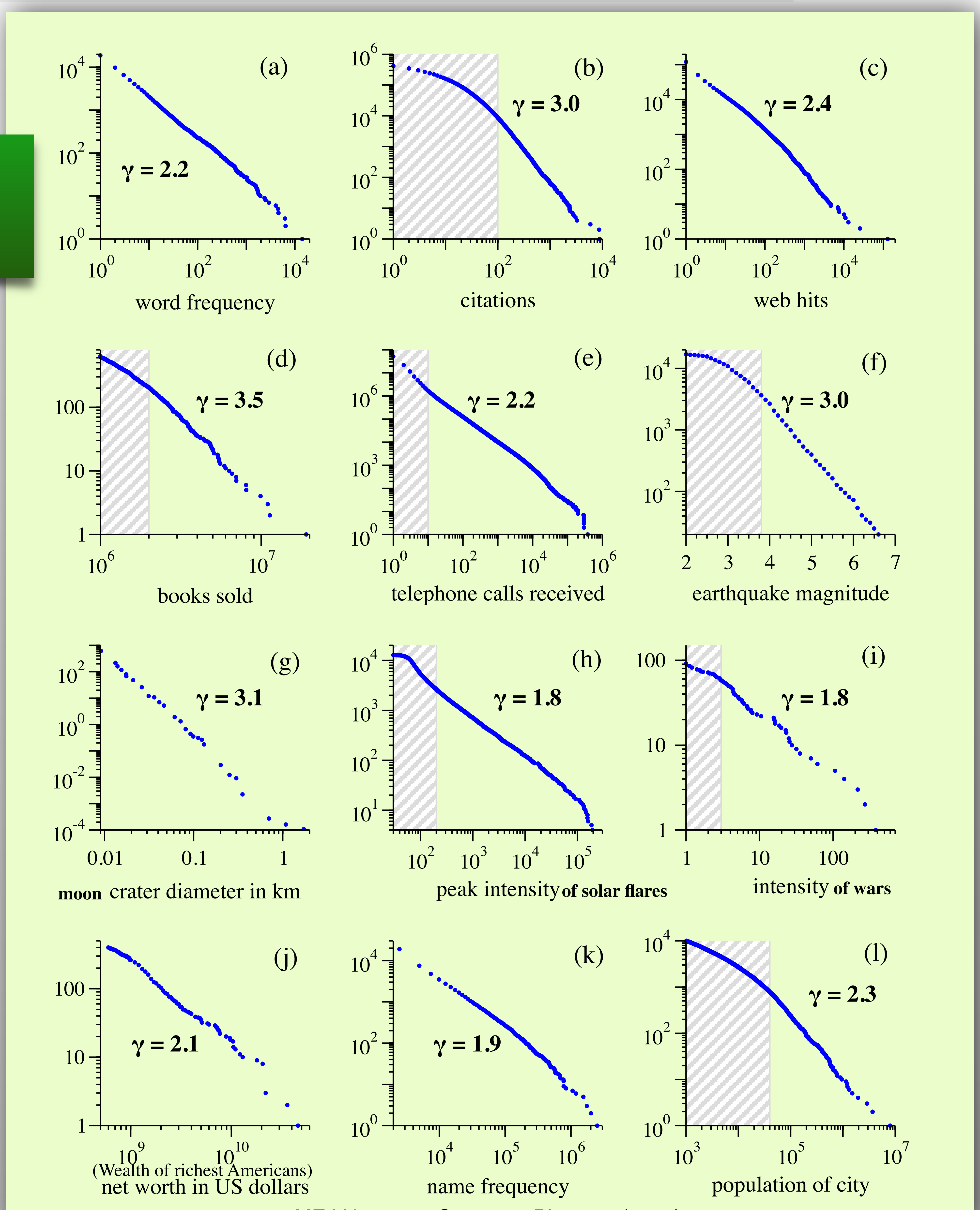




Energy Spectrum of Cosmic Rays

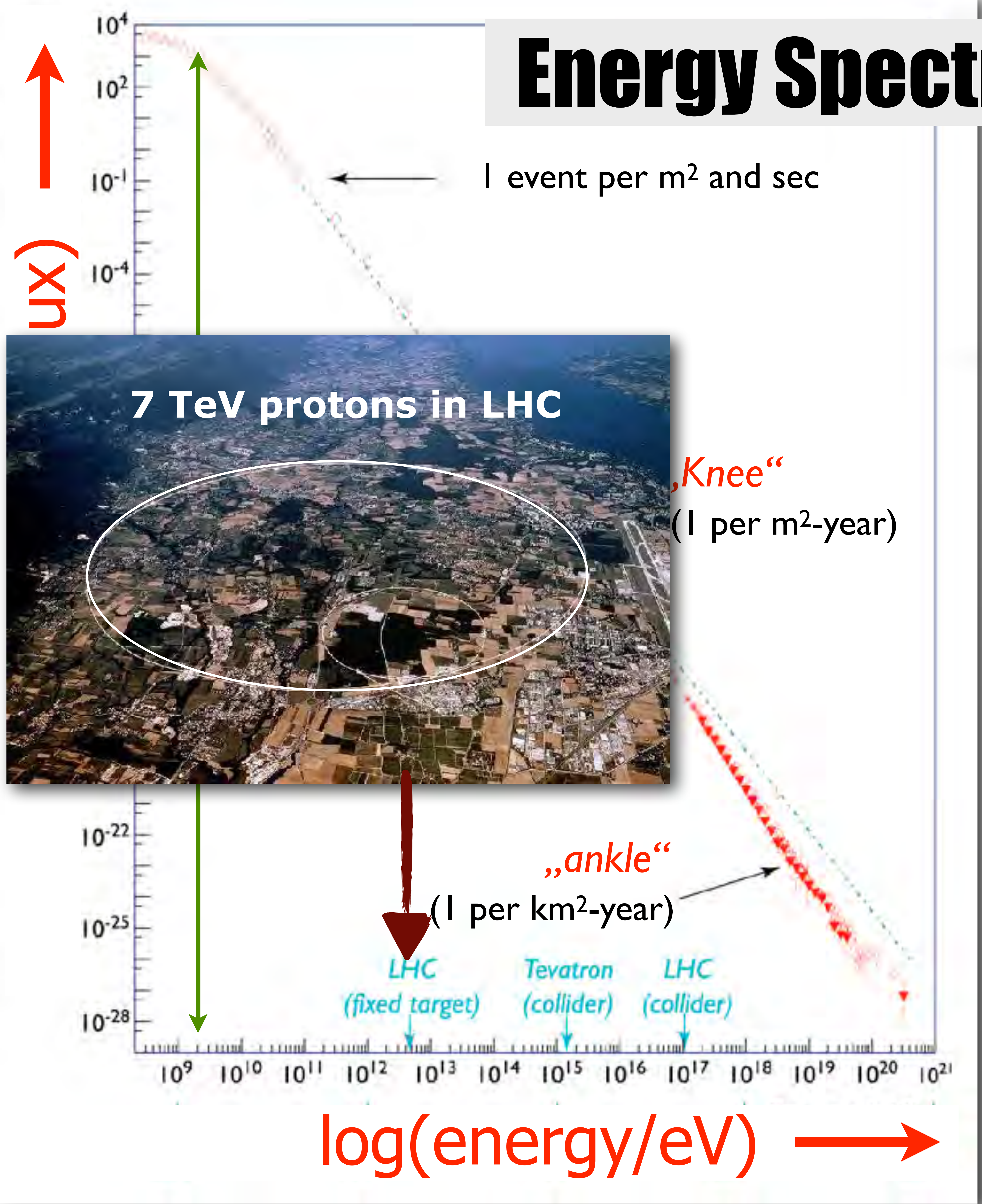


Power laws are quite common ….

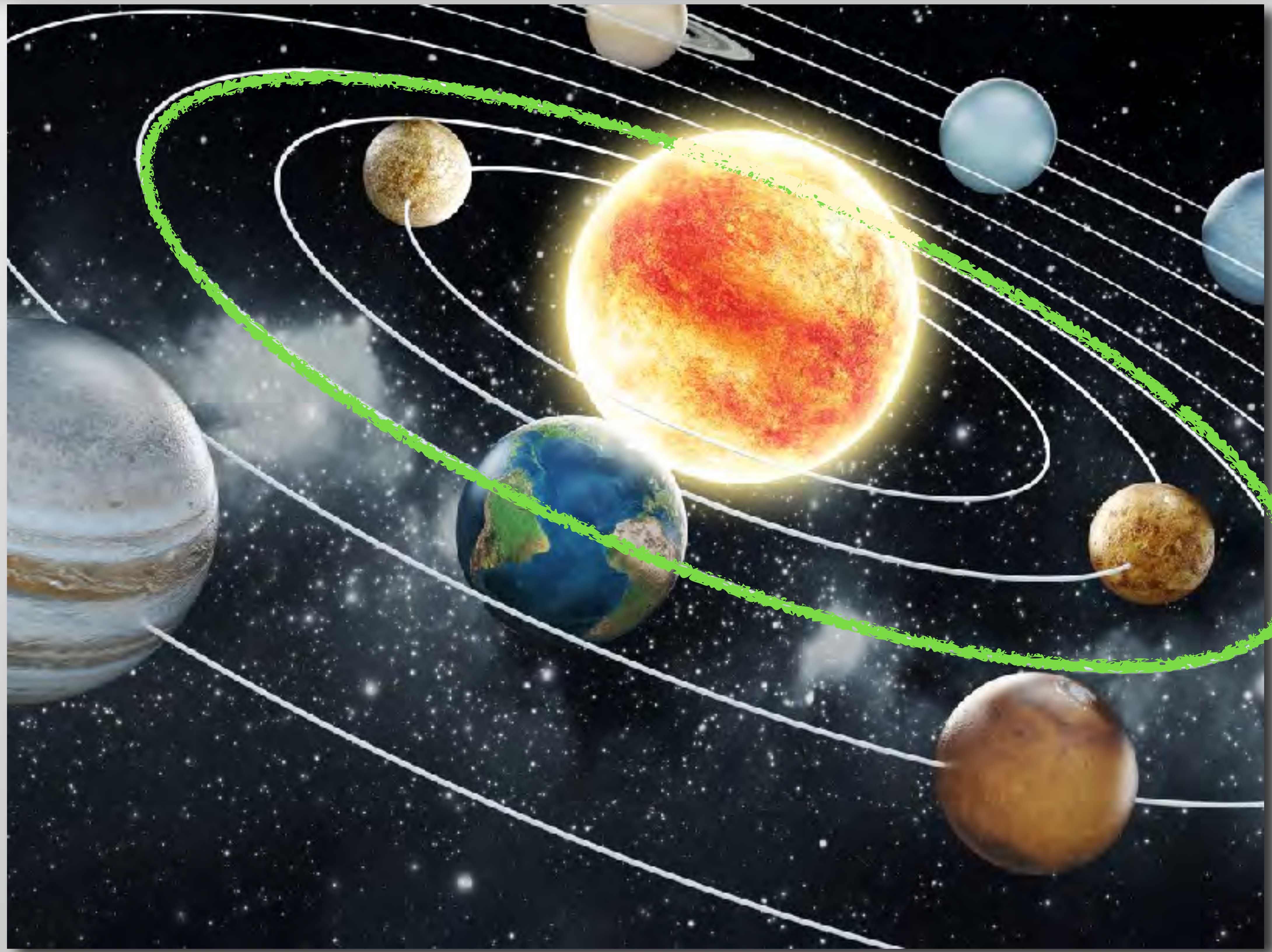


MEJ Newman, Contemp. Phys. 46 (2005) 323

Energy Spectrum of Cosmic Rays



10^{20} eV protons in LHC would require size of Earth's orbit around the Sun



Themes of UHECR Physics

- **Cosmic Particle Acceleration**
  - How and where are cosmic rays accelerated?
  - Does Nature impose any energy limits?
  - How do CRs propagate through space?
  - What is their impact on the environment?
- **Probing Extreme Environments**
  - Processes close to supermassive black holes or GRBs?
  - Processes in relativistic jets, winds and radio-lobes?
  - Exploring cosmic magnetic fields
- **Physics Frontiers – beyond the SM**
  - Lorentz invariance violation; Smoothness of Space-Time
  - Particles beyond SM ?
  - New particle physics at $\sqrt{s}$=150 TeV ?

Putative

Cosmic Particle Accelerators

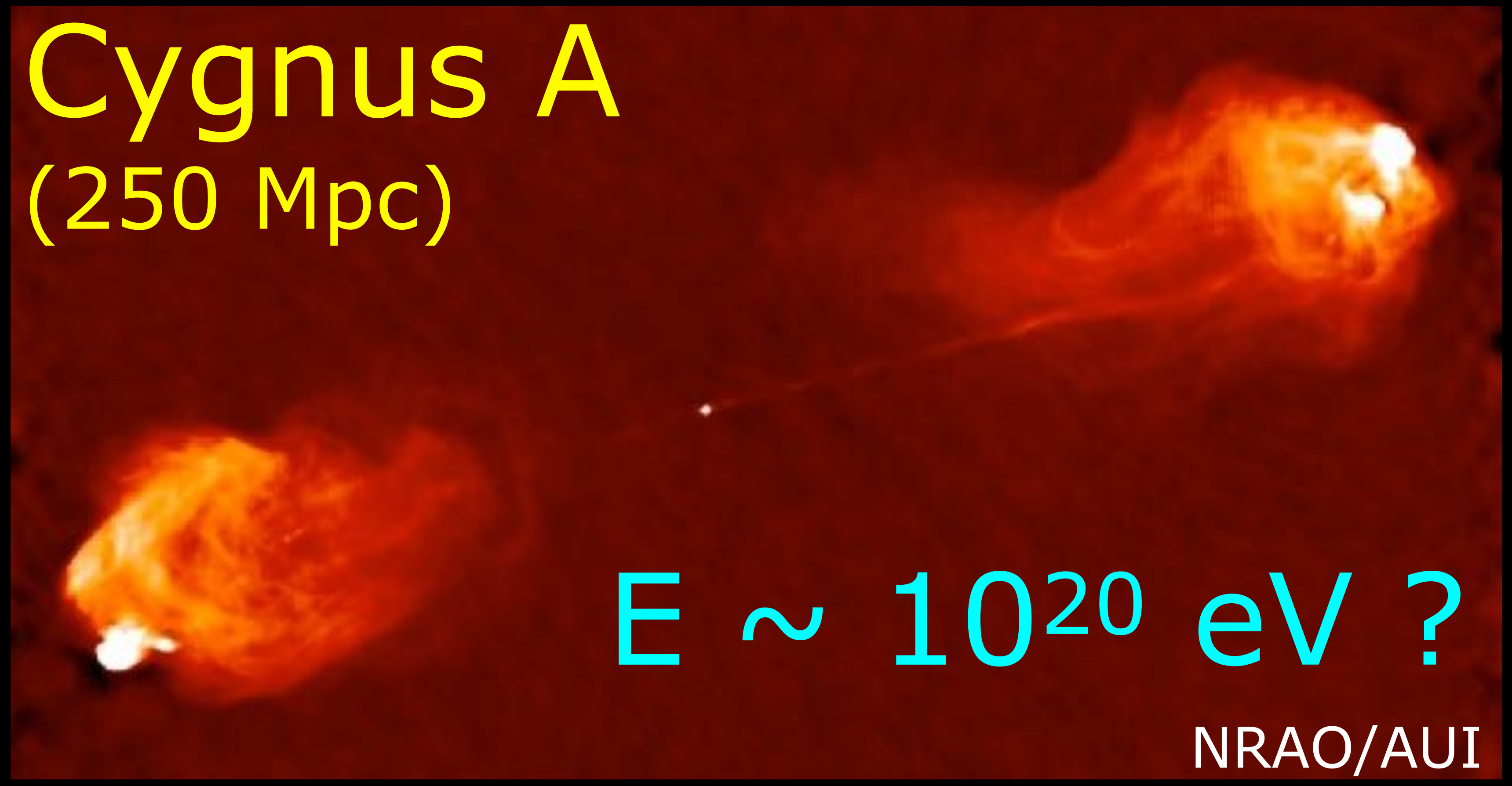
Supernova Remnants

SNR509
(50 kpc)

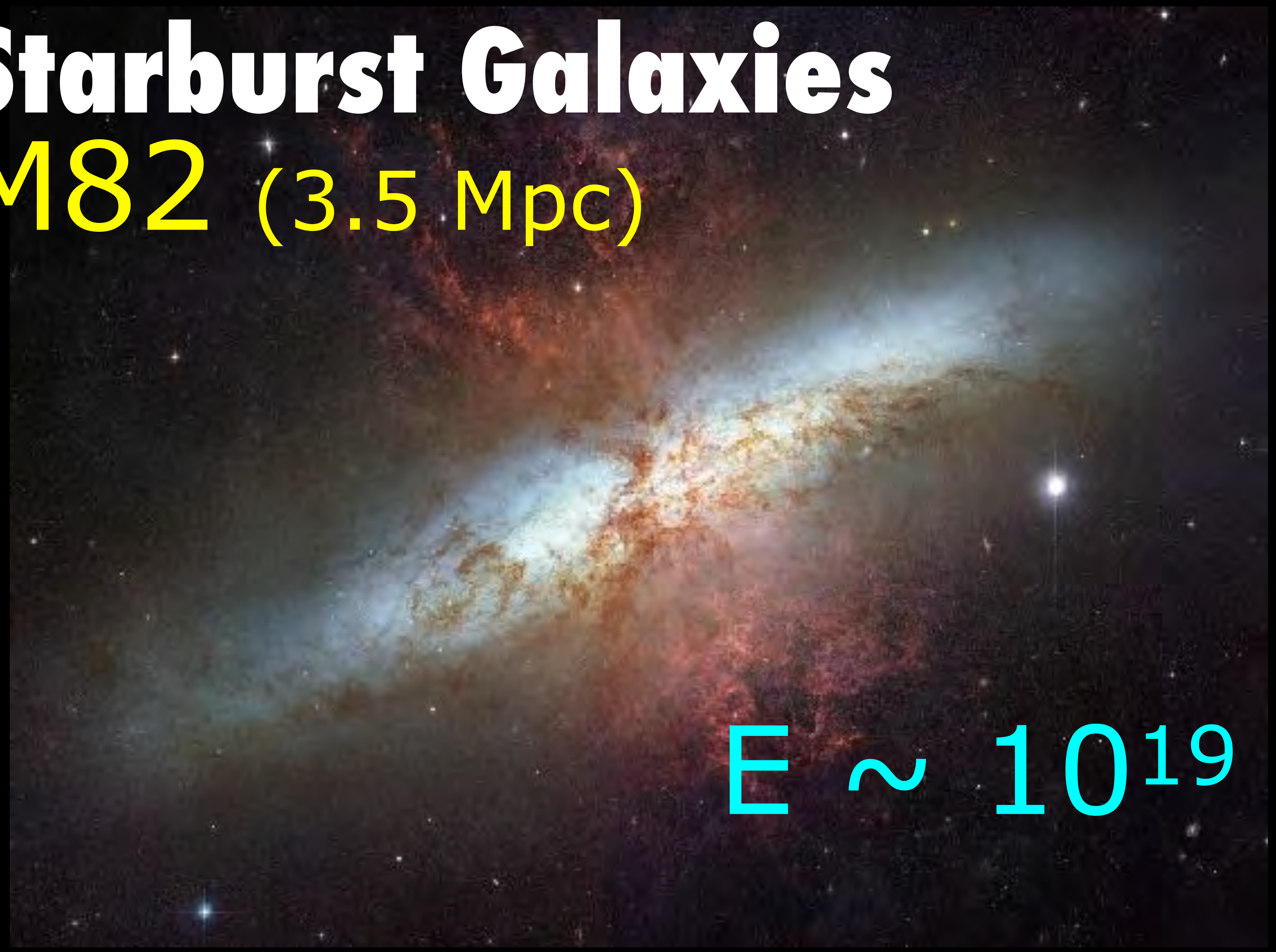
$E < 10^{15}$ eV

particle acceleration at shock waves

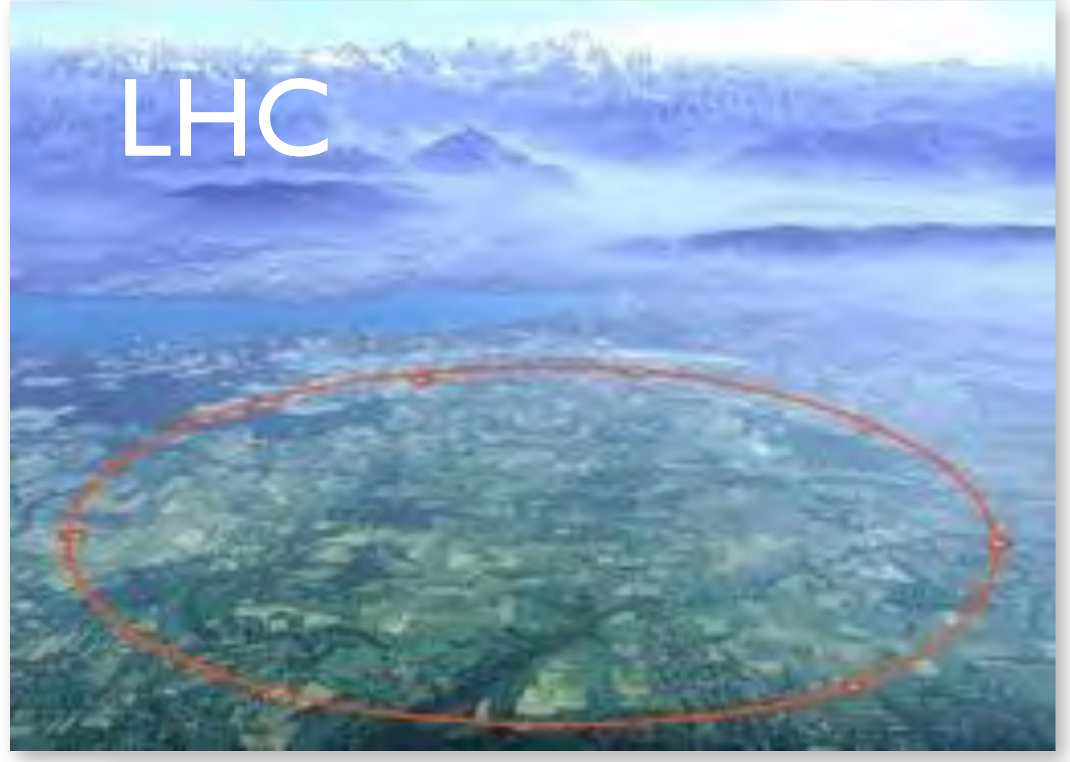
AGN and their Jets/Lobes

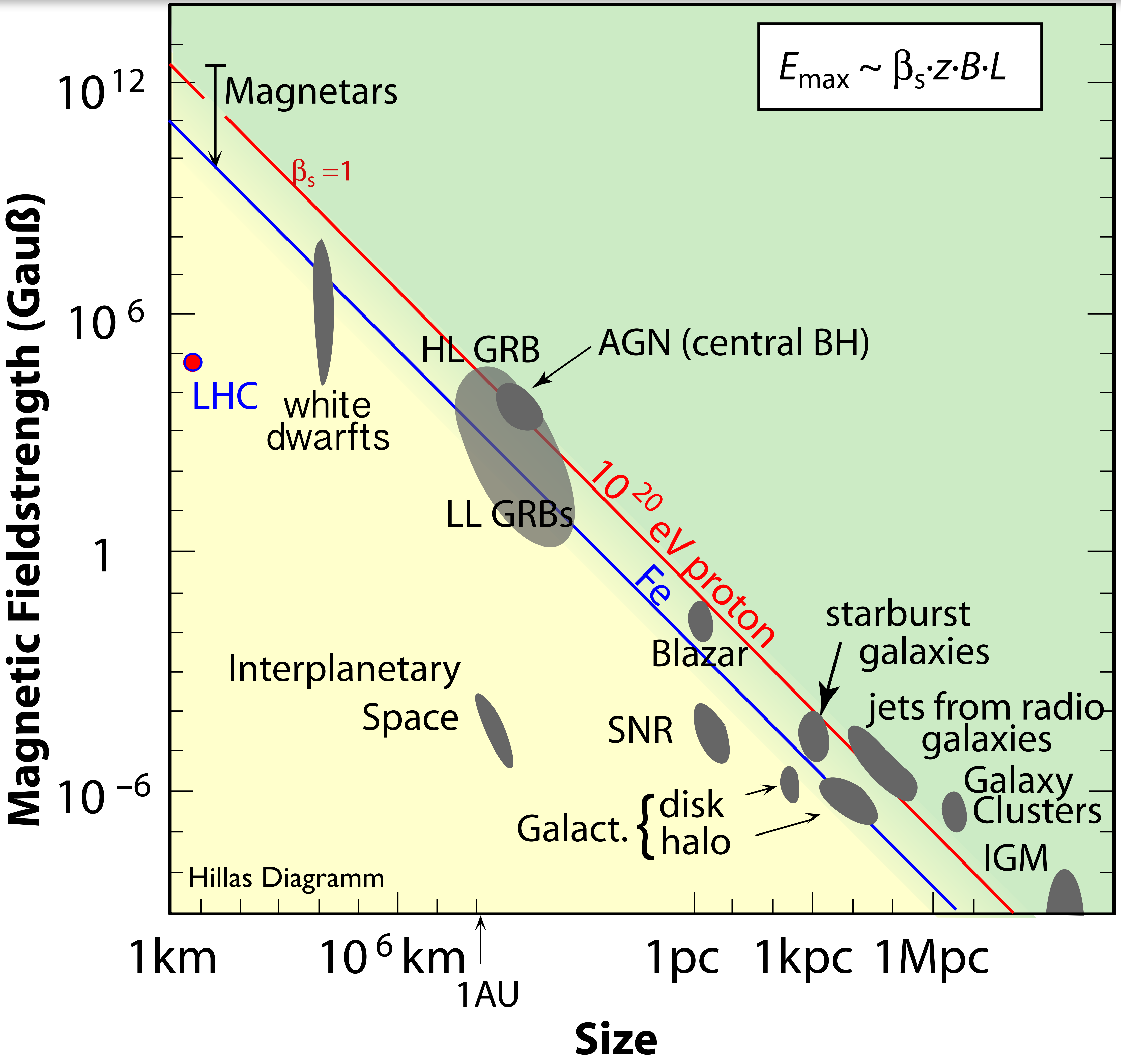


Starburst Galaxies



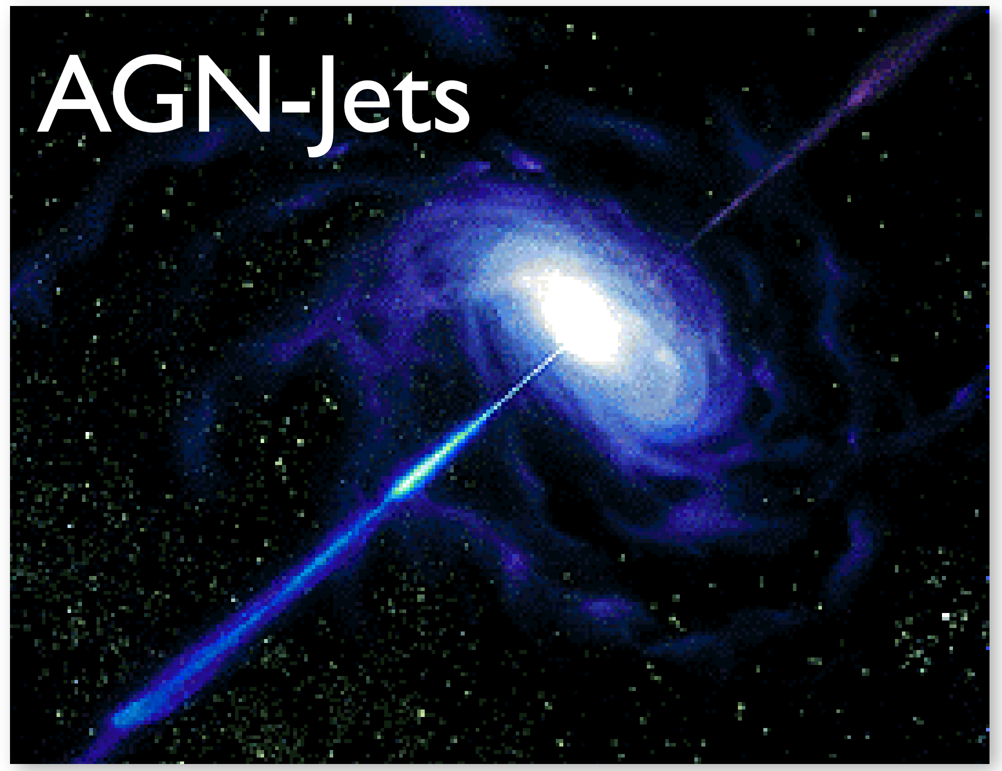
Requirement: B × Size ≳ 1 Gauss × 1 parsec





Realistic constraints more severe

- small acceleration efficiency
- synchrotron & adiabatic losses
- interactions in source region



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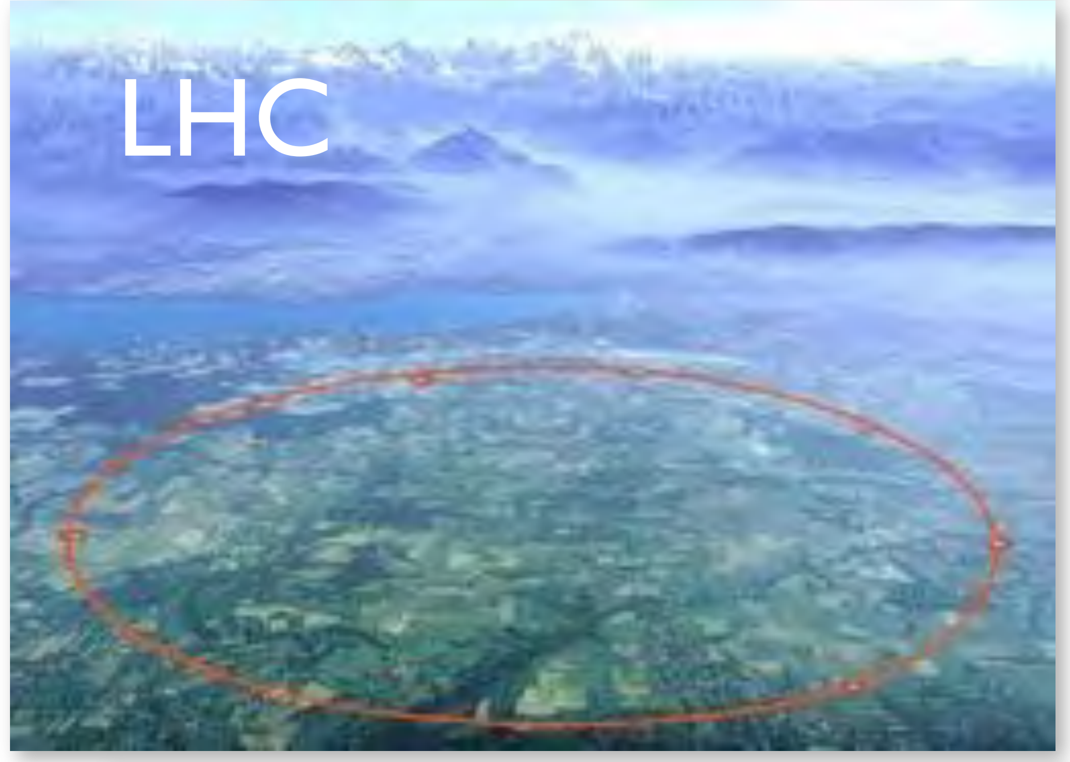
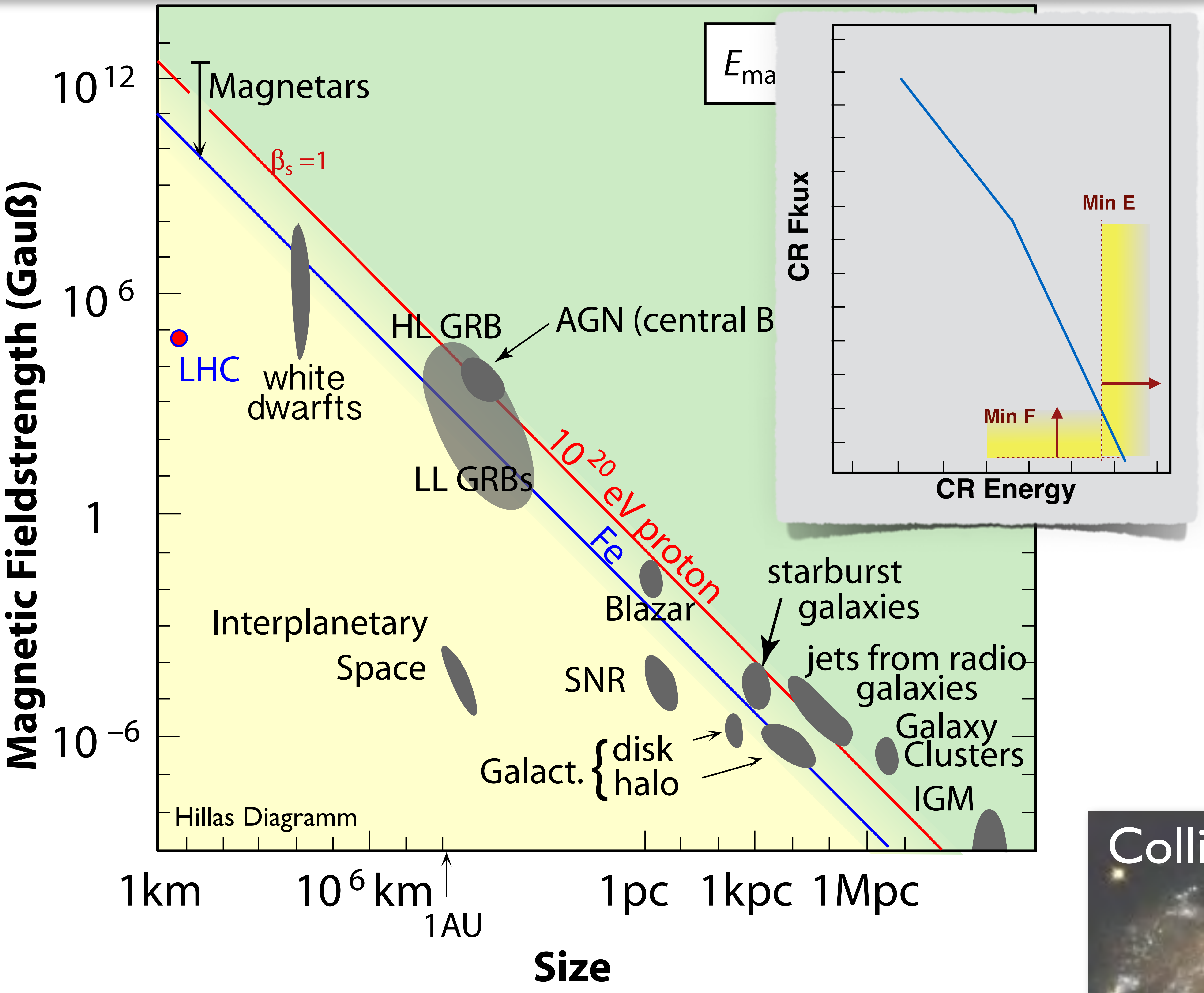
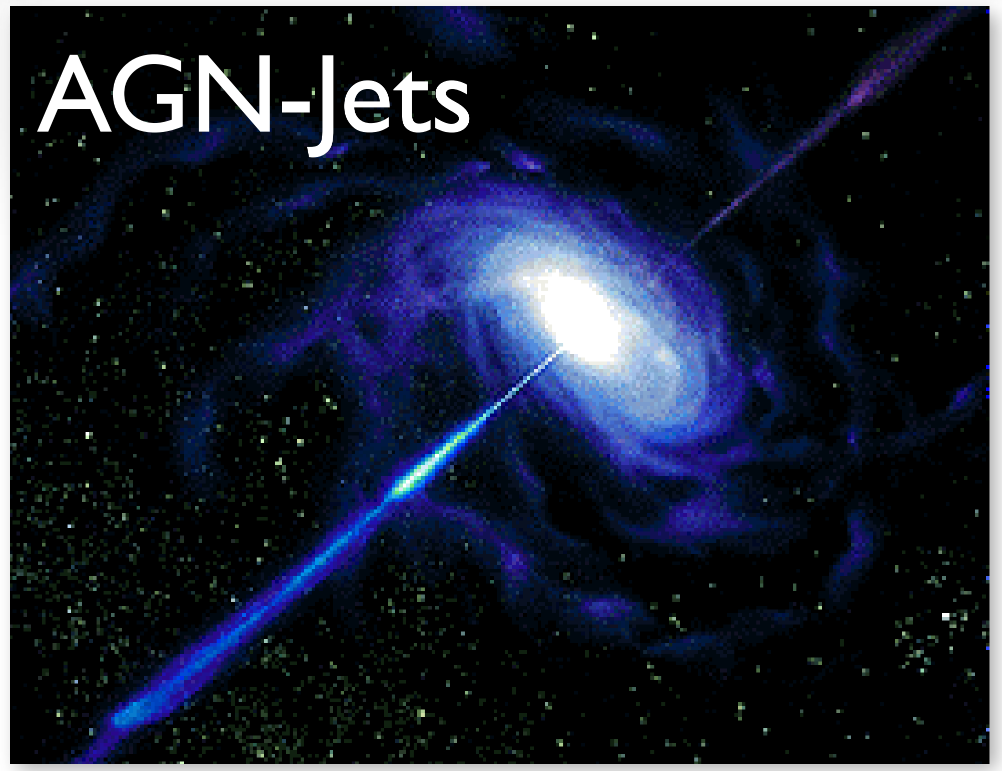


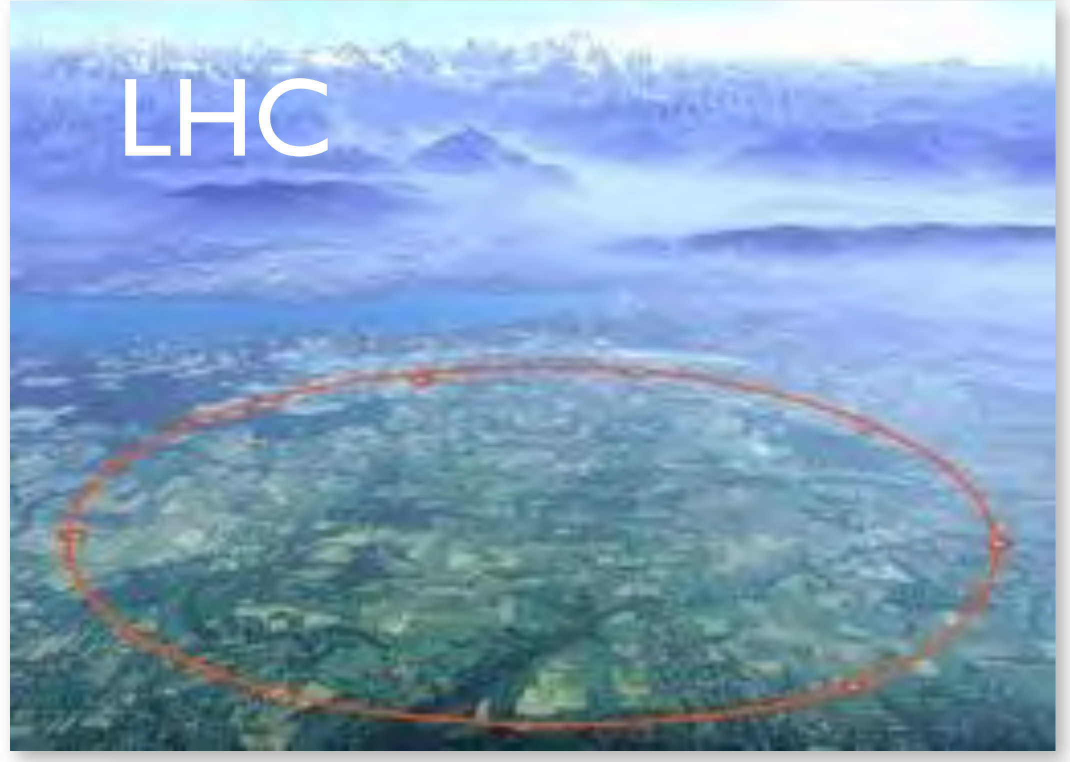


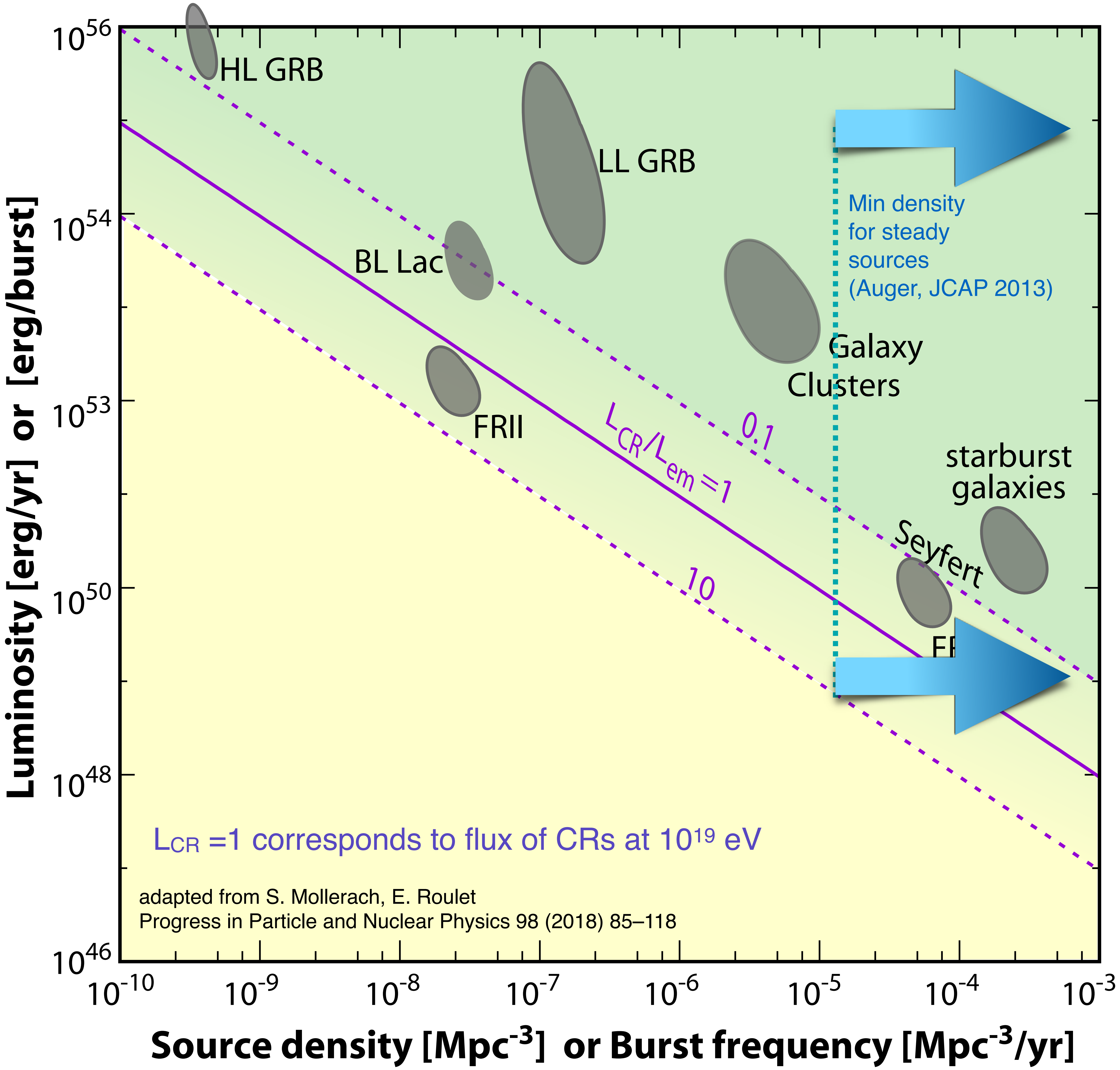


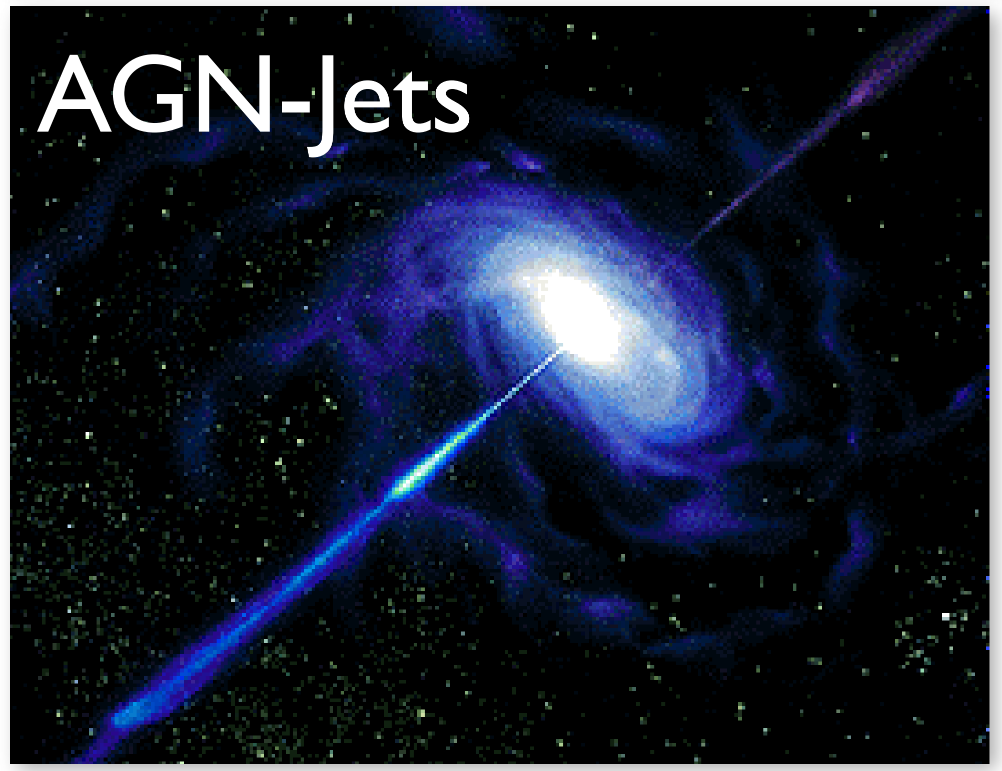
Realistic constraints more severe

- small acceleration efficiency
- synchrotron & adiabatic losses
- interactions in source region

Source Luminosity vs Source Density

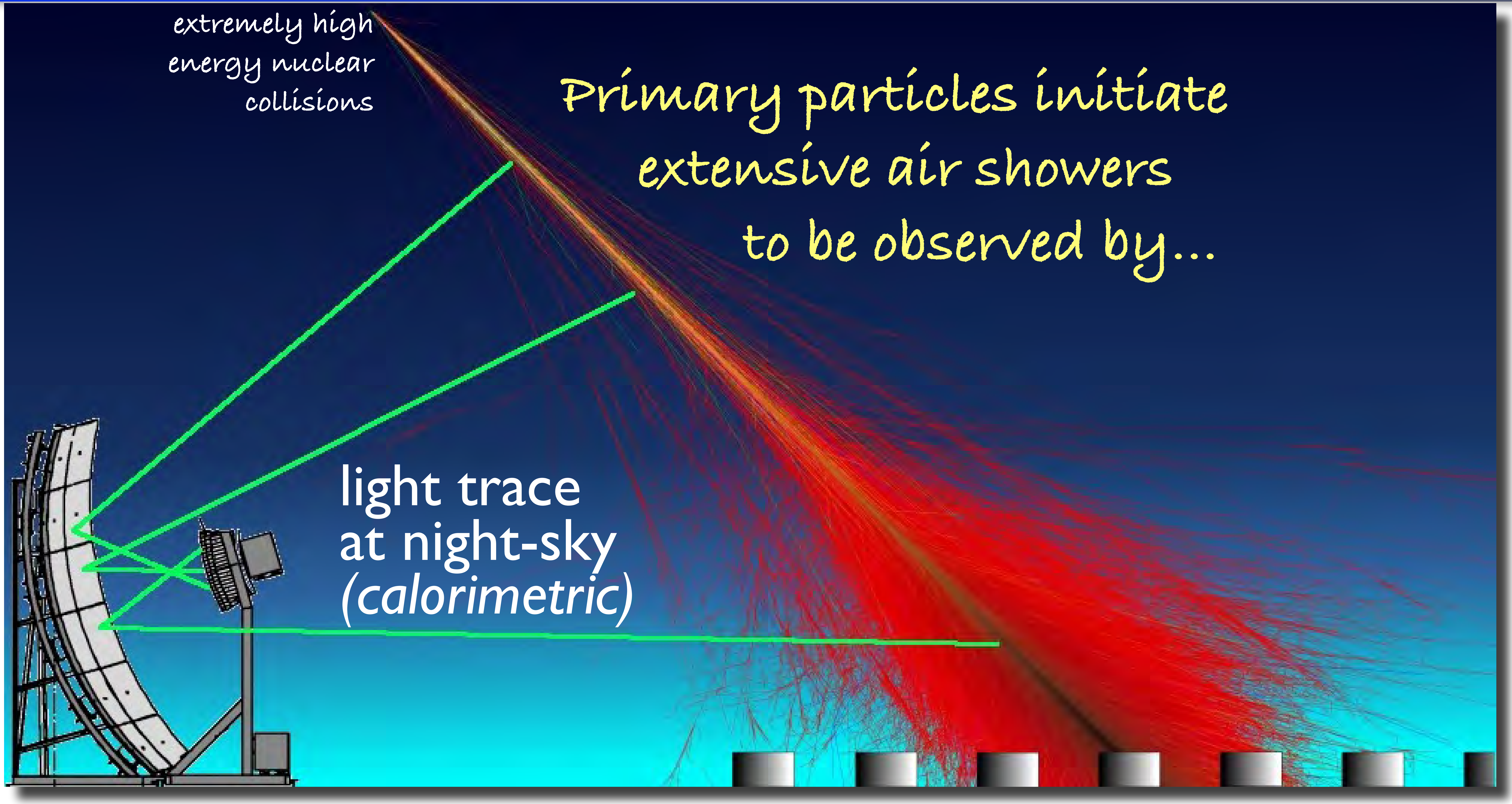






UHECR Observations

Observation of UHECR



Fluorescence light

Particle-density and
-composition at ground

Also:
Detection of Radio- & Microwave-Signals

Pierre Auger Observatory

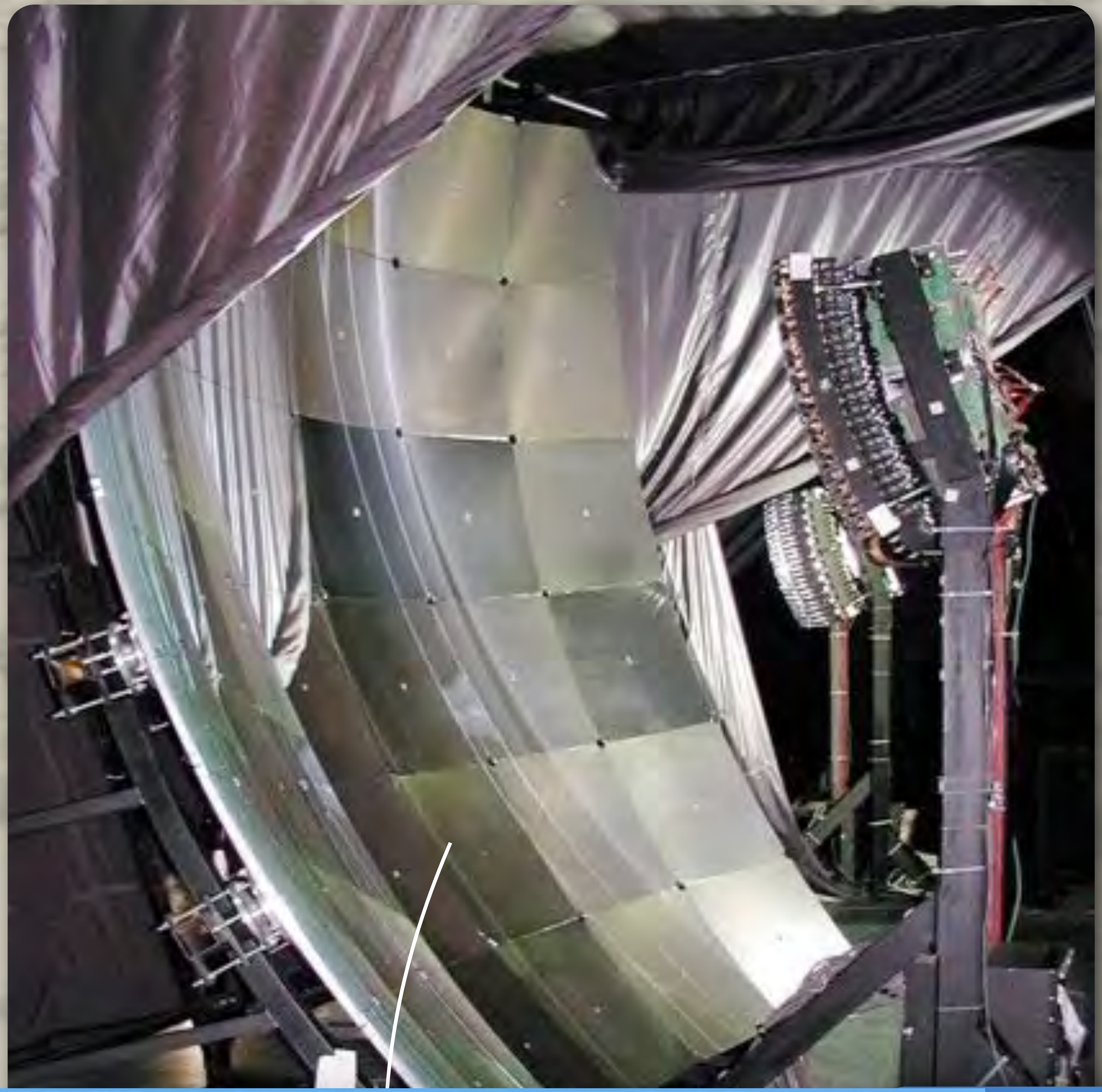
Province Mendoza, Argentina

1660 detector stations on 1.5 km grid

27 fluores. telescopes at periphery

153 radio antennas over 17 km^2

Auger Hybrid Observatory

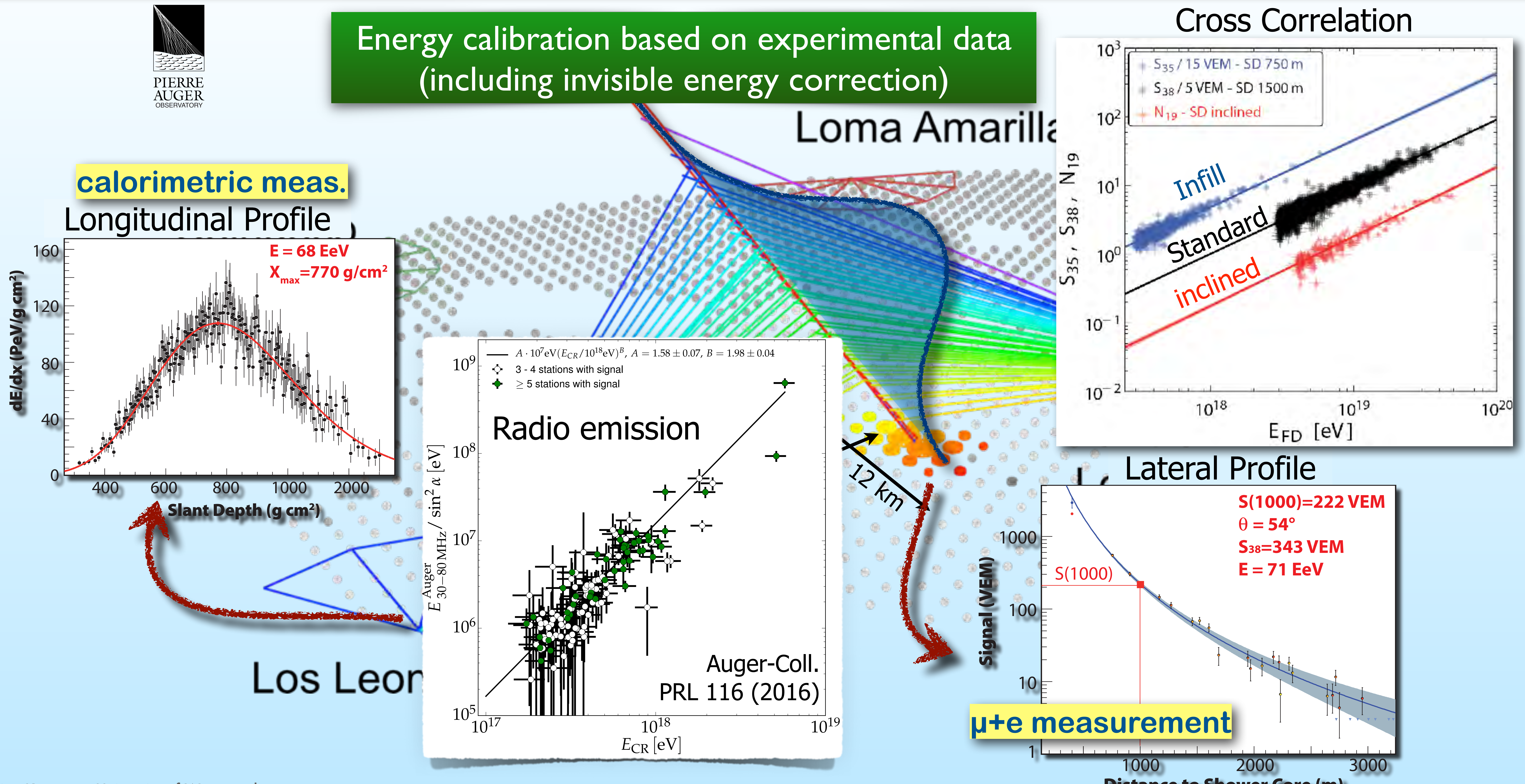


3000 km² area, Argentina
27 fluorescence telescopes plus
153 Radio Antennas

...1660 Water Cherenkov tanks

Event Example in Auger Observatory

Energy calibration based on experimental data
(including invisible energy correction)



Pierre Auger Collaboration

~450 Collaborators; 92 Institutions, 18 Countries:

Argentina
Australia
Brazil
Czech Republic
France
Germany (Aachen, Bonn, Hamburg, KIT, Siegen, Wuppertal)
Italy
Mexico
Netherlands
Poland
Portugal
Romania
Slovenia
Spain
UK
USA

Associated: **Colombia**, **Belgium**

EoI: **Peru**, **Bolivia**

New members are welcome!

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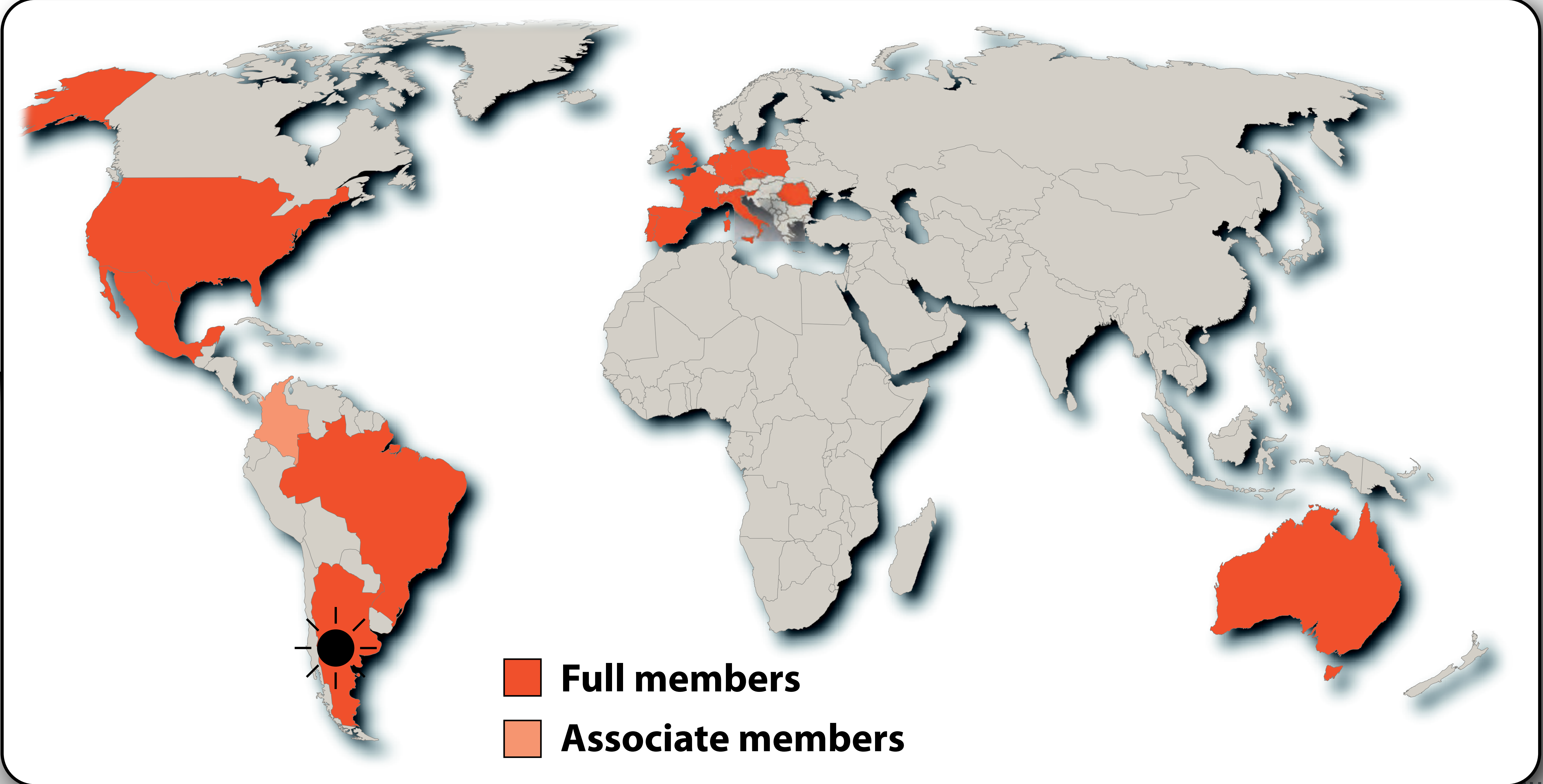
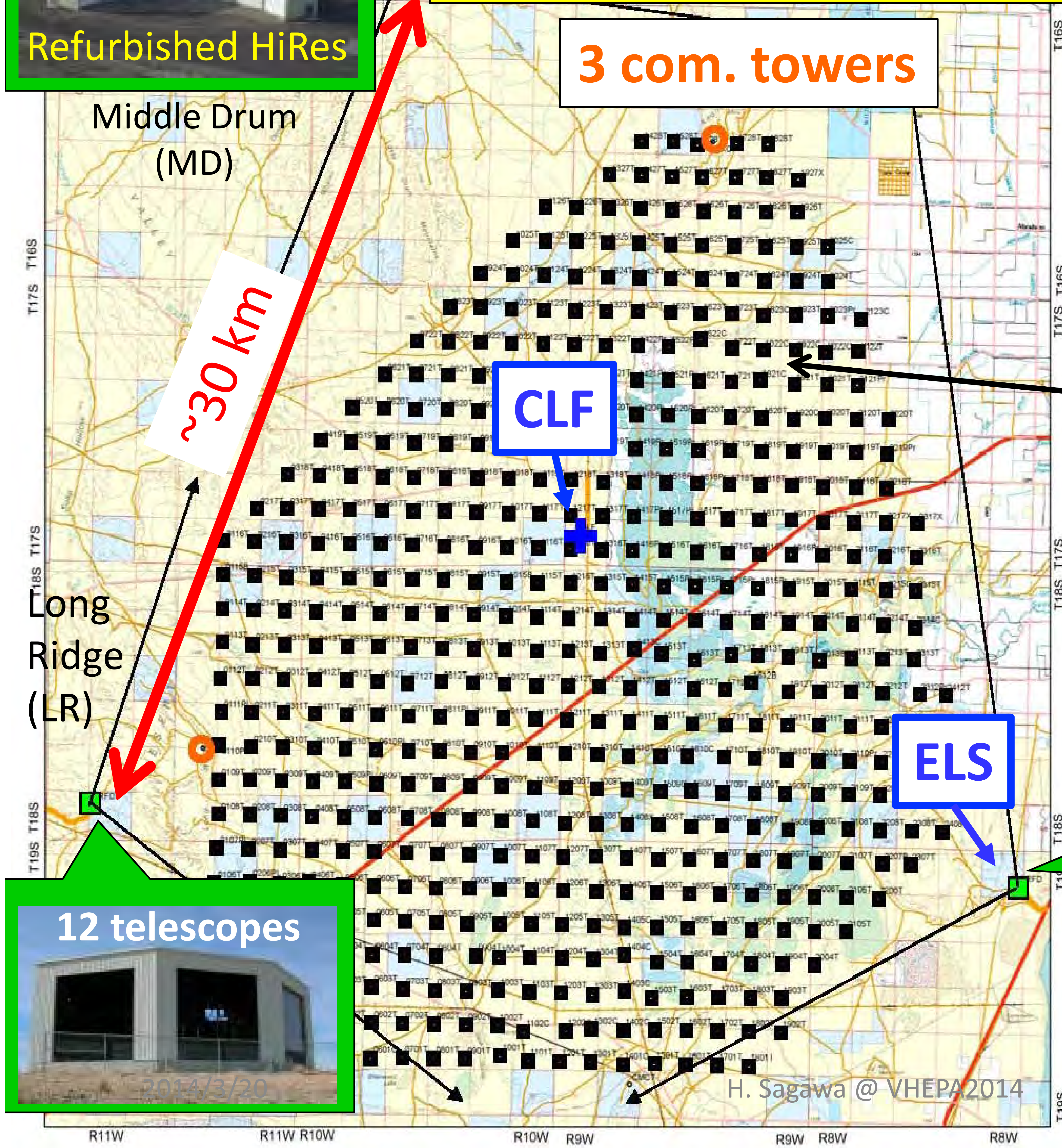


TA detector in Utah

39.3°N, 112.9°W
~1400 m a.s.l.



Surface Detector (SD)

507 plastic scintillator SDs

1.2 km spacing

~700 km²

to be upgraded to 2800 km² with coarser spacing

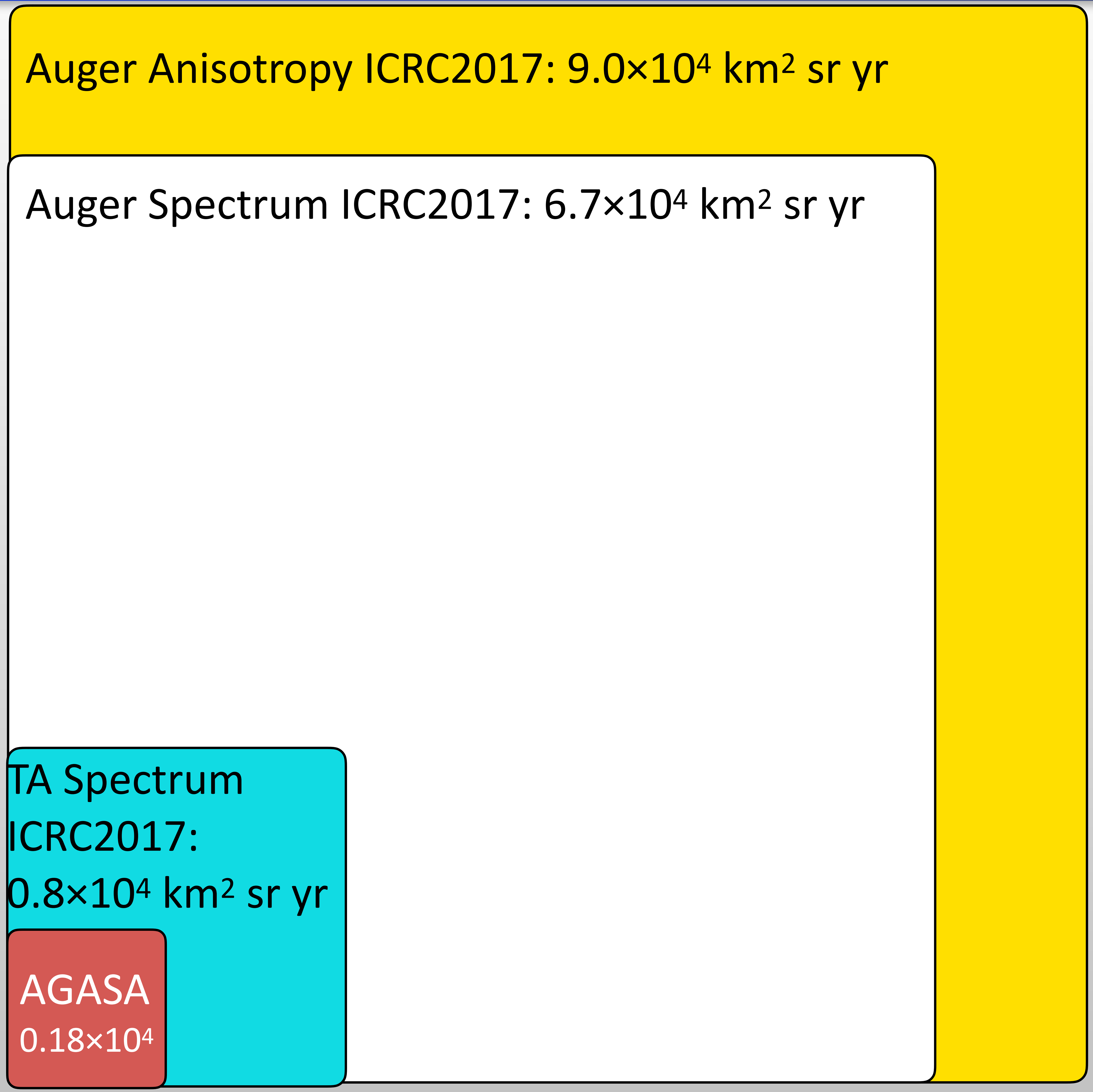
Fluorescence Detector(FD)

3 stations

38 telescopes

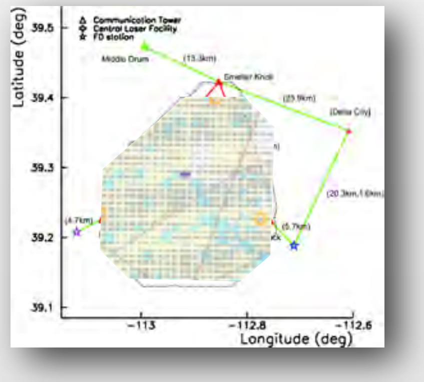
FD and SD: fully operational since 2008/May 4

UHE Exposure in Comparison

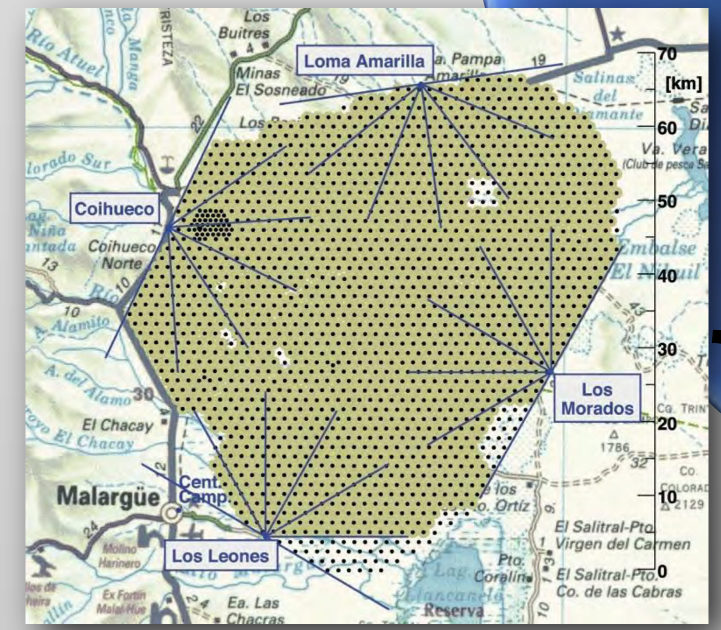


Combined Auger and TA cover full sky

Telescope Array (TA)
Delta, UT, USA
507 detector stations, 680 km²
36 fluorescence telescopes

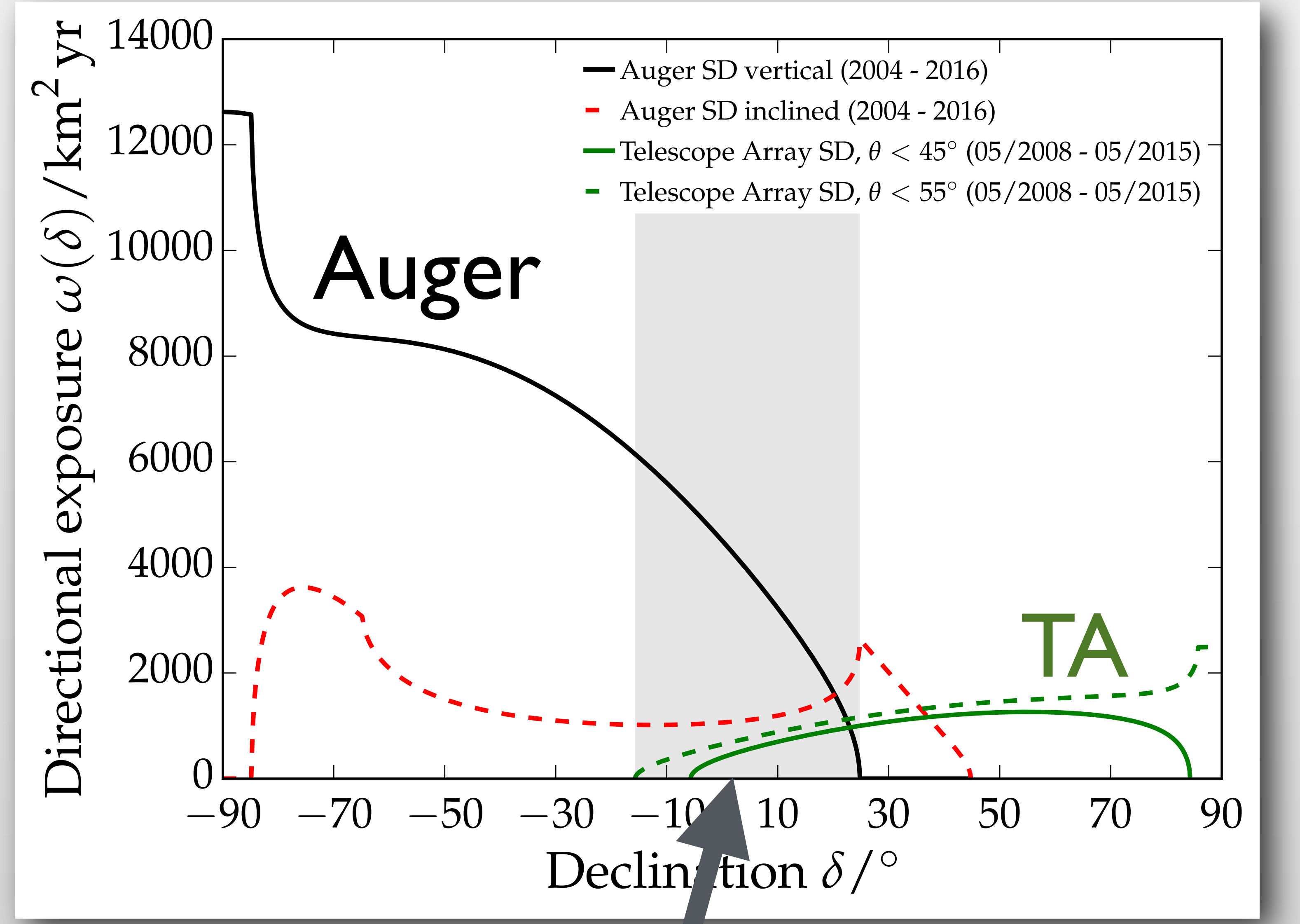


same scale



Pierre Auger Observatory
Province Mendoza, Argentina
1660 detector stations, 3000 km²
27 fluorescence telescopes

Auger: started 01/2004
TA: started 05/2008

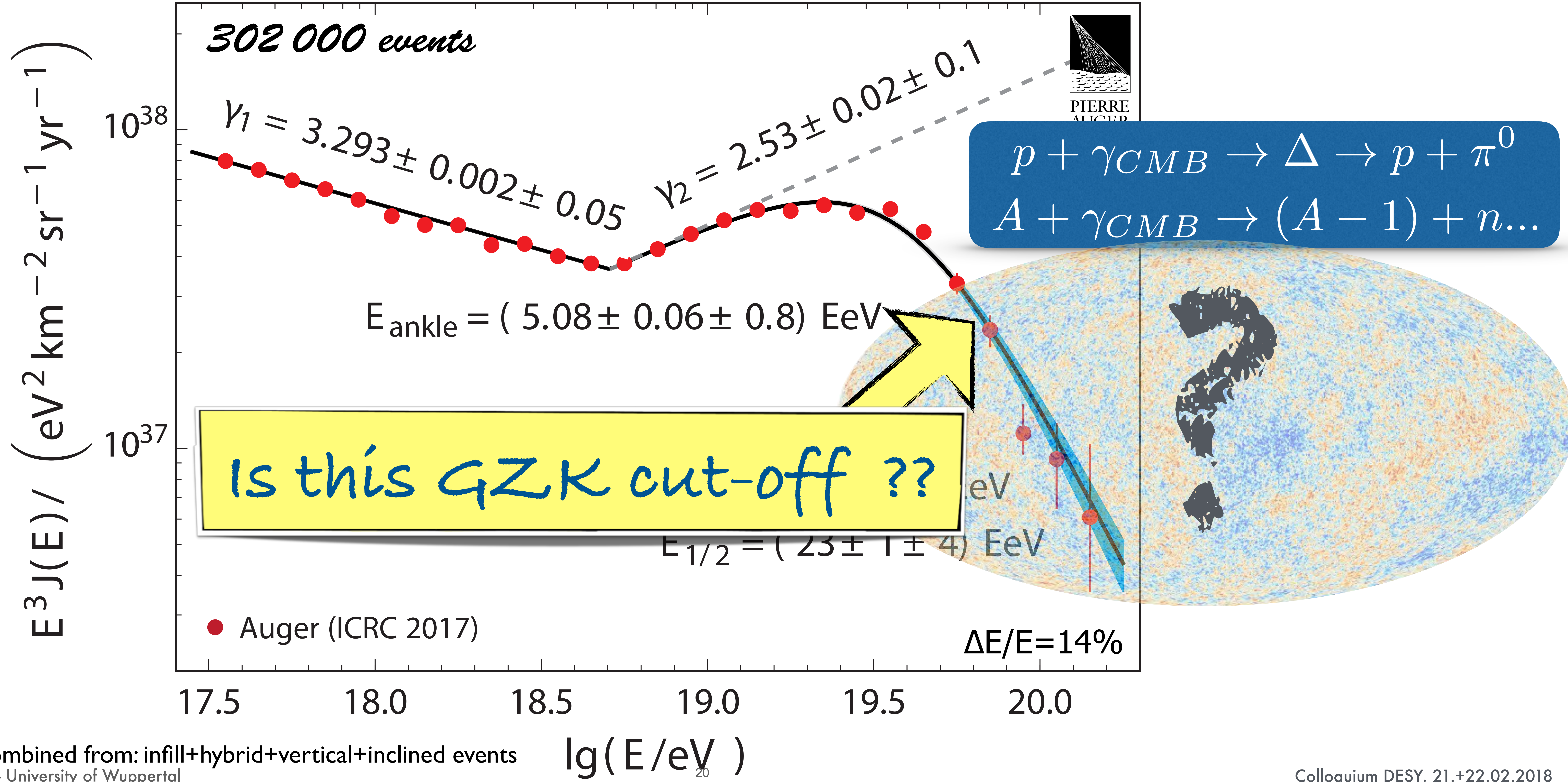


Auger and TA can see the same sky

UHECR
Energy Spectrum

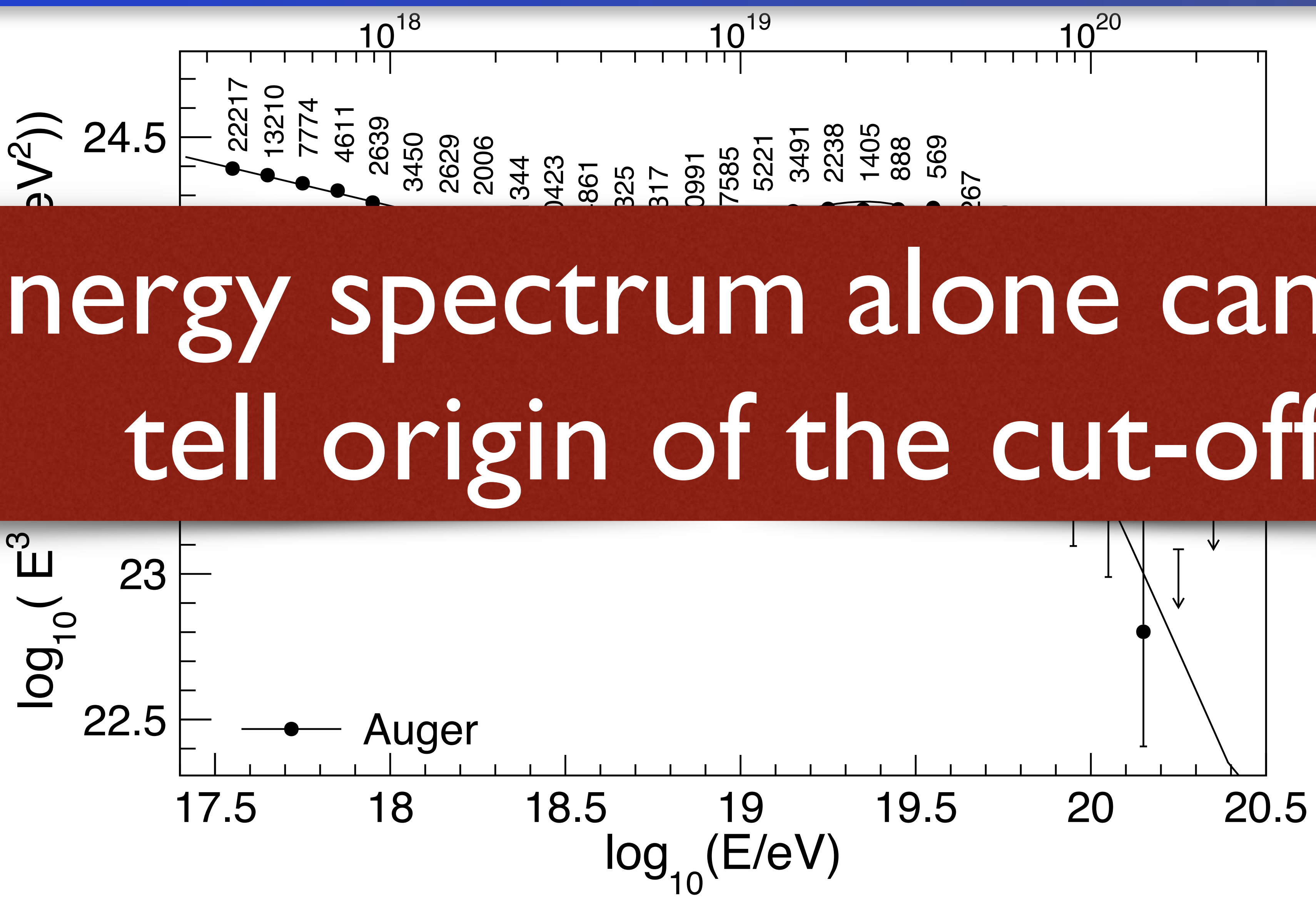
End of the CR-Spectrum (0°-80°)

arXiv:1708.06592
Update from: PRL 101, 061101 (2008), Physics Letters B 685 (2010) 239



combined from: infill+hybrid+vertical+inclined events

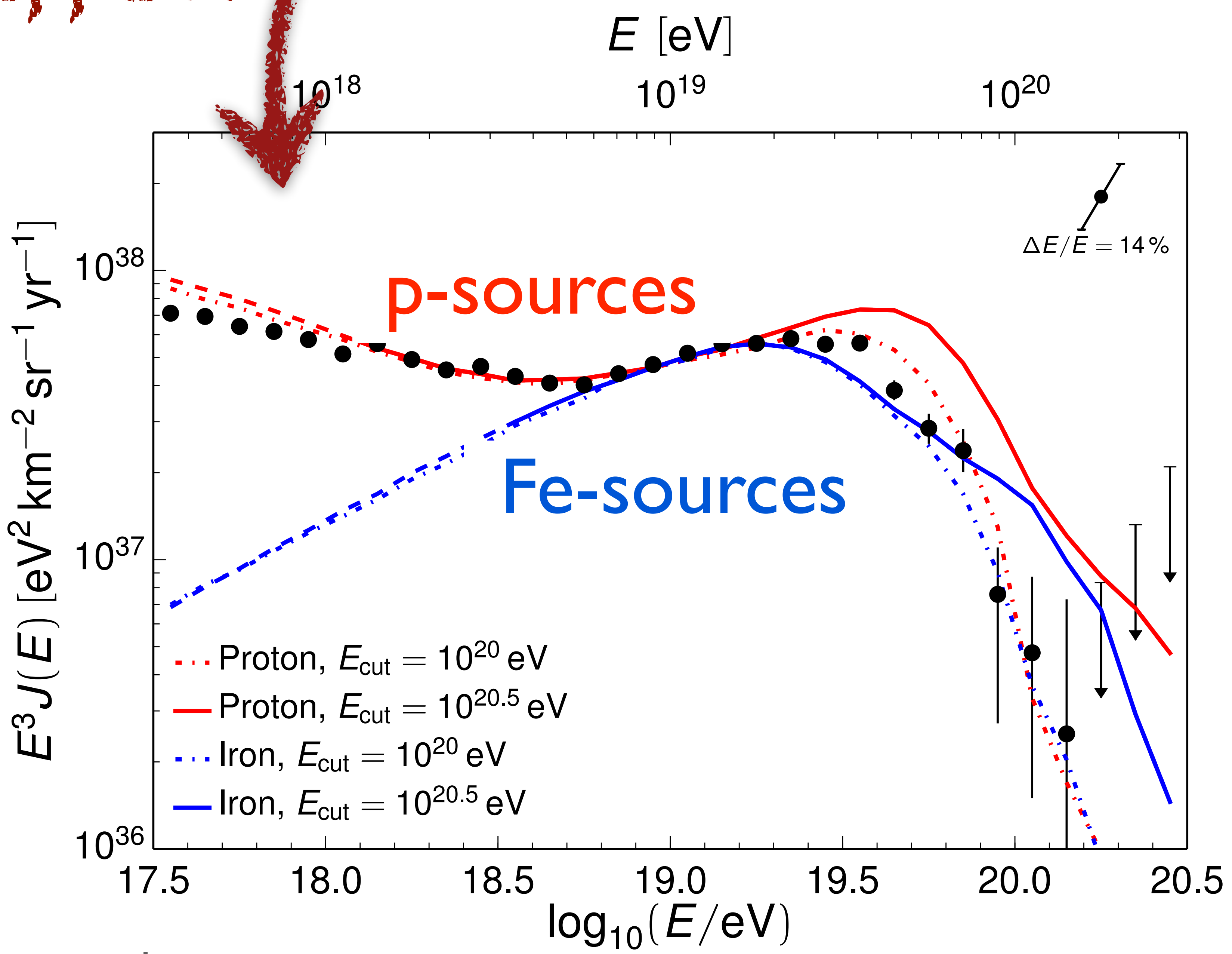
GZK-Effect or Exhausted Sources?



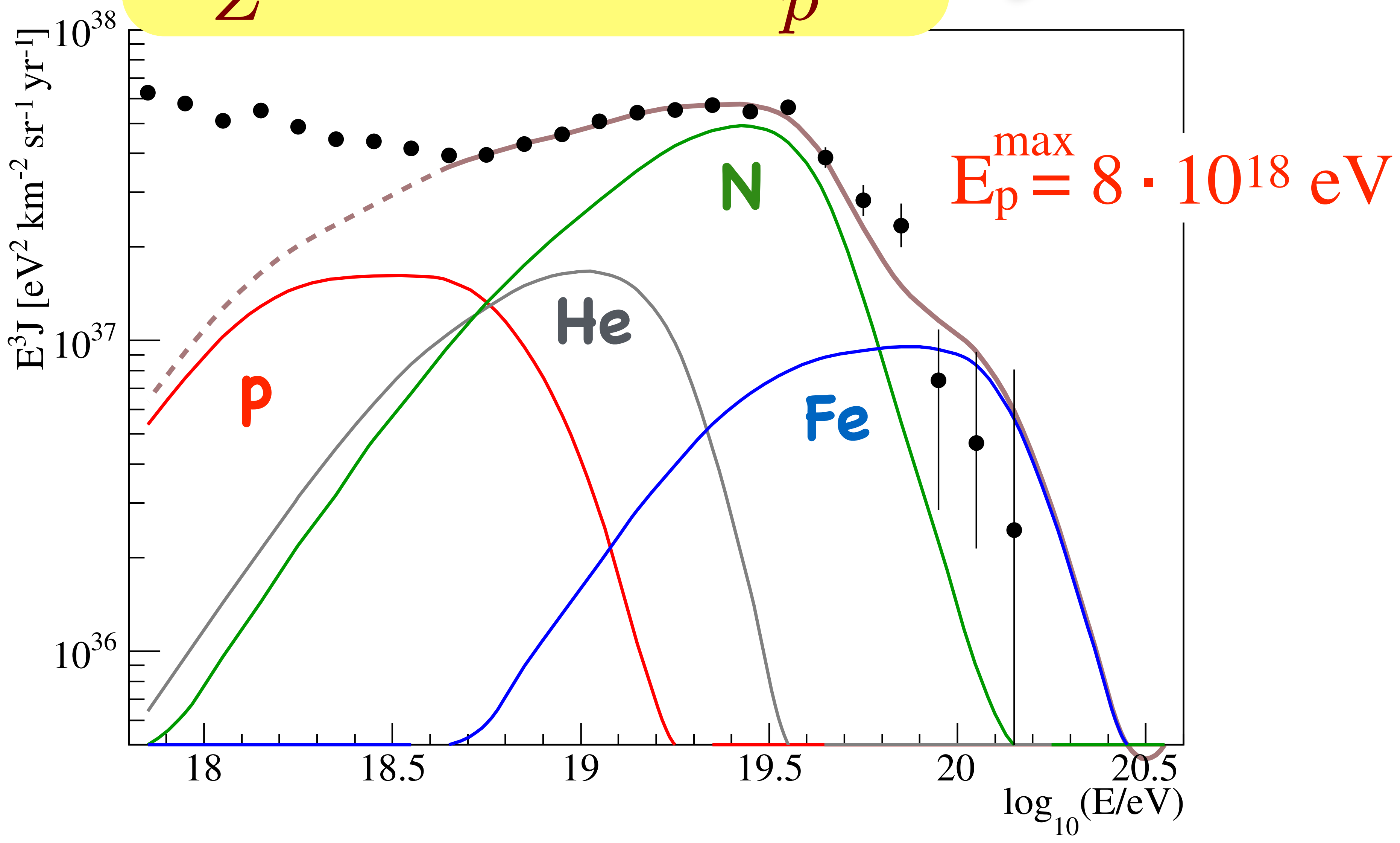
Energy spectrum alone cannot tell origin of the cut-off

GZK-effect

exhausted sources

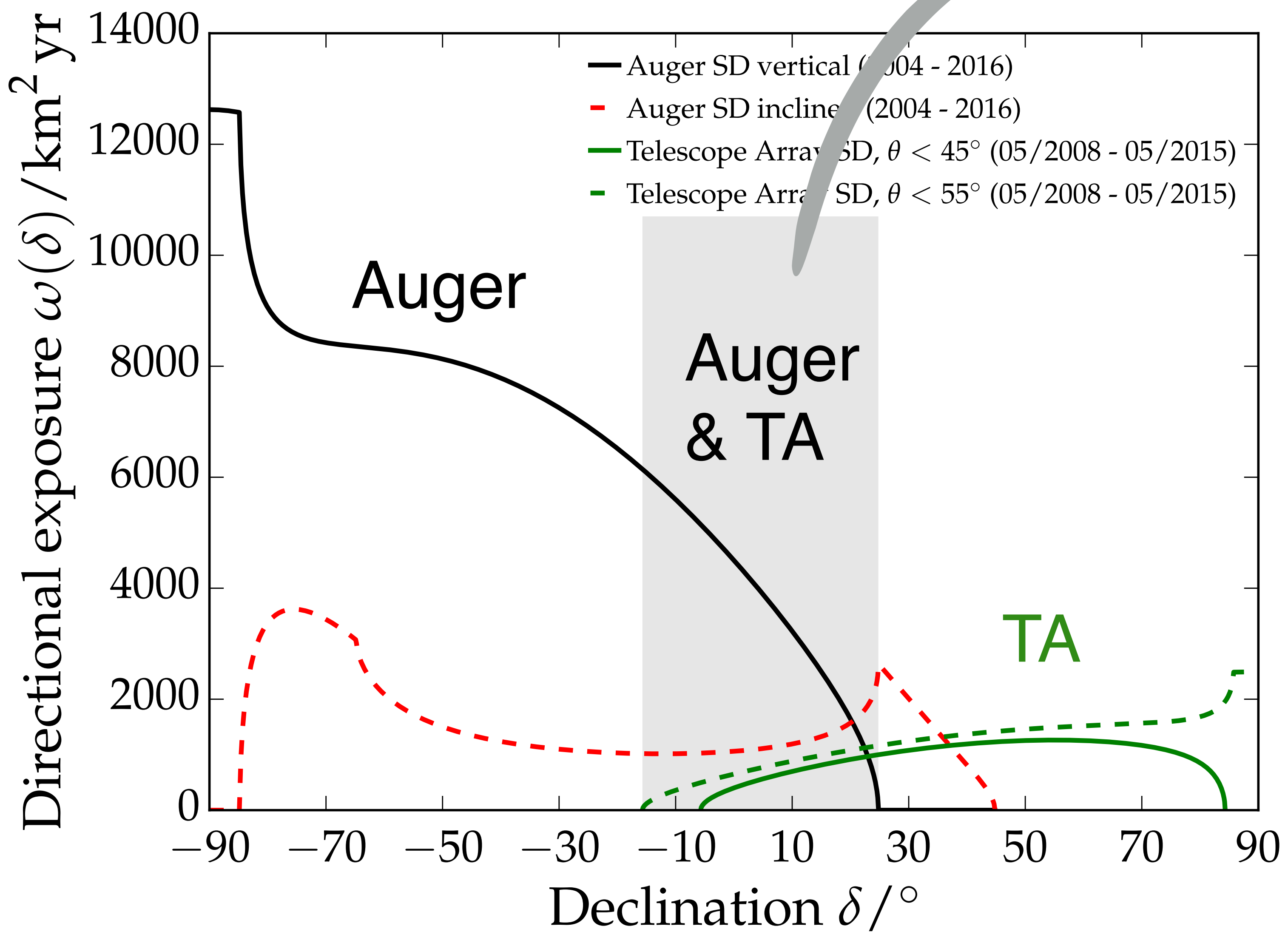


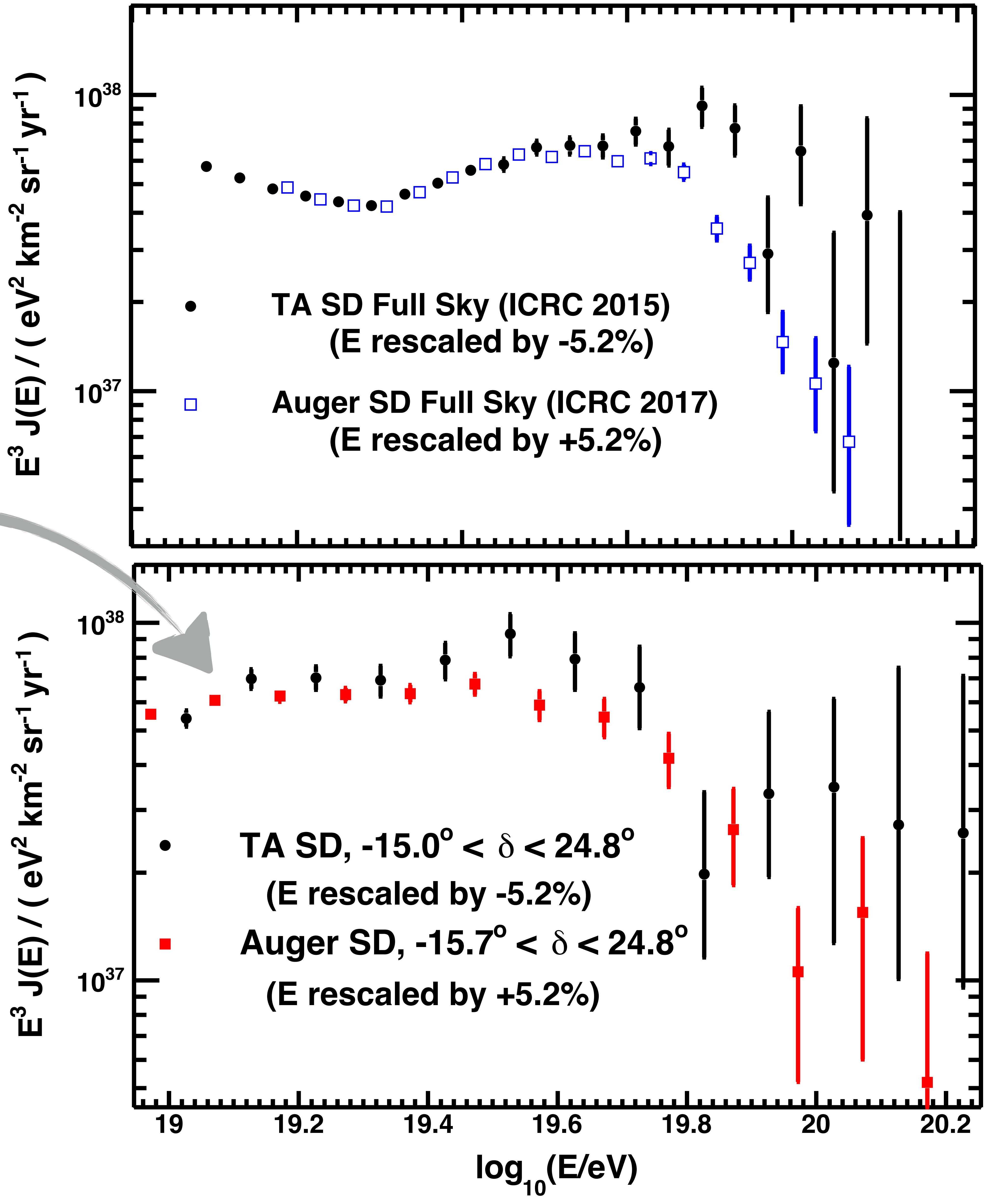
$E_Z^{max} \propto Z \times E_p^{max}$



Energy Spectra : Comparison Auger & TA

Auger-TA Working Group
@ ICRC2017



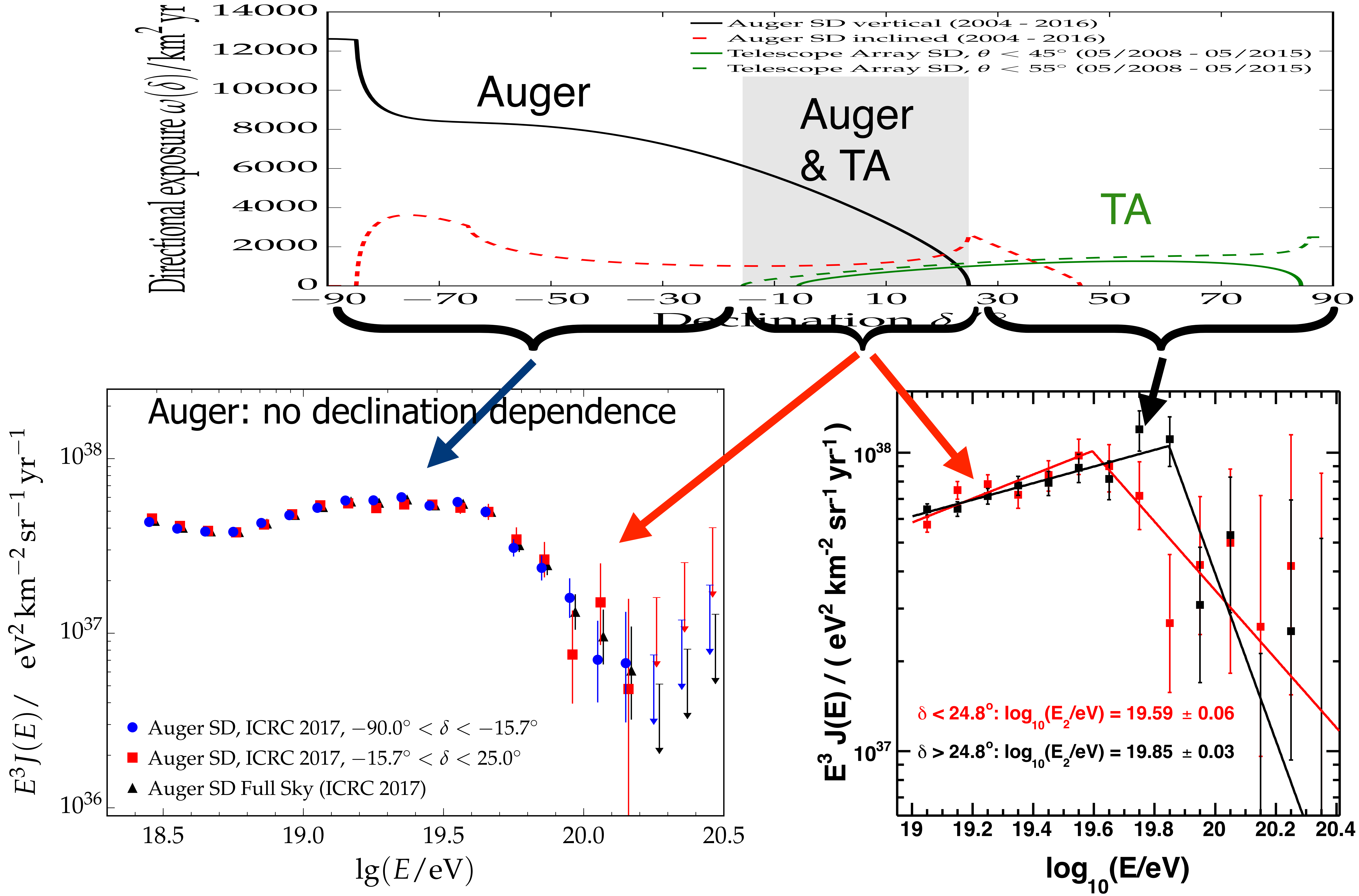


Each in their
full sky

Instrumental
or astrophysical
effect?

Joint
declination band

Energy Spectrum in Declination Bands



Auger-TA Working Group @ ICRC2017

TA-collaboration; arXiv:1801.07820 → 3.5σ

Globus et al, arXiv:1707.01177 → would yield much larger anisotropies than observed

Need

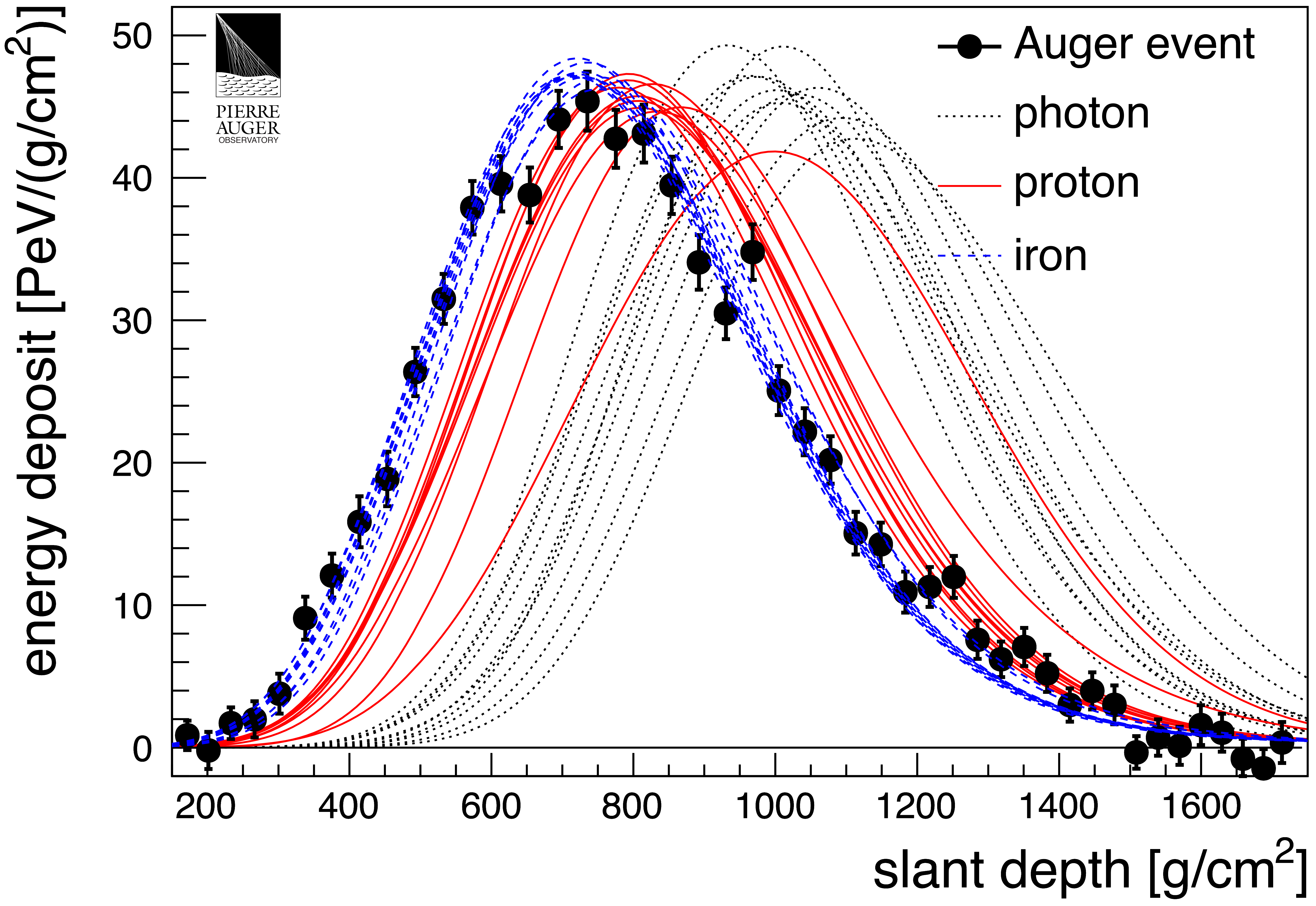
Mass Composition

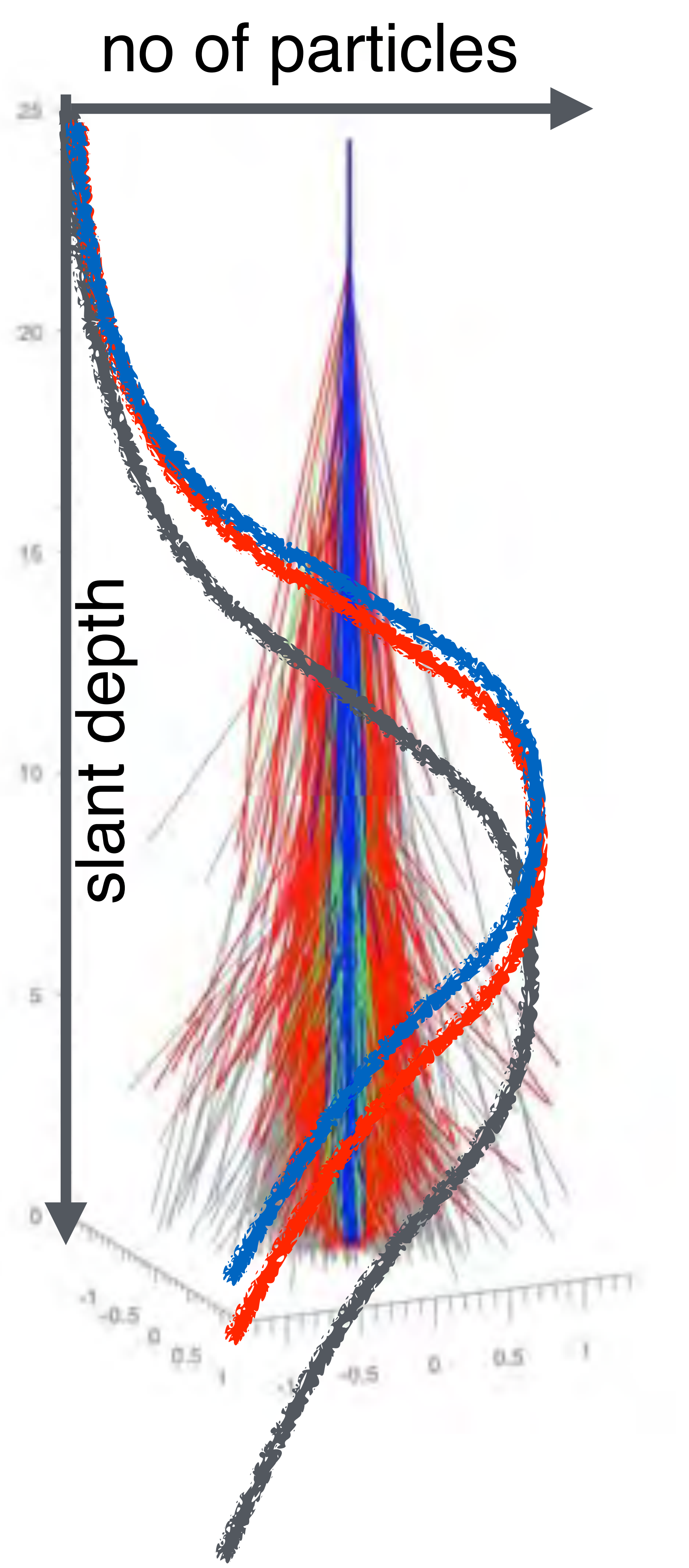
to disentangle GZK-suppression
from maximum energy scenario

Longitudinal Shower Development → Primary Mass

KHK, Unger, APP 35 (2012)
EPOS 1.99 Simulations

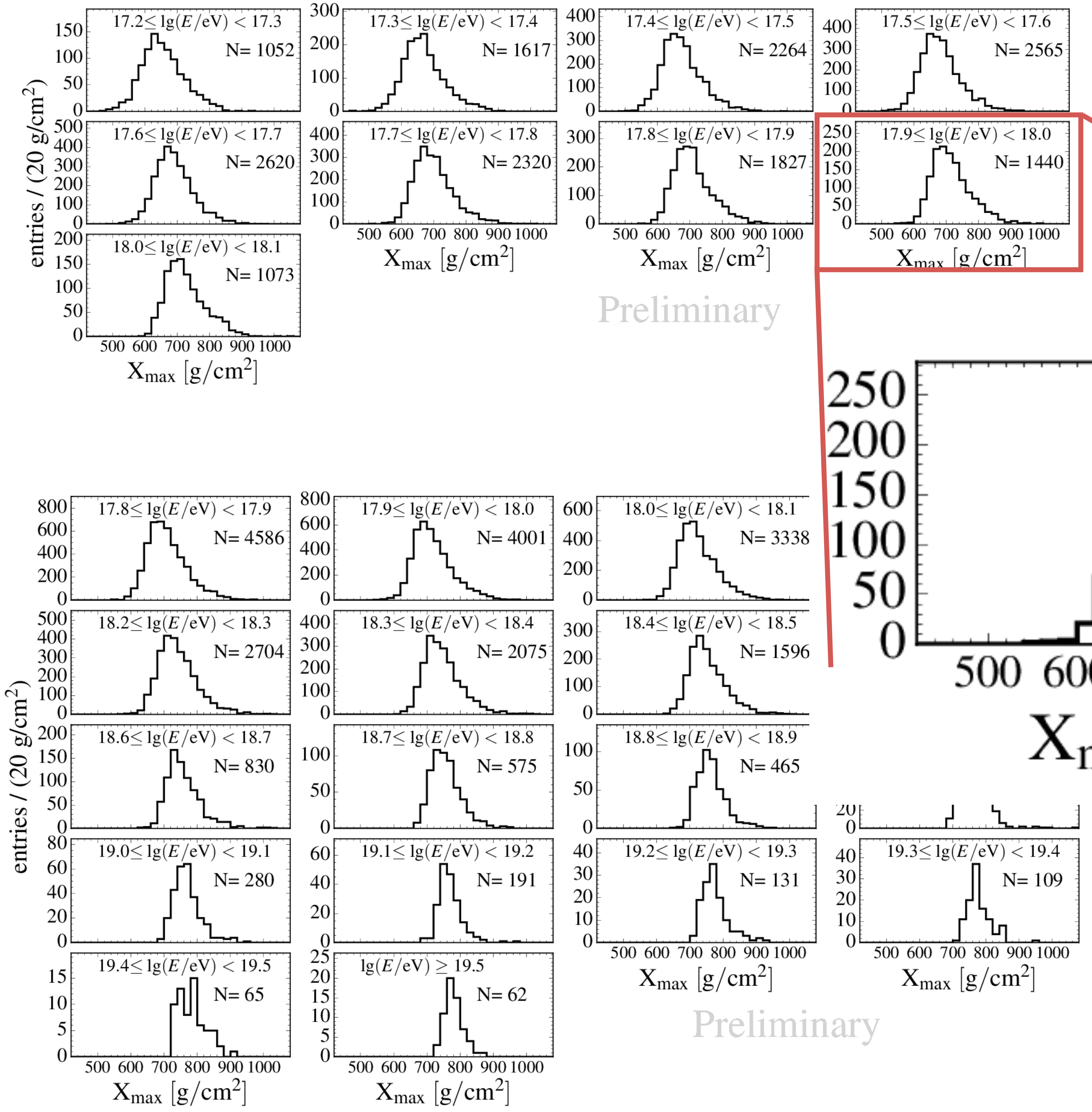
Example of a $3 \cdot 10^{19}$ eV EAS event in FD

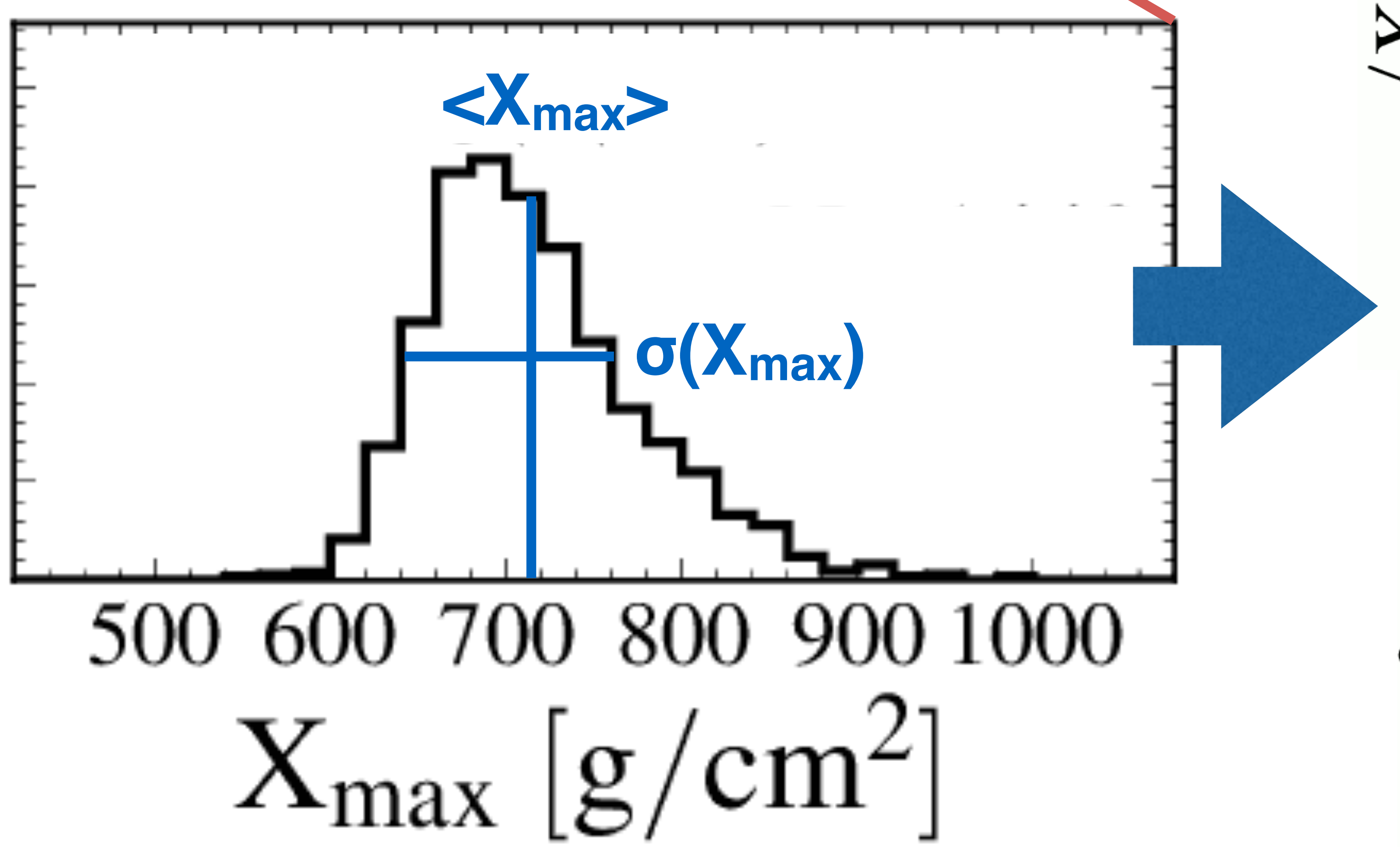


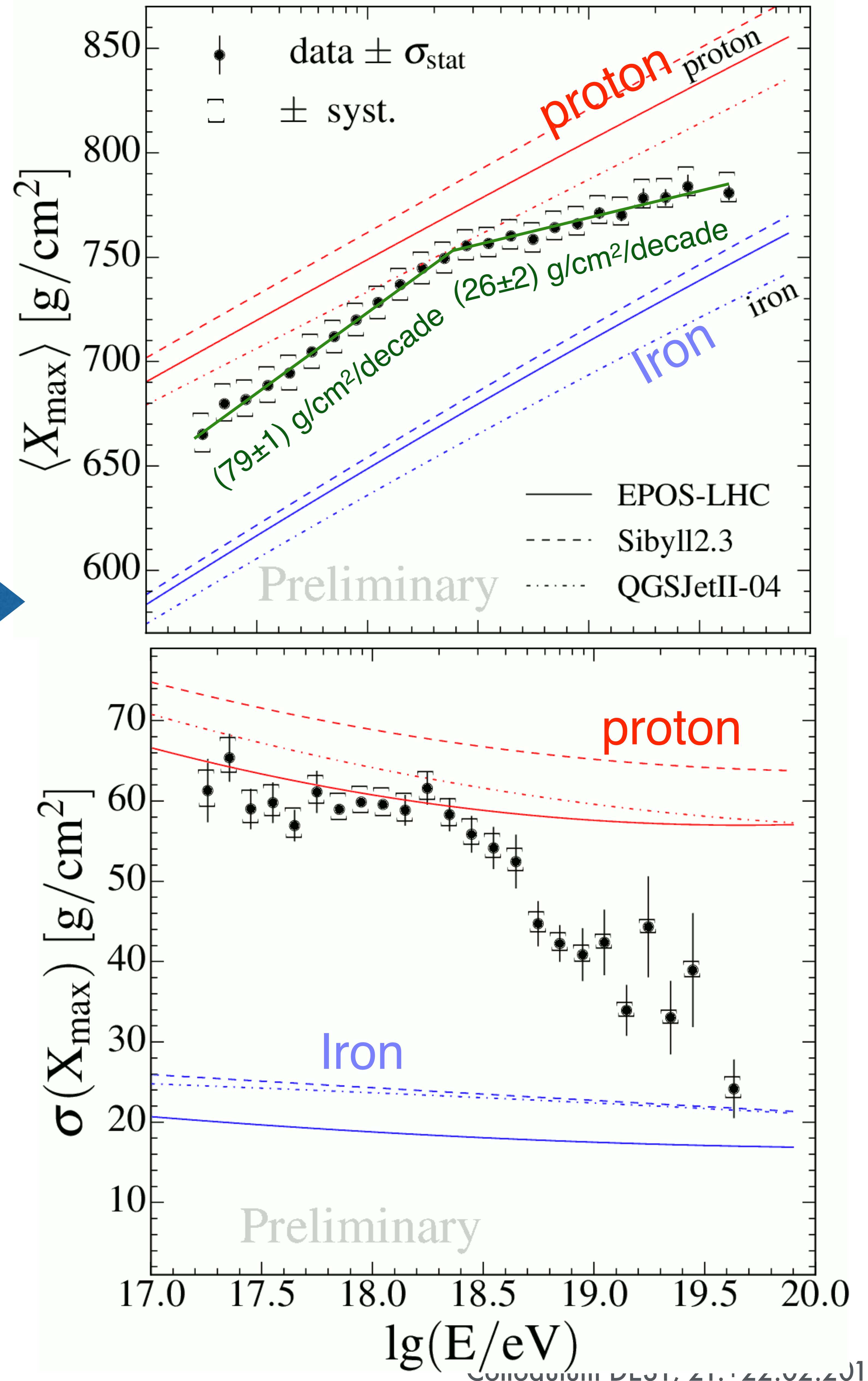


X_{max} Distributions

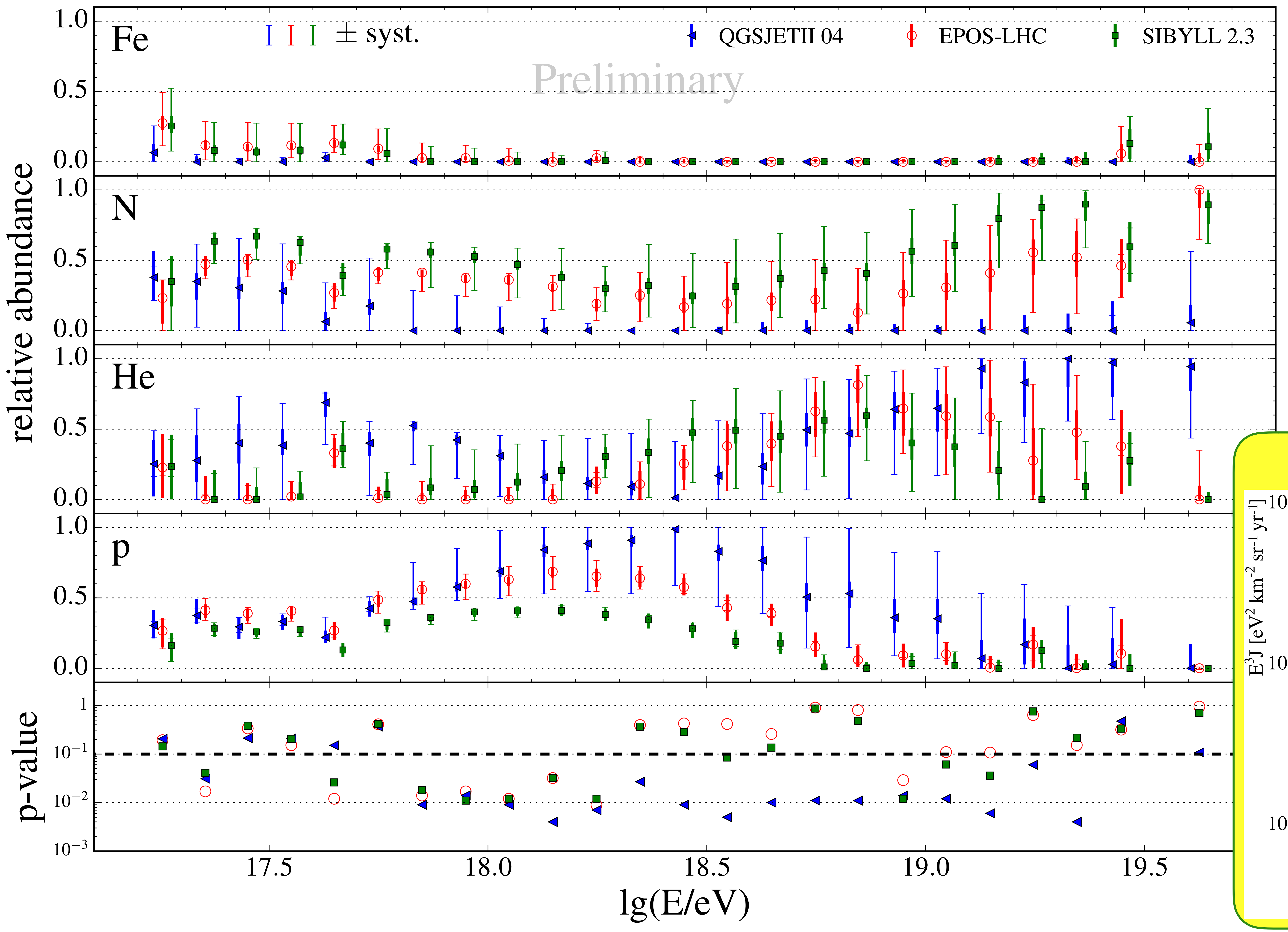
Bellido (Auger)
@ ICRC2017







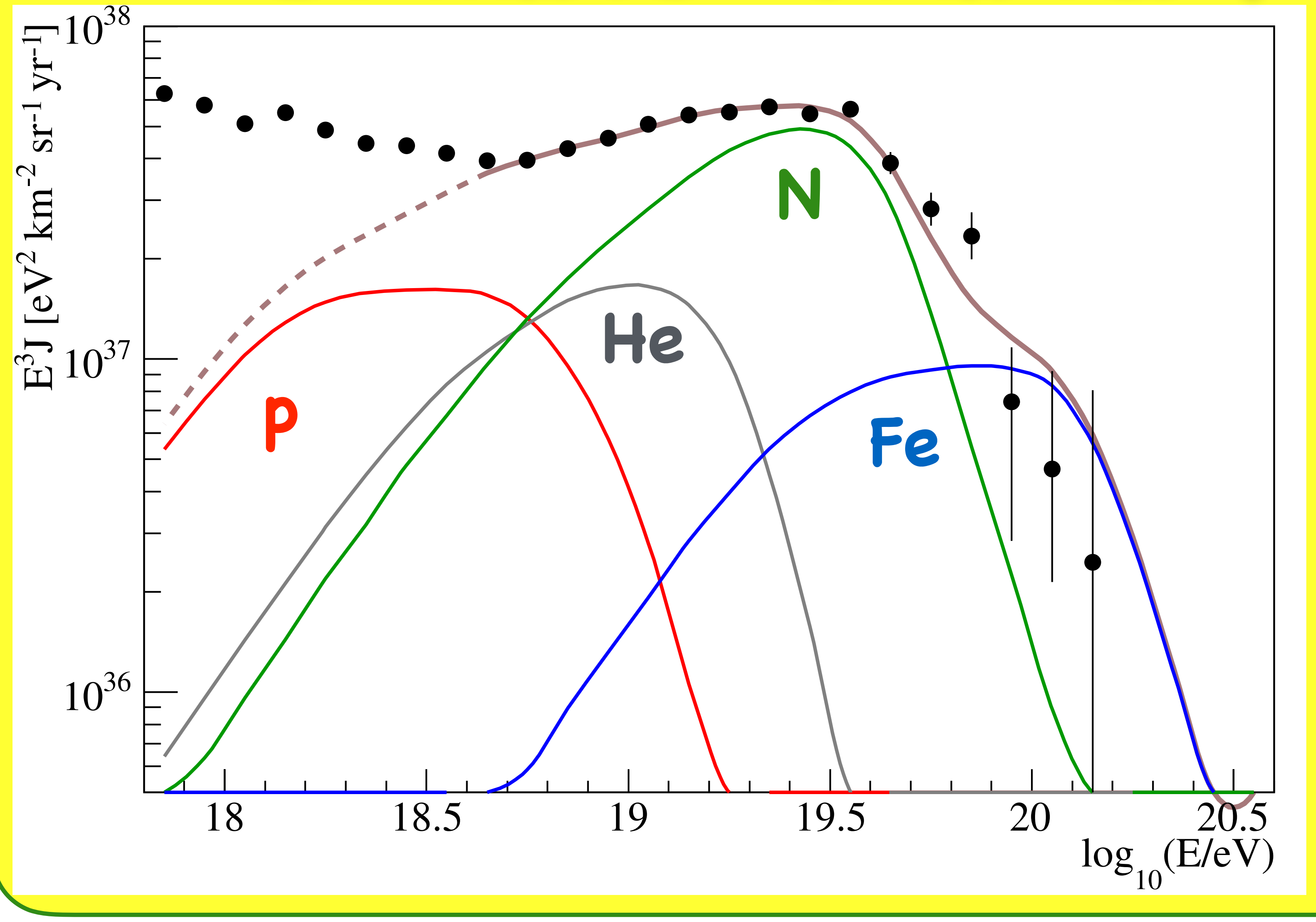
Mass Fractions



Bellido (Auger)
@ ICRC2017

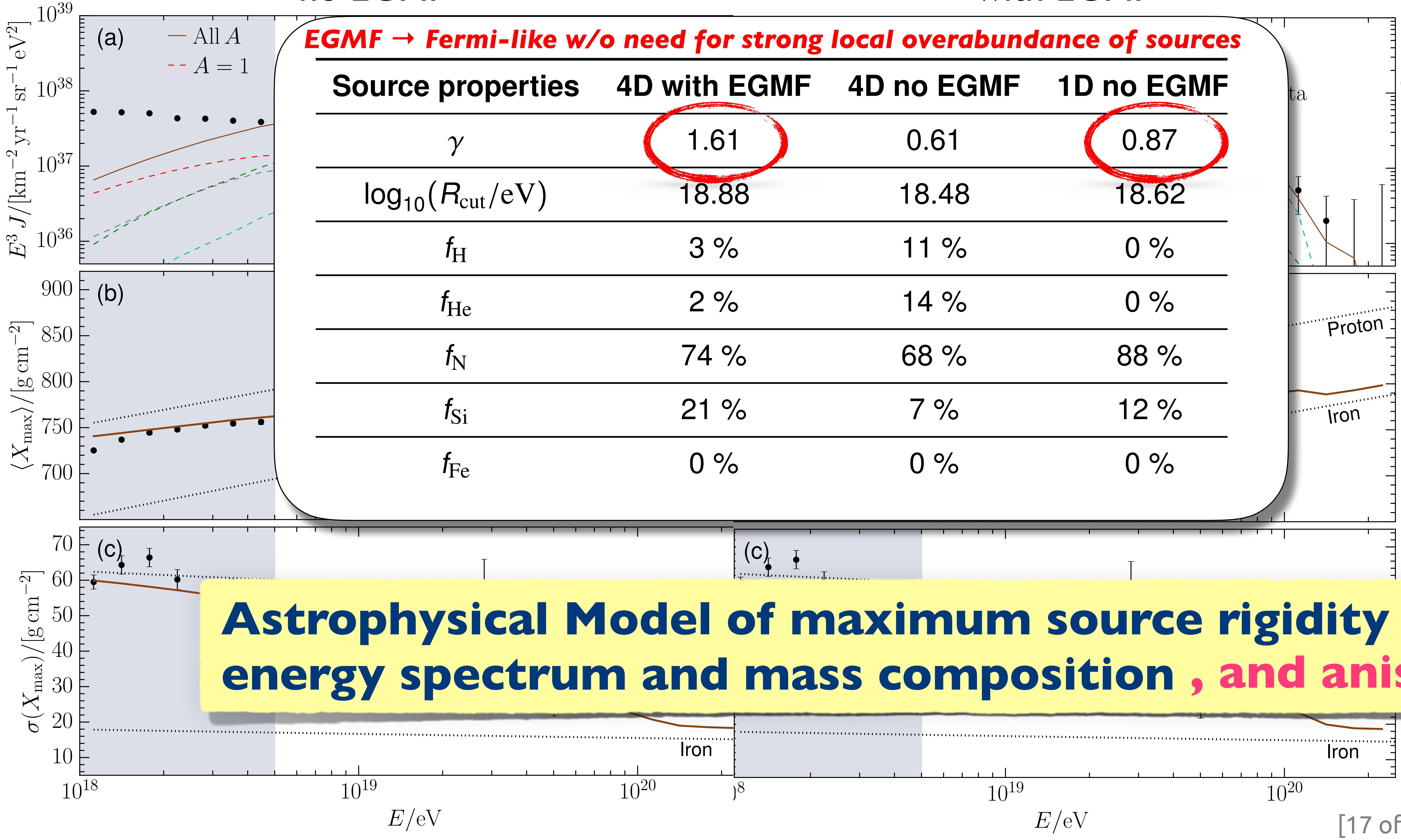
remember....

exhausted sources ?



Combined Fit of E-Spectr. and X_{max} Distr.

no EGMF   with EGMF



EGMF → Fermi-like w/o need for strong local overabundance of sources

Source properties	4D with EGMF	4D no EGMF	1D no EGMF
γ	1.61	0.61	0.87
$\log_{10}(R_{cut}/\mathrm{eV})$	18.88	18.48	18.62
f_H	3 %	11 %	0 %
f_{He}	2 %	14 %	0 %
f_N	74 %	68 %	88 %
f_{Si}	21 %	7 %	12 %
f_{Fe}	0 %	0 %	0 %

rigidity-dependent cutoff at source:
$E_{max} = R_{cut} \times Z$,

power law injection $E^{-\gamma}$,

mass fractions f_A.

propagation with CRPropa 3,

Gilmore12 EBL,

Dolag12 LSS

Astrophysical Model of maximum source rigidity describes energy spectrum and mass composition , and anisotropy!

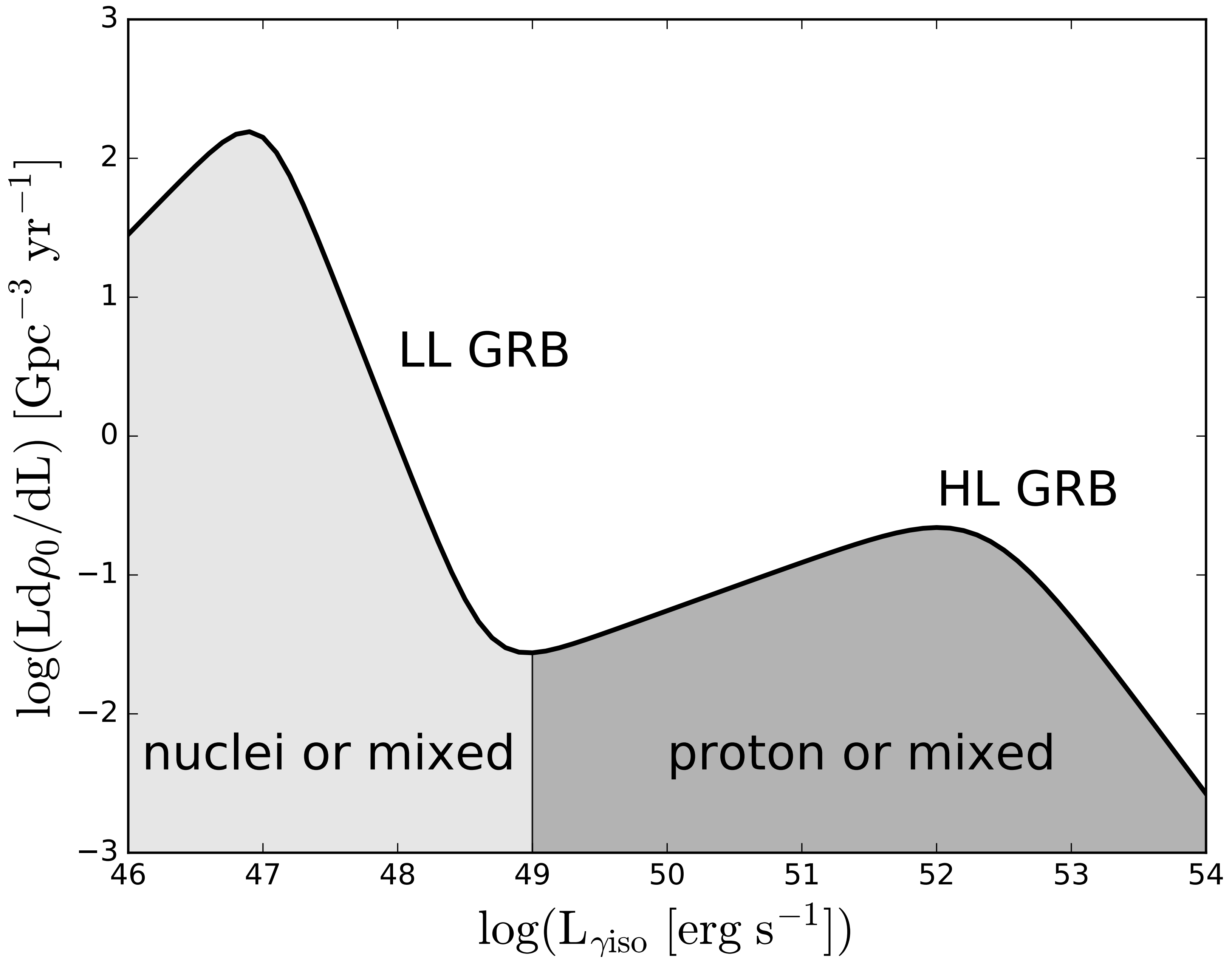
[17 of

Wittkowski (Auger) @ ICRC2017

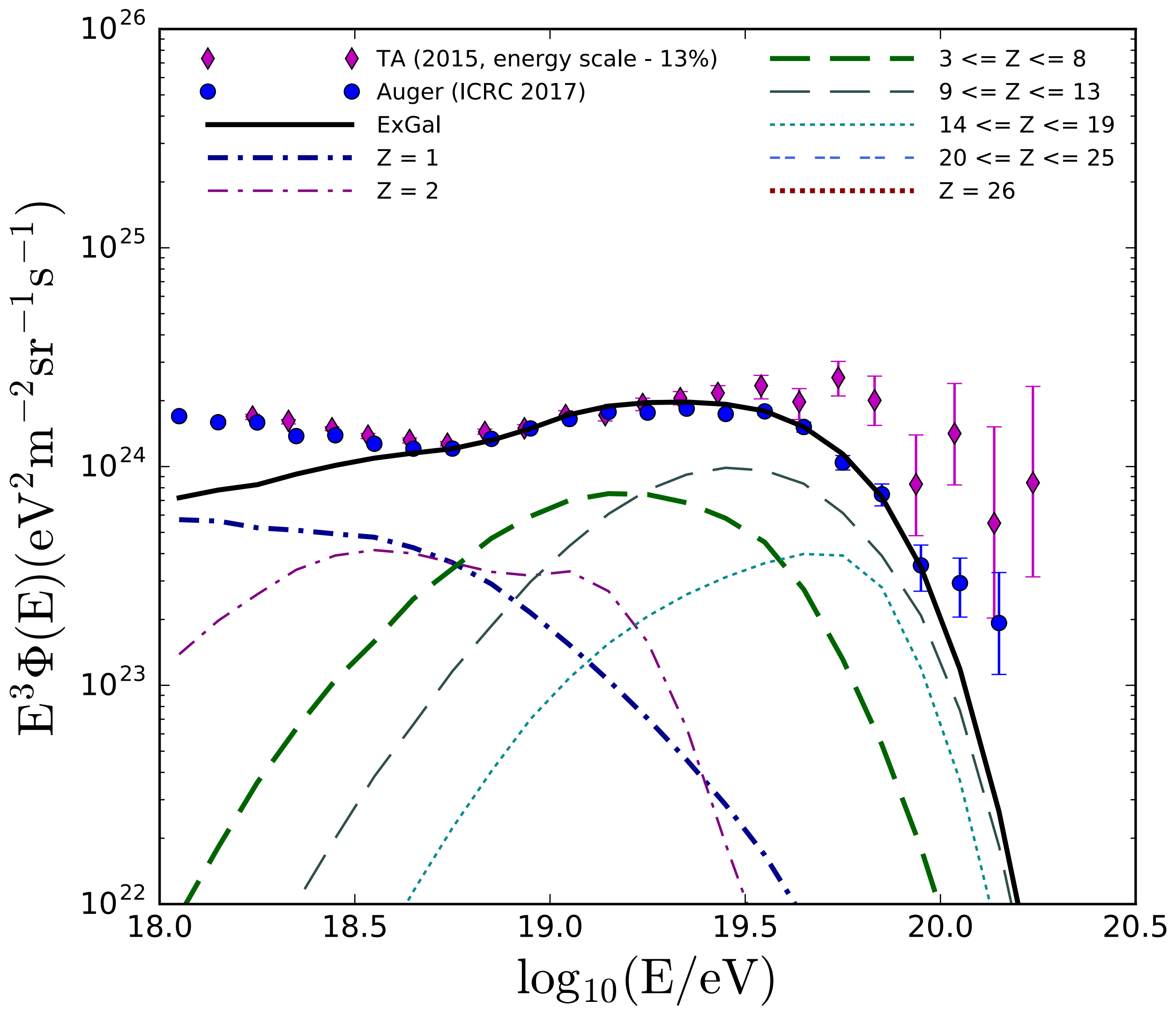
Models with specific source classes: LL/HL-GRBs

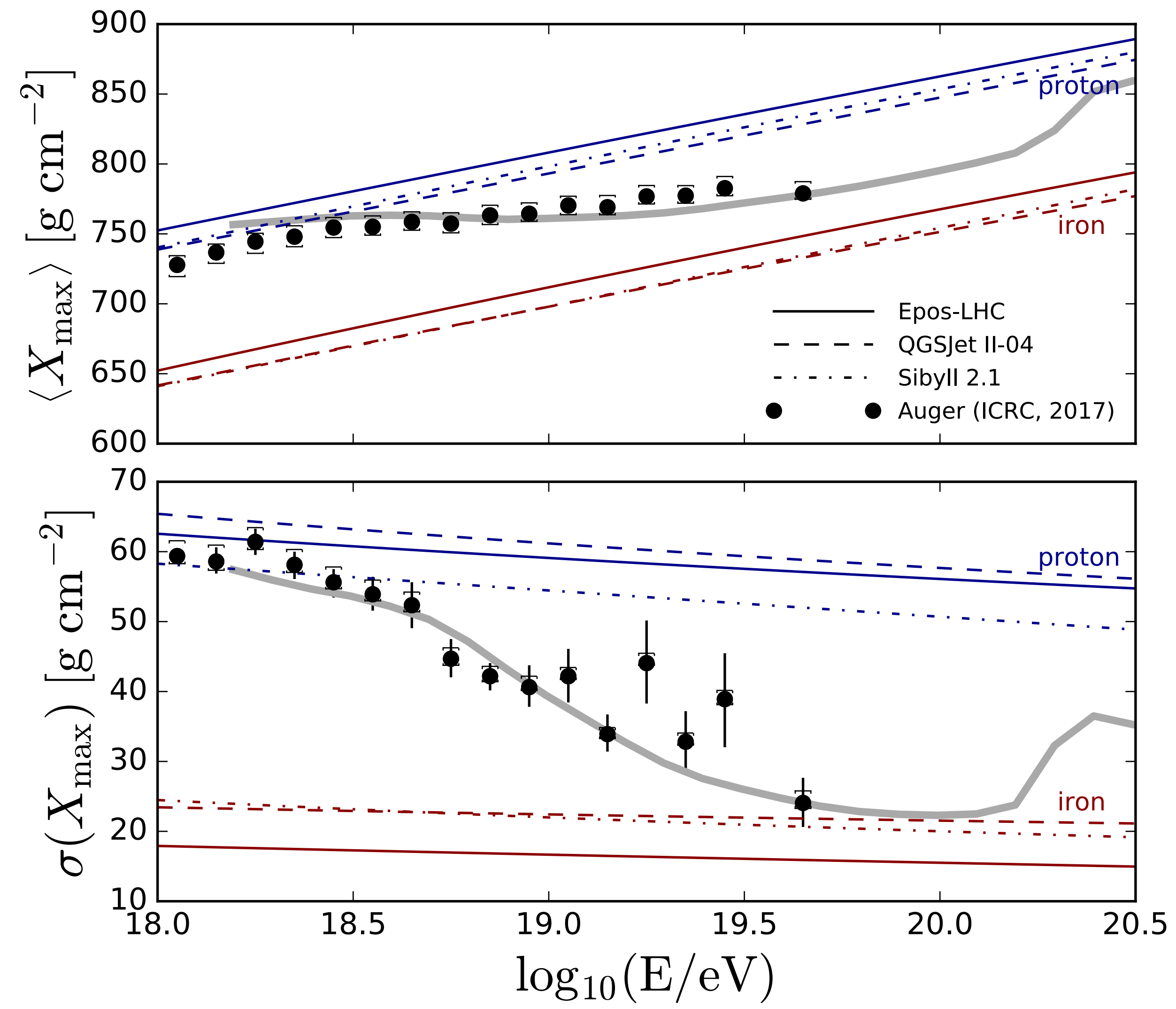
Zhang, Murase, et al.,
arXiv:1712.09984

Low and High Luminosity GRBs
as sources of UHECR

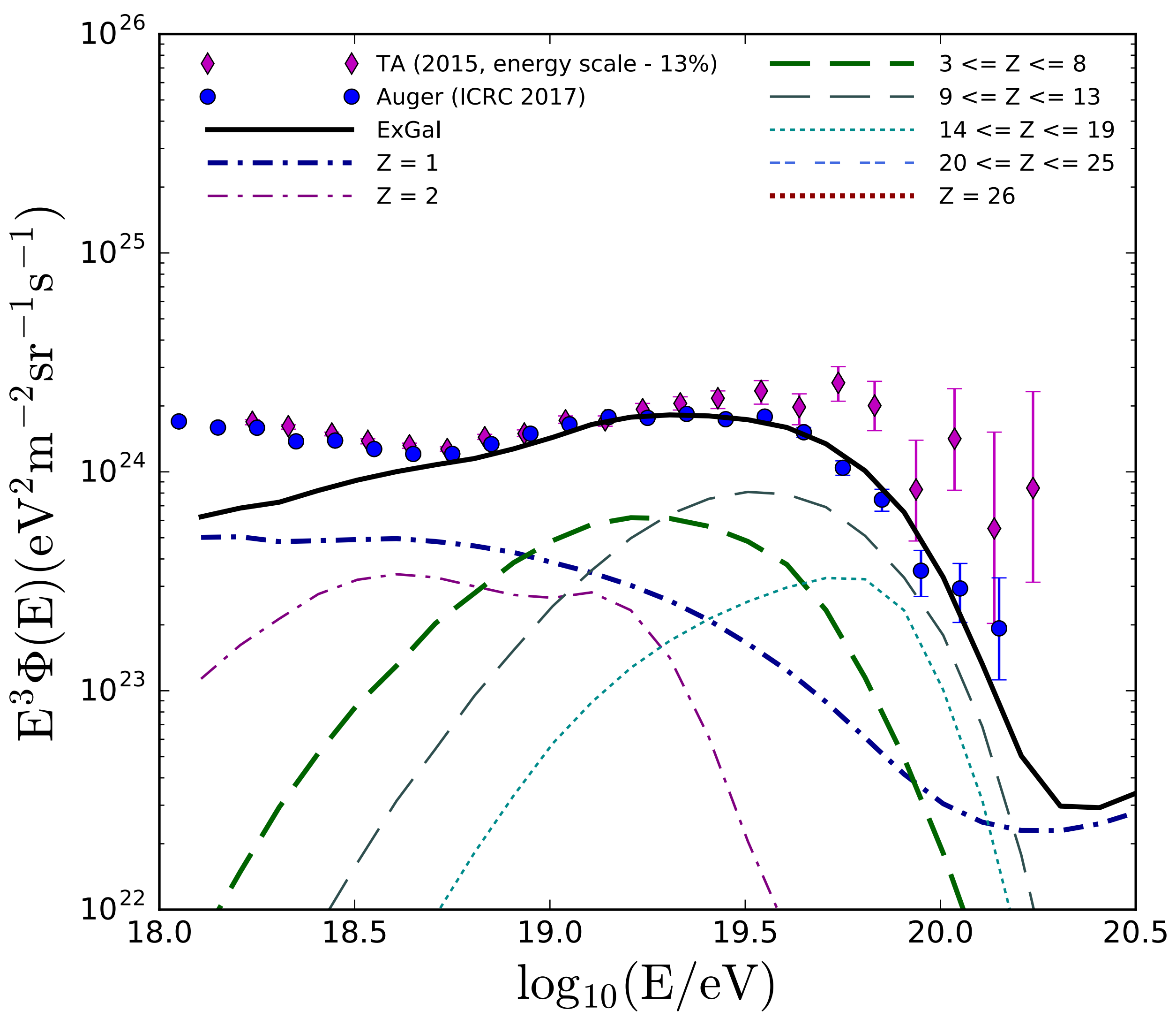


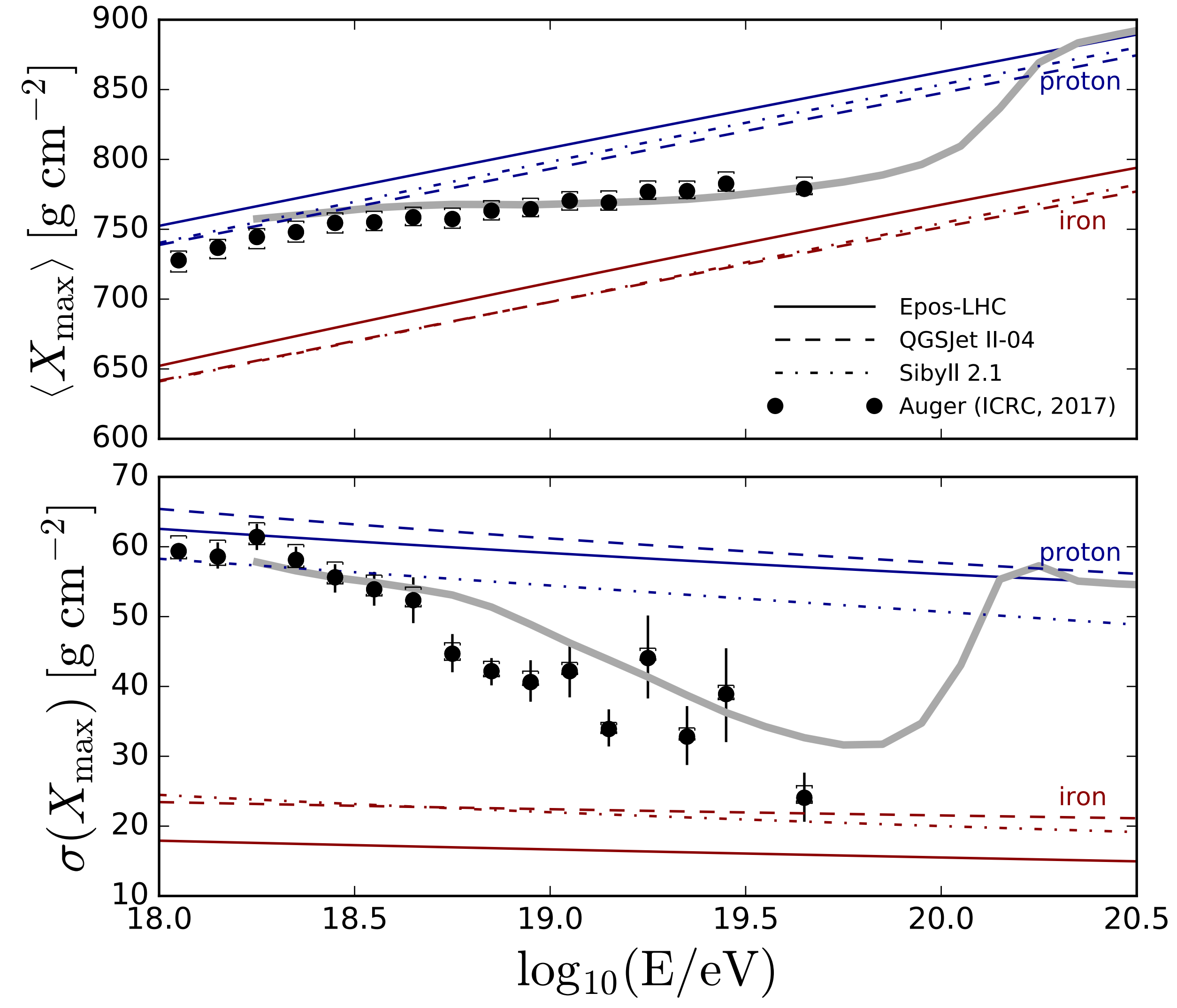
LL GRBs with Si-rich progenitor





LL GRBs + HL GRBs





Auger data and all these models suggest seeing exhausting sources

$\Rightarrow$ Strongly Suppressed Cosmogenic Neutrino Fluxes

Recall:

- If flux suppression above $5 \cdot 10^{19}$ eV is due to GZK-effect: expect cosmogenic neutrinos & photons

$$\boxed{\begin{aligned} p + \gamma_{CMB} \rightarrow \Delta &\rightarrow p + \pi^0 \rightarrow \gamma \\ &\rightarrow n + \pi^+ \rightarrow \nu \end{aligned}}$$

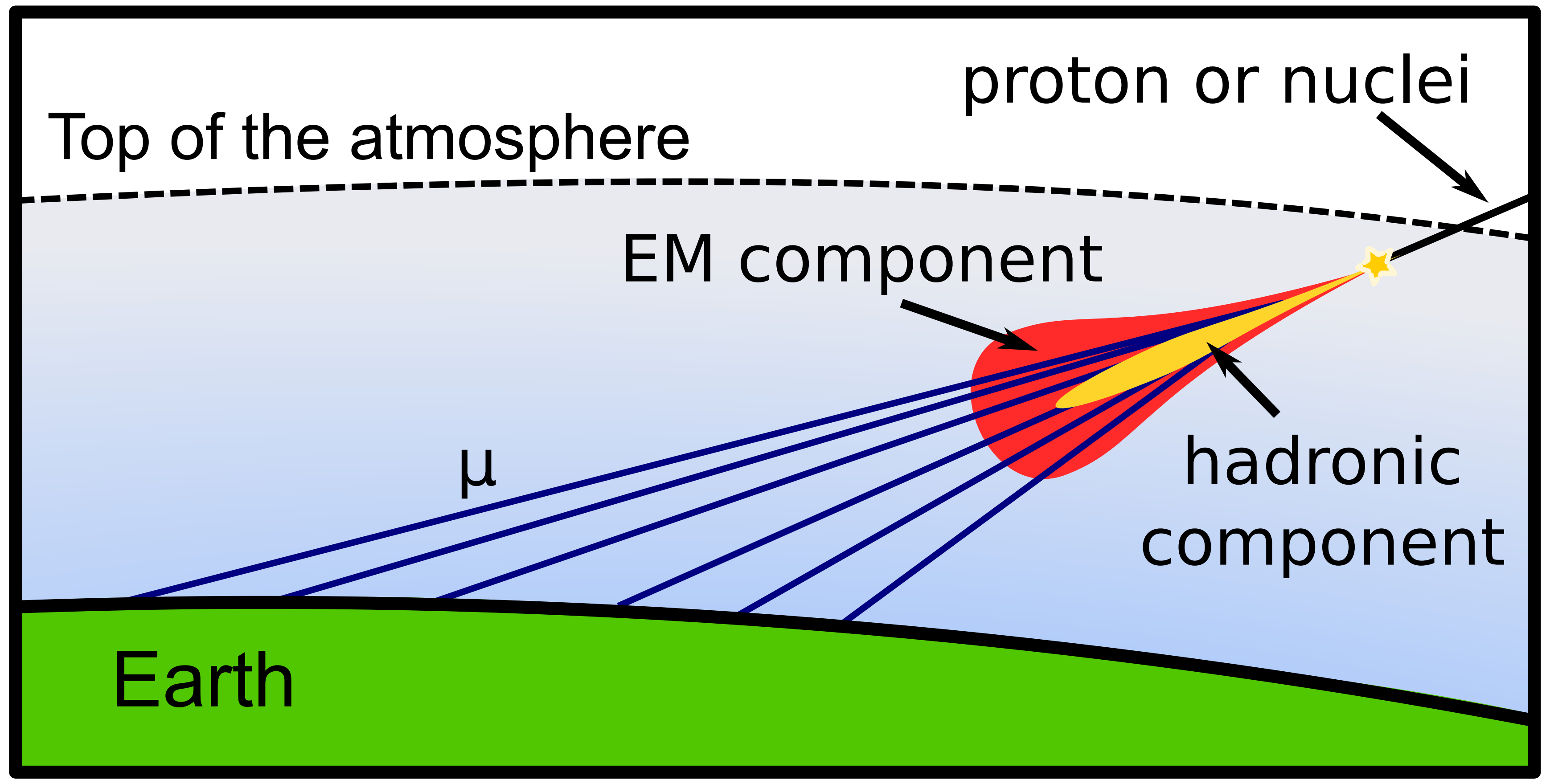
smoking gun...

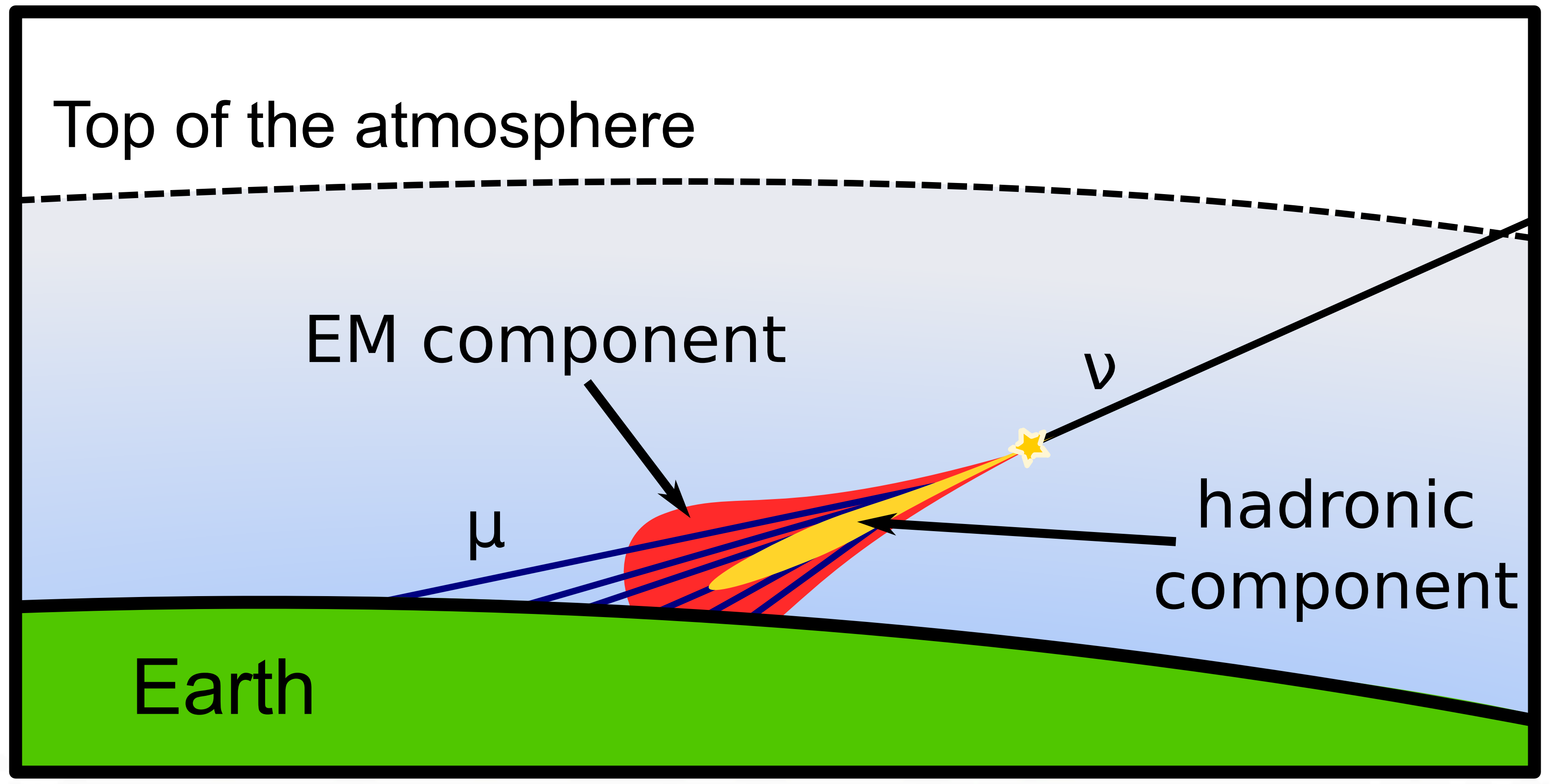
- If due to source exhaustion: neutrinos & photons strongly suppressed

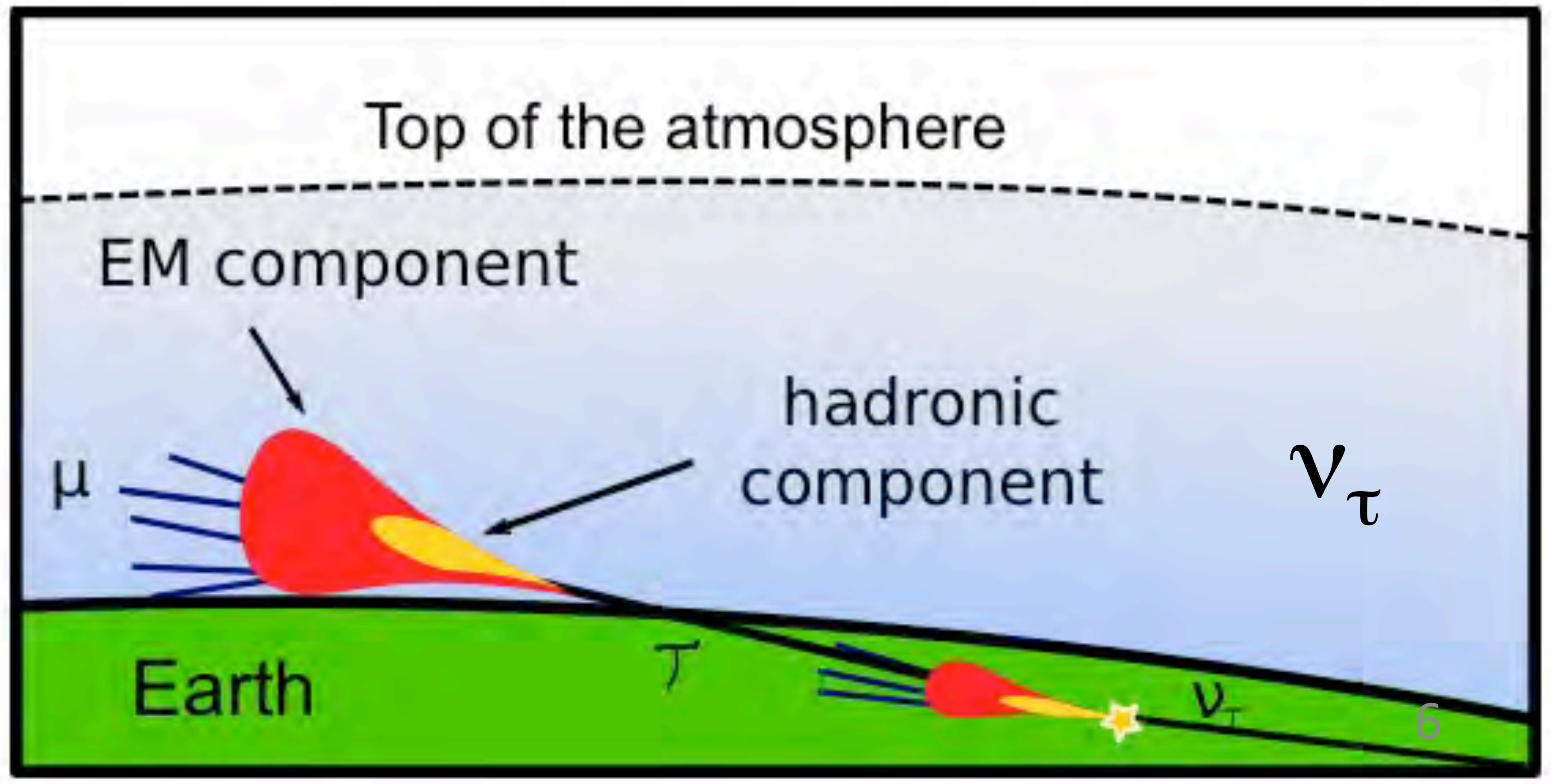
Search for EeV Neutrinos in inclined showers

- Protons & nuclei initiate showers high in the atmosphere.
  - Shower front at ground:
    - mainly composed of muons
    - electromagnetic component absorbed in atmosphere.
- Neutrinos can initiate “deep” showers close to ground.
  - Shower front at ground: electromagnetic + muonic components

Searching for neutrinos ⇒ searching for inclined showers with electromagnetic component

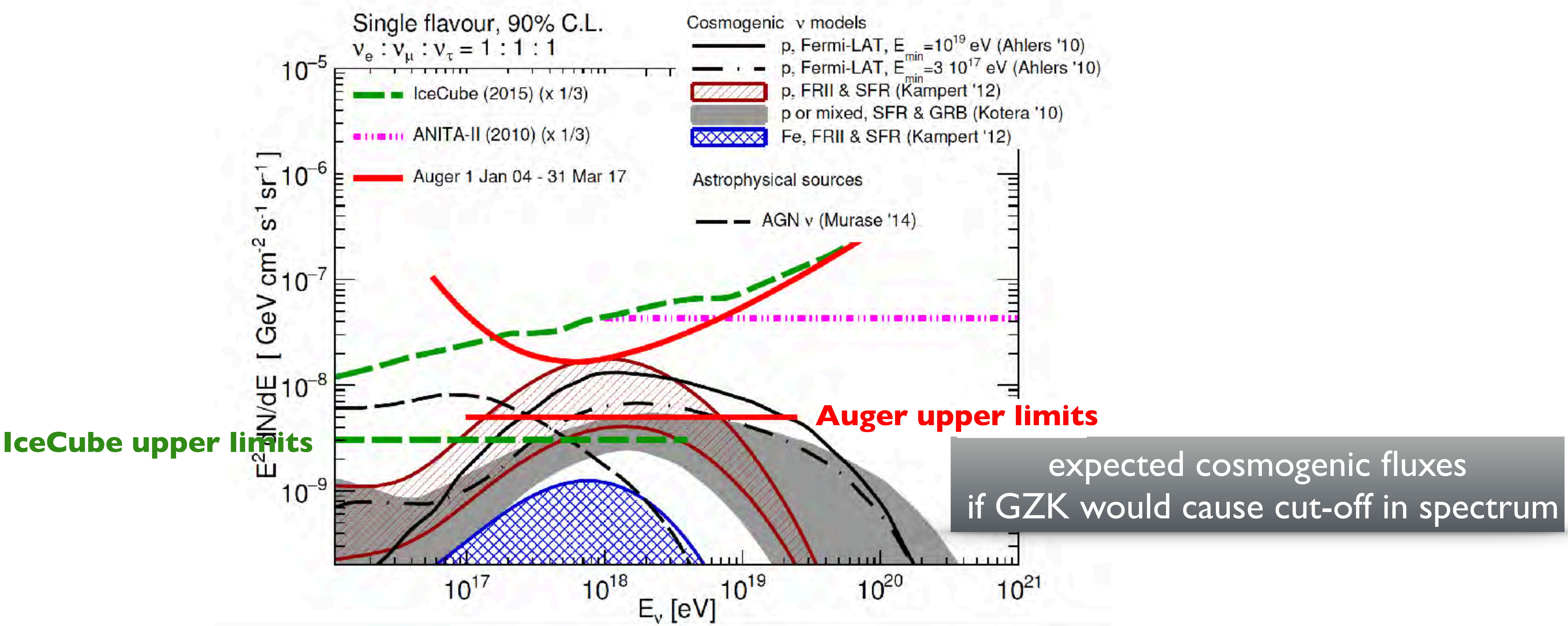






EeV Neutrino Limits challenge GZK

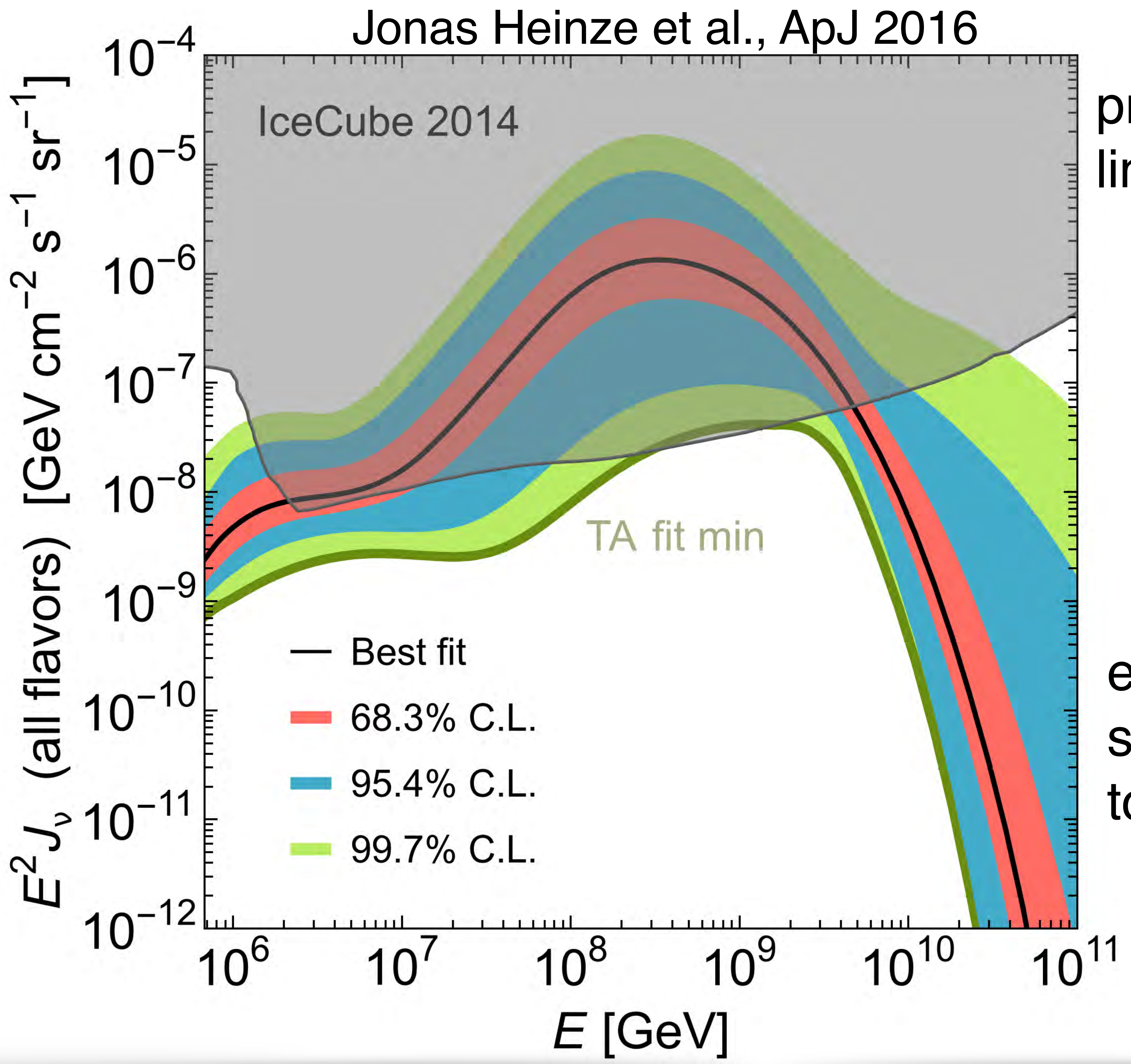
Auger Collaboration, PRD 91, 092008 (2015); update ICRC2017



Would have expected to see 1-7 GZK neutrinos (for different models), have seen none

Neutrino upper limits start to constrain cosmogenic neutrino fluxes of p-sources

EeV Neutrino Limits challenge GZK

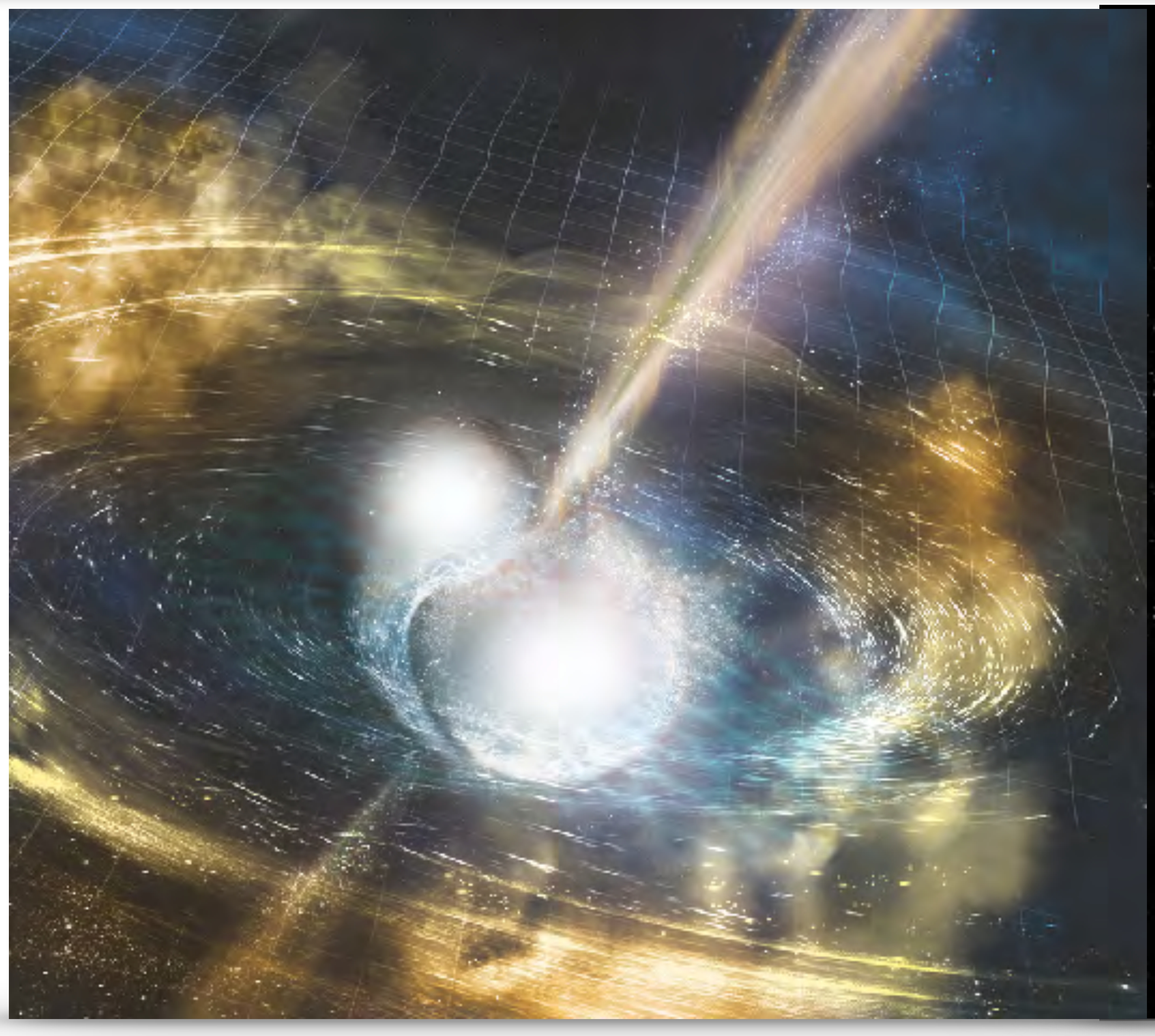


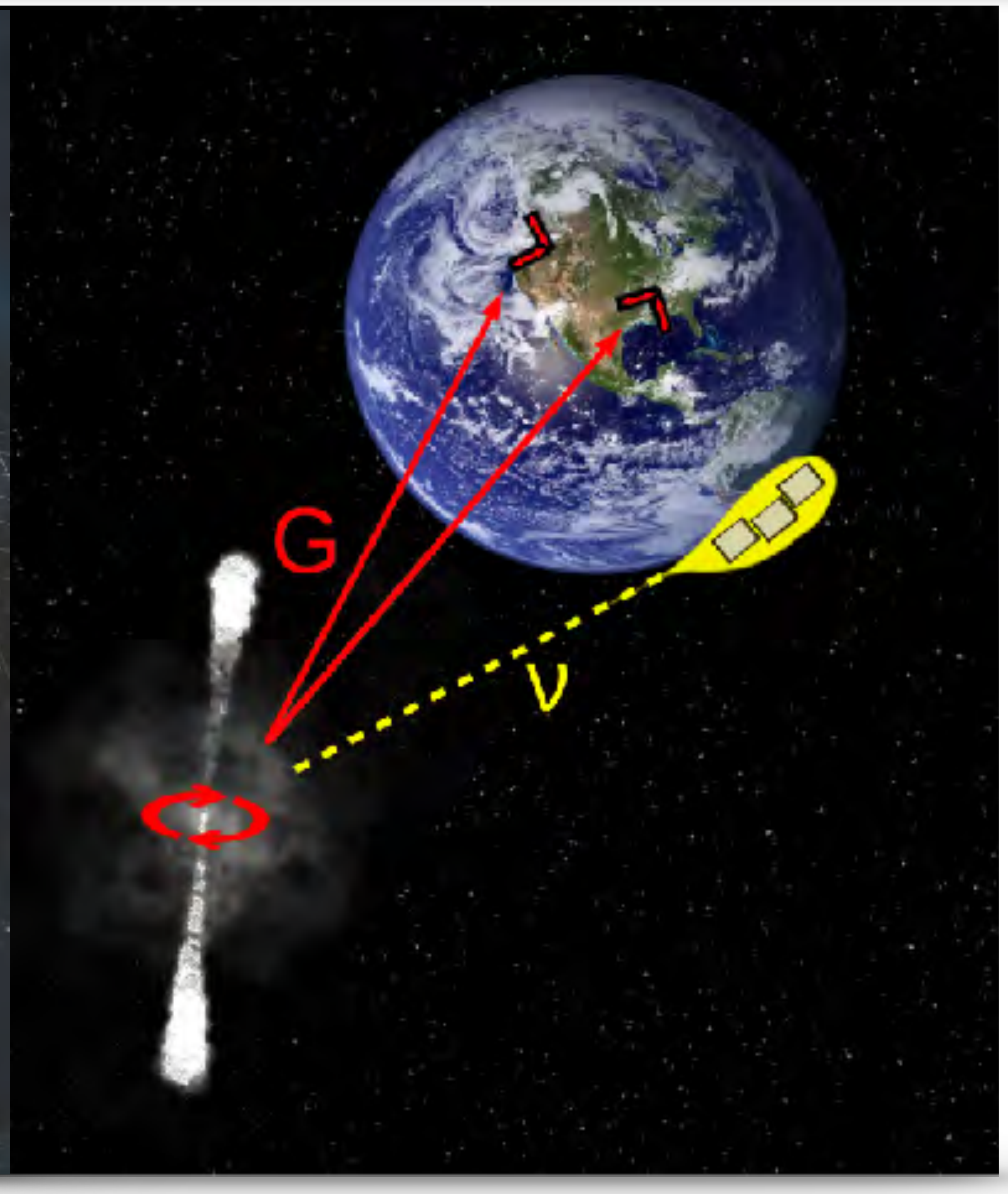
preliminary exclusion limits from IceCube

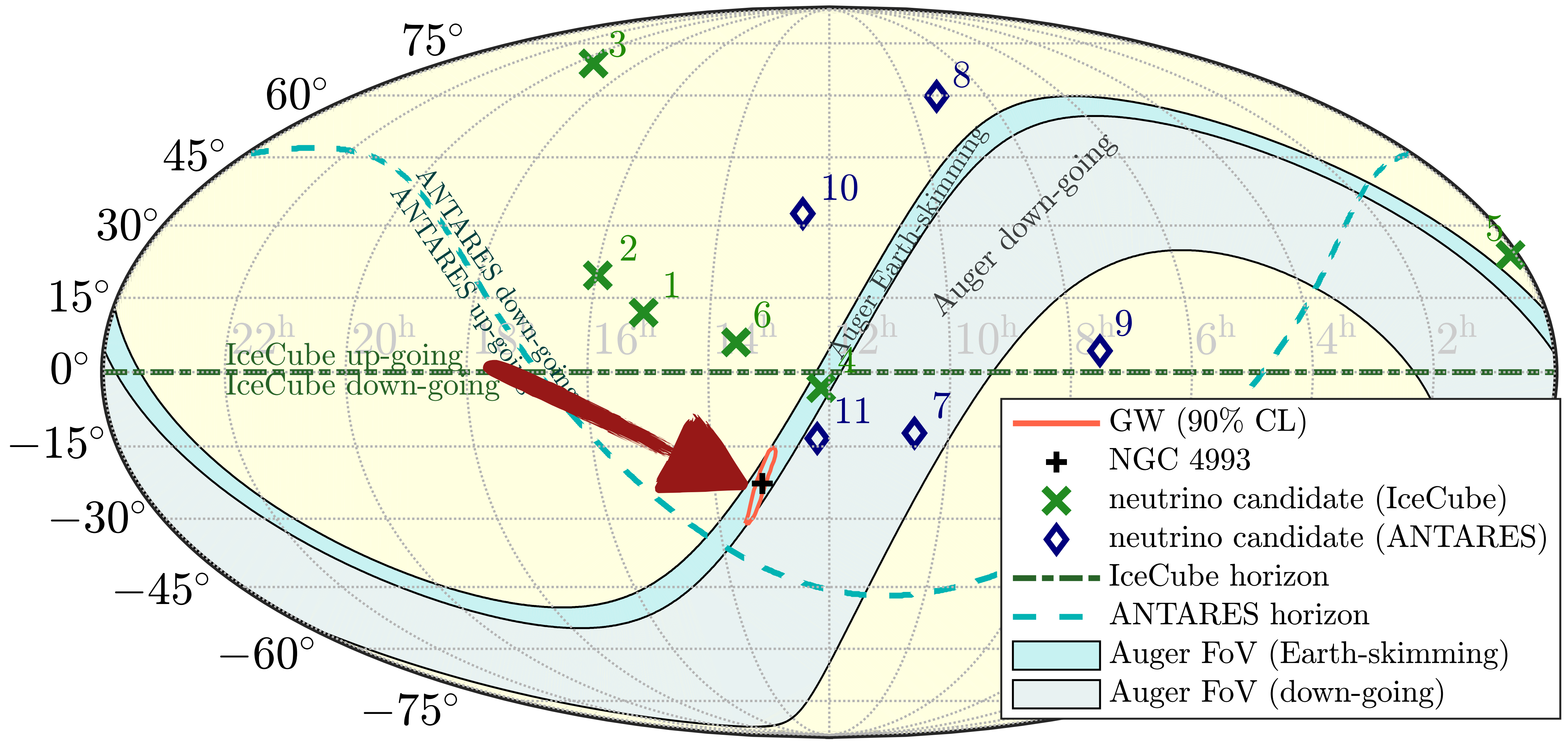
expected ν-fluxes if flux suppression of TA is due to GZK-effect

Pure proton with GZK suppression challenged by upper limits to neutrino fluxes

Neutrino Upper Limits for GW170817





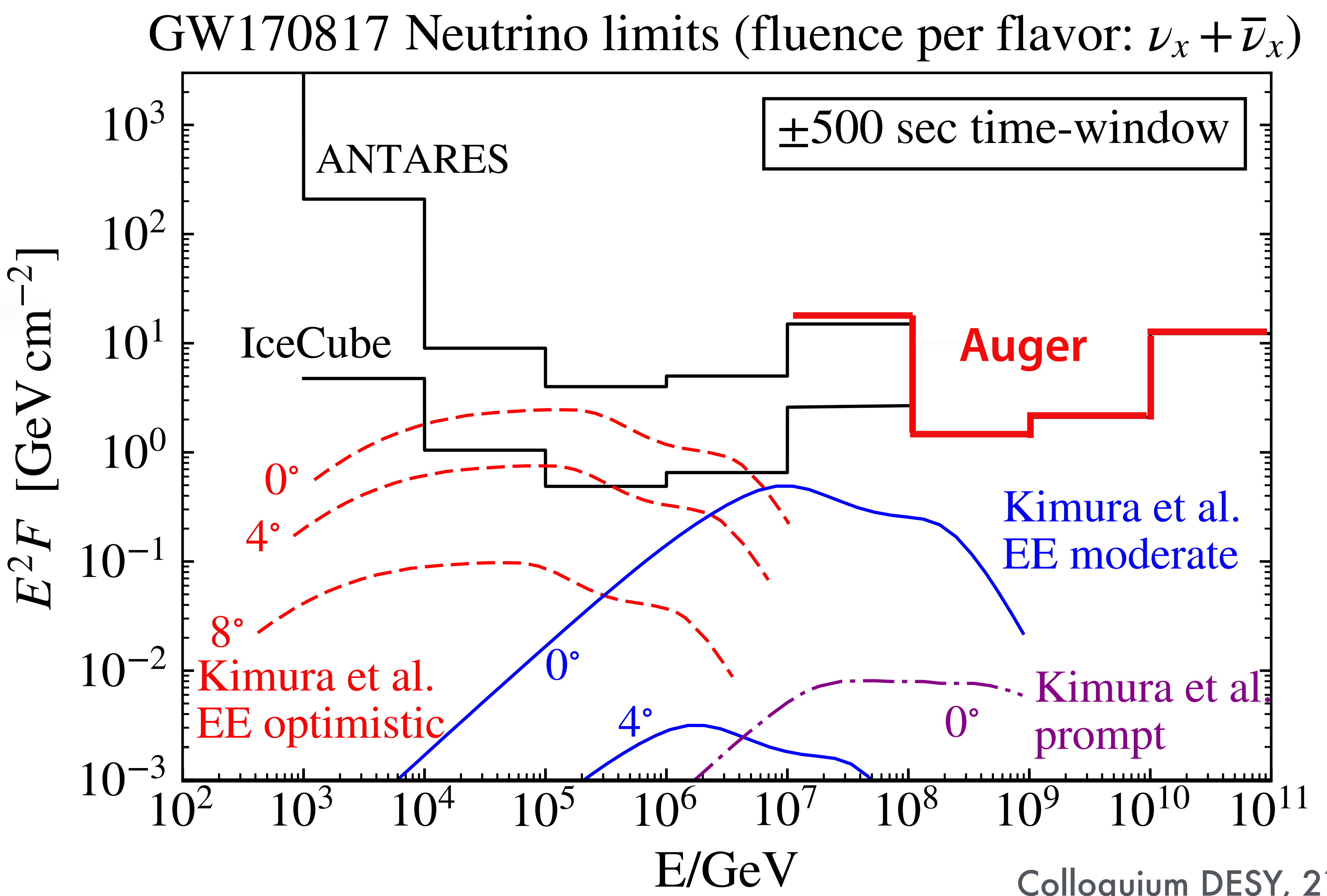


THE ASTROPHYSICAL JOURNAL LETTERS, 848:L12 (59pp), 2017 October 20
https://doi.org/10.3847/2041-8213/aa91c9

Multi-messenger Observations of a Binary Neutron Star Merger

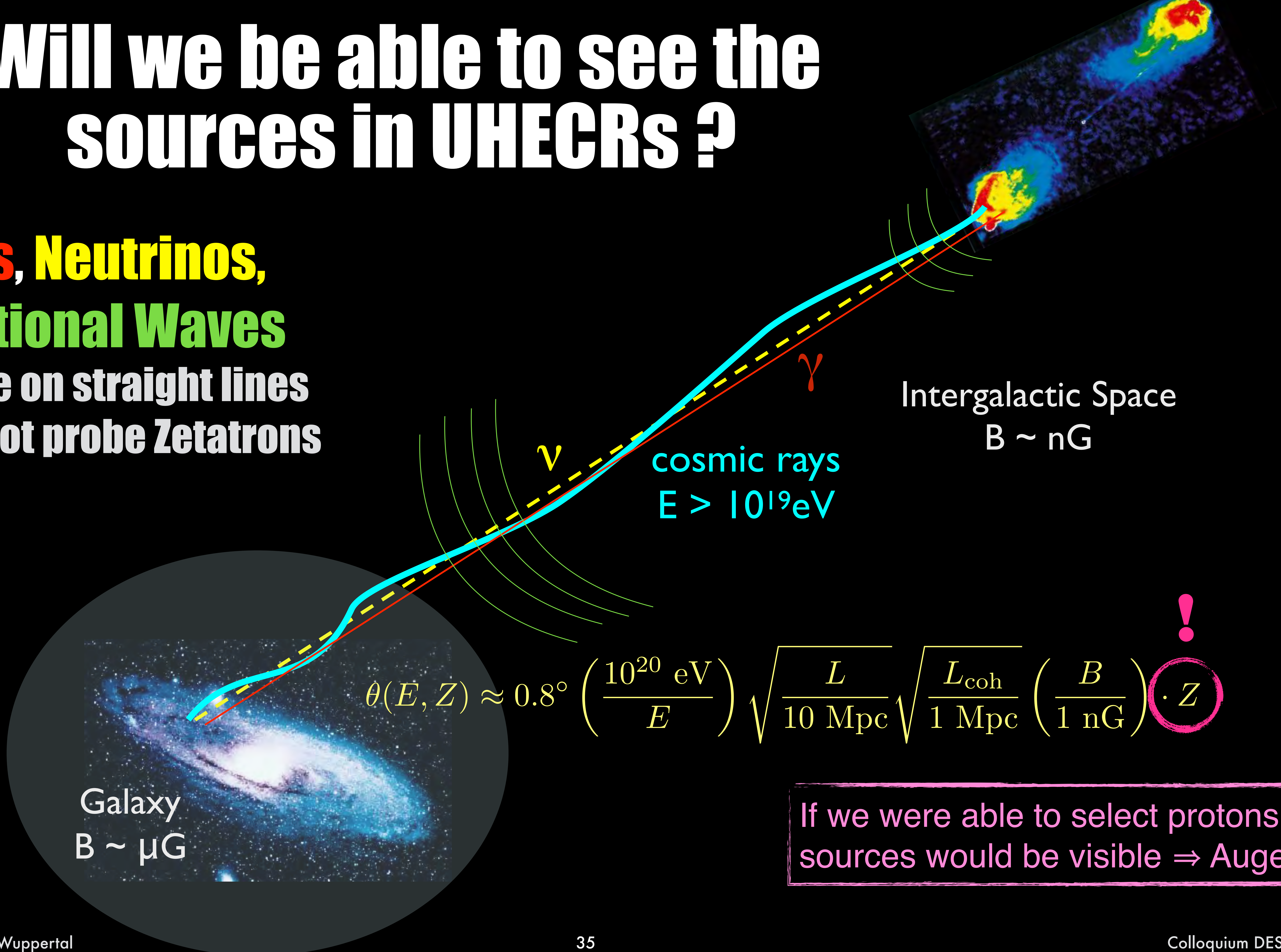
Scientific Breakthrough of 2017:
Birth of Multimessenger Astronomy

LIGO, ANTARES, IceCube, Auger,
The Astrophys. J. Lett. 850 (2017) L35



Will we be able to see the sources in UHECRs ?

Photons, Neutrinos, Gravitational Waves
propagate on straight lines
but may not probe Zetatrons



$$\theta(E,Z) \approx 0.8^\circ \left(\frac{10^{20}\ \mathrm{eV}}{E}\right) \sqrt{\frac{L}{10\ \mathrm{Mpc}}} \sqrt{\frac{L_{\mathrm{coh}}}{1\ \mathrm{Mpc}}} \left(\frac{B}{1\ \mathrm{nG}}\right) \cdot Z$$

!

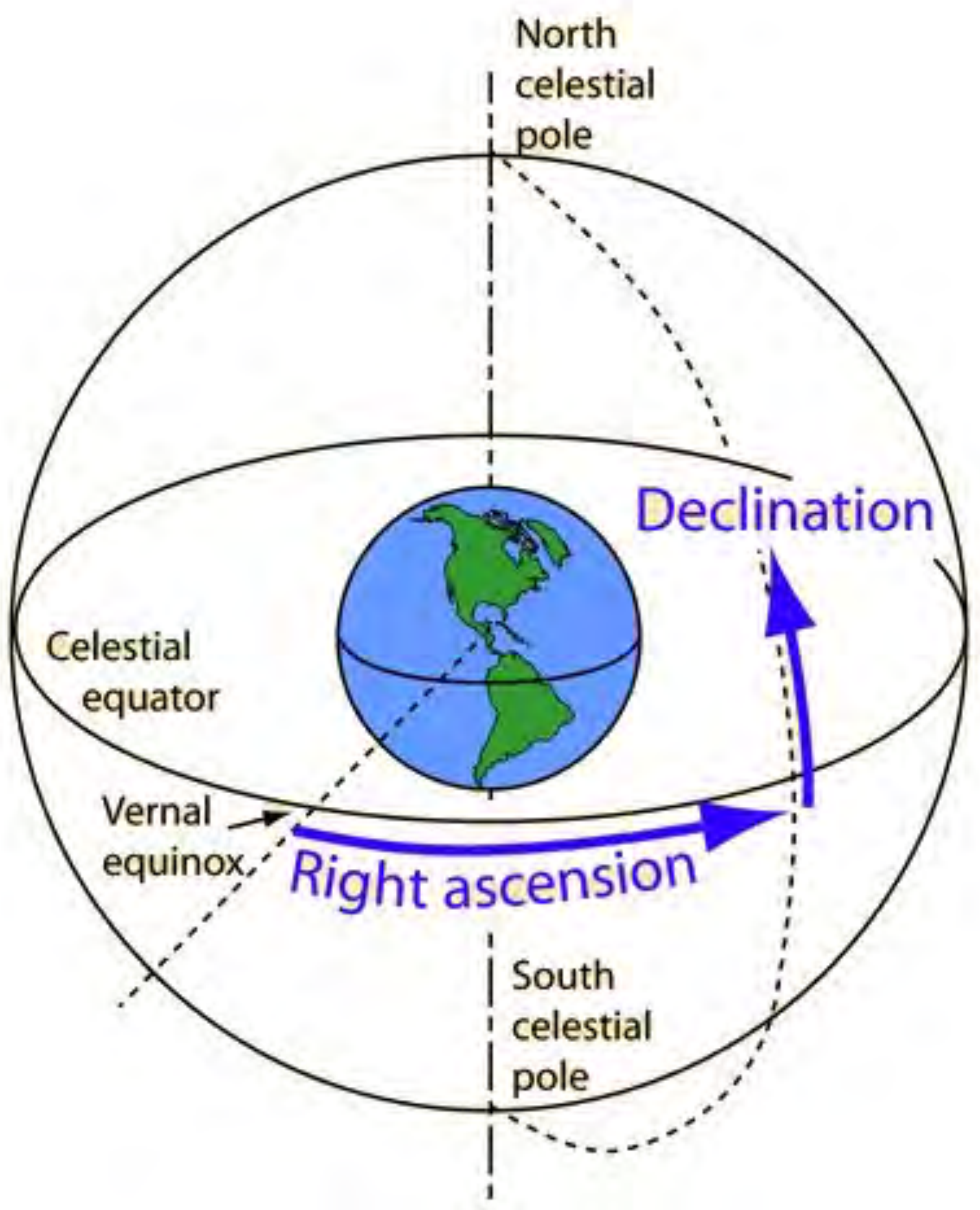
If we were able to select protons, sources would be visible ⇒ AugerPrime

Rayleigh Analysis in Right Ascension α

Fourier coefficients: $a_\alpha = \frac{2}{\mathcal{N}} \sum_{i=1}^{N} w_i \cos\alpha_i \;;\; b_\alpha = \frac{2}{\mathcal{N}} \sum_{i=1}^{N} w_i \sin\alpha_i \;;\; \mathcal{N} = \sum_{i=1}^{N} w_i$

weights w_{i} account for small variations in coverage and tilt of the array

Auger Collaboration, Science 357 (2017) 1266

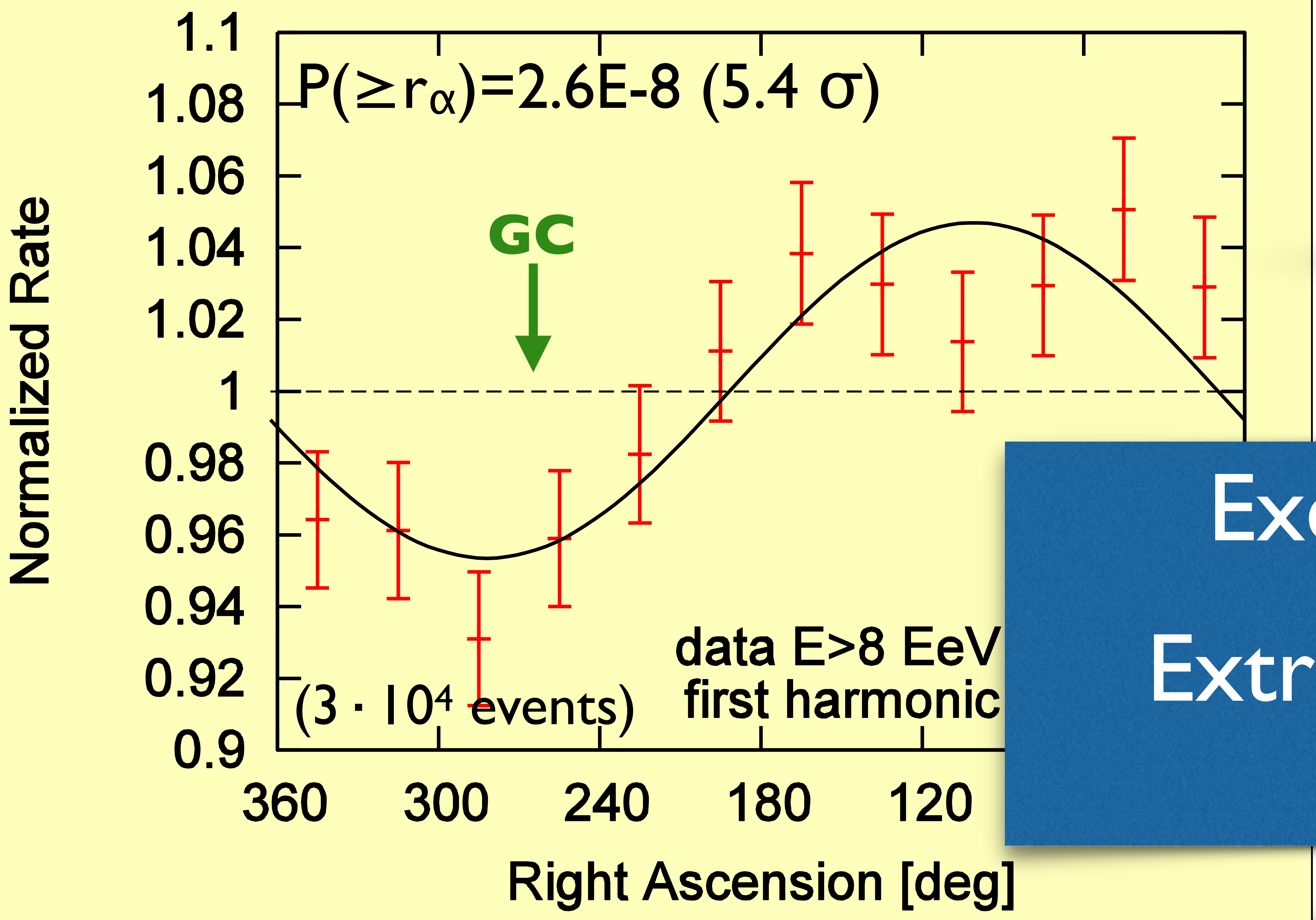


Data compatible with isotropy for 4 ≤ E ≤ 8 EeV

E > 8 EeV

$r_\alpha = \sqrt{a_\alpha^2 + b_\alpha^2}$ amplitude $r_\alpha = 0.047 \pm 0.007$

$\tan\varphi_\alpha = \frac{b_\alpha}{a_\alpha}$ phase $\varphi_\alpha = 100° \pm 10°$



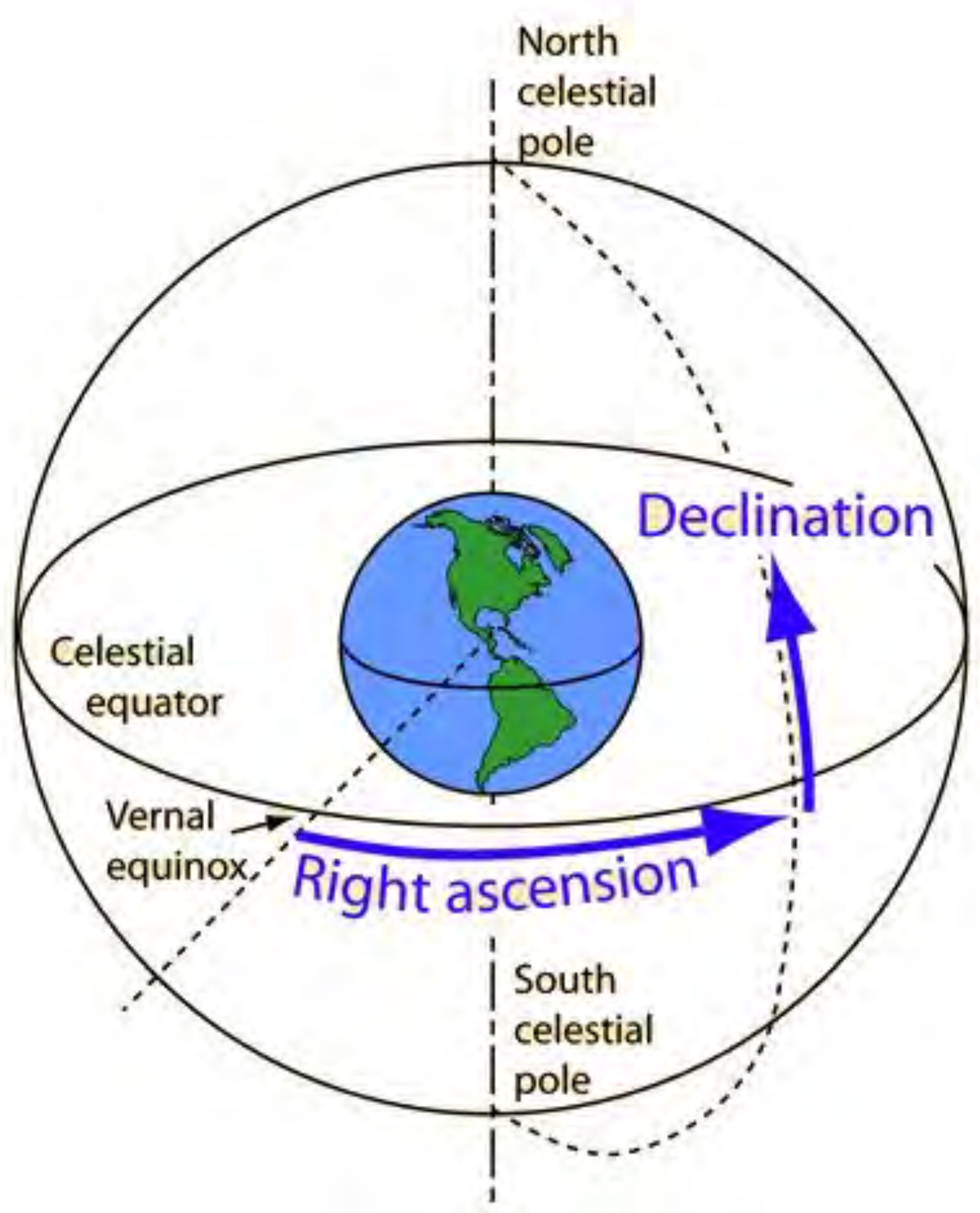
Excess direction ⇒

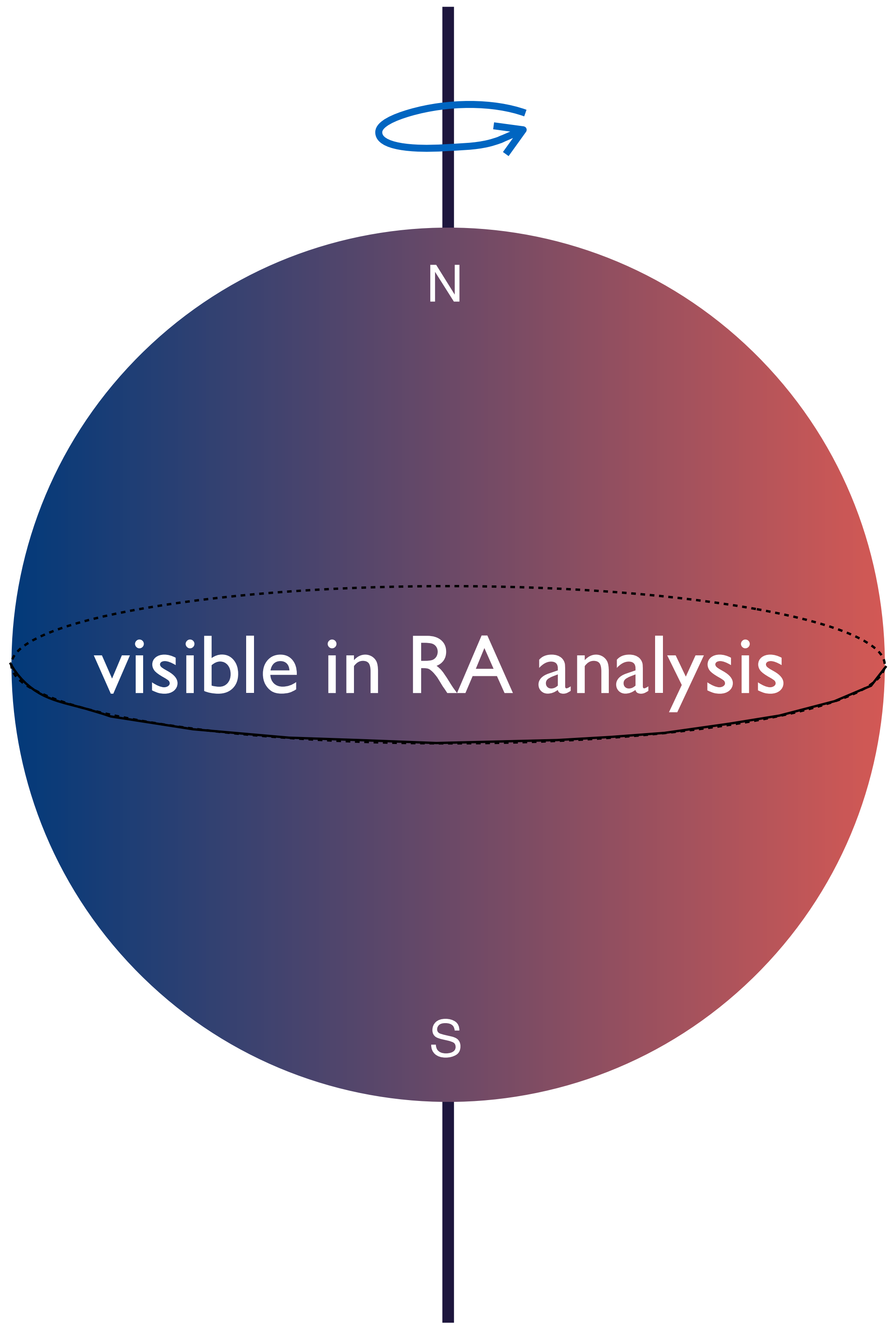
Extragalactic Origin of UHECR

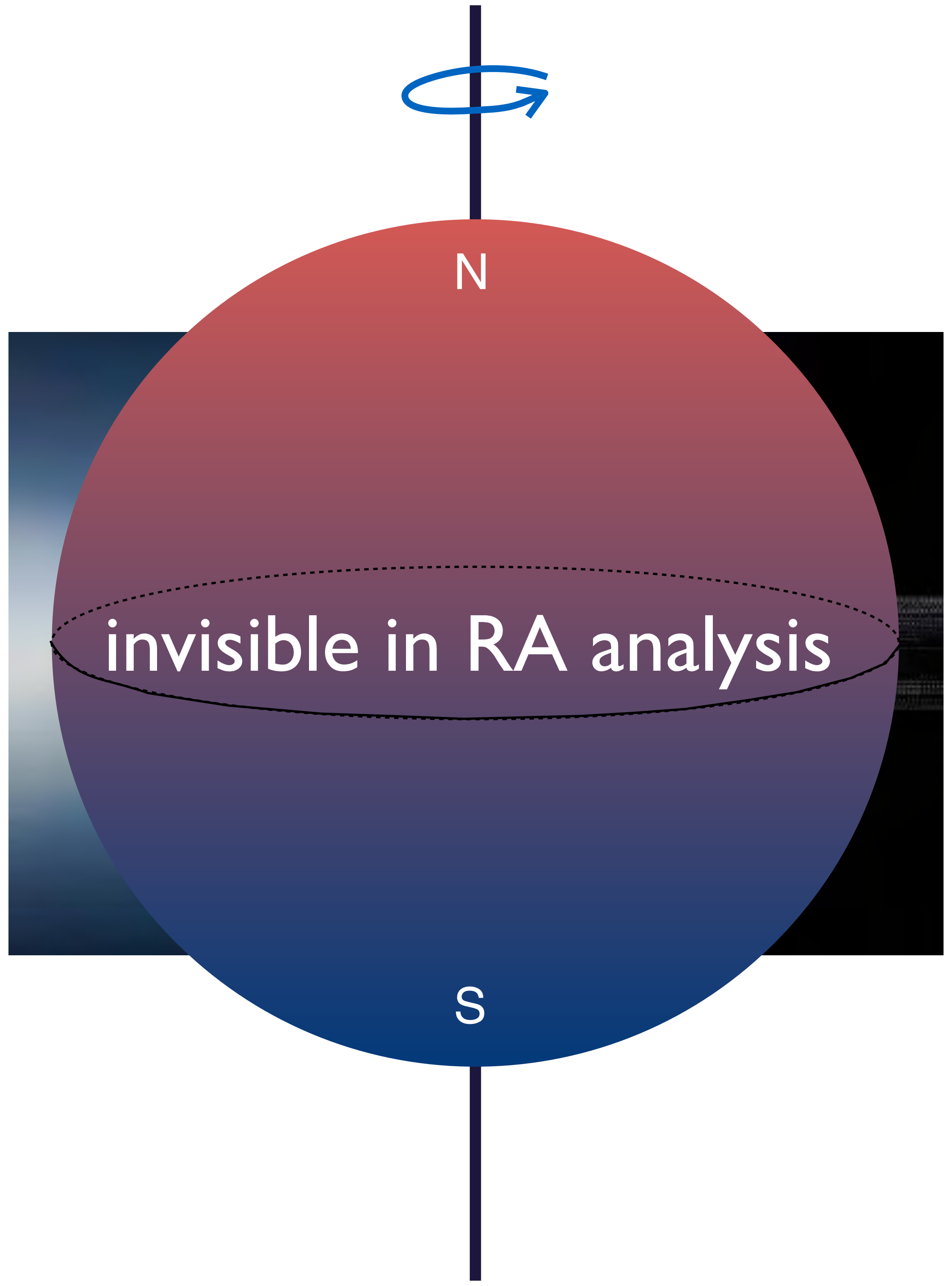
Reconstruction of 3-dim Dipole

Fourier coefficients: $$a_\alpha = \frac{2}{\mathcal{N}} \sum_{i=1}^{N} w_i \cos\alpha_i \;\; ; \;\; b_\alpha = \frac{2}{\mathcal{N}} \sum_{i=1}^{N} w_i \sin\alpha_i \;\; ; \;\; \mathcal{N} = \sum_{i=1}^{N} w_i$$

weights w_i account for small variations in coverage and tilt of the array



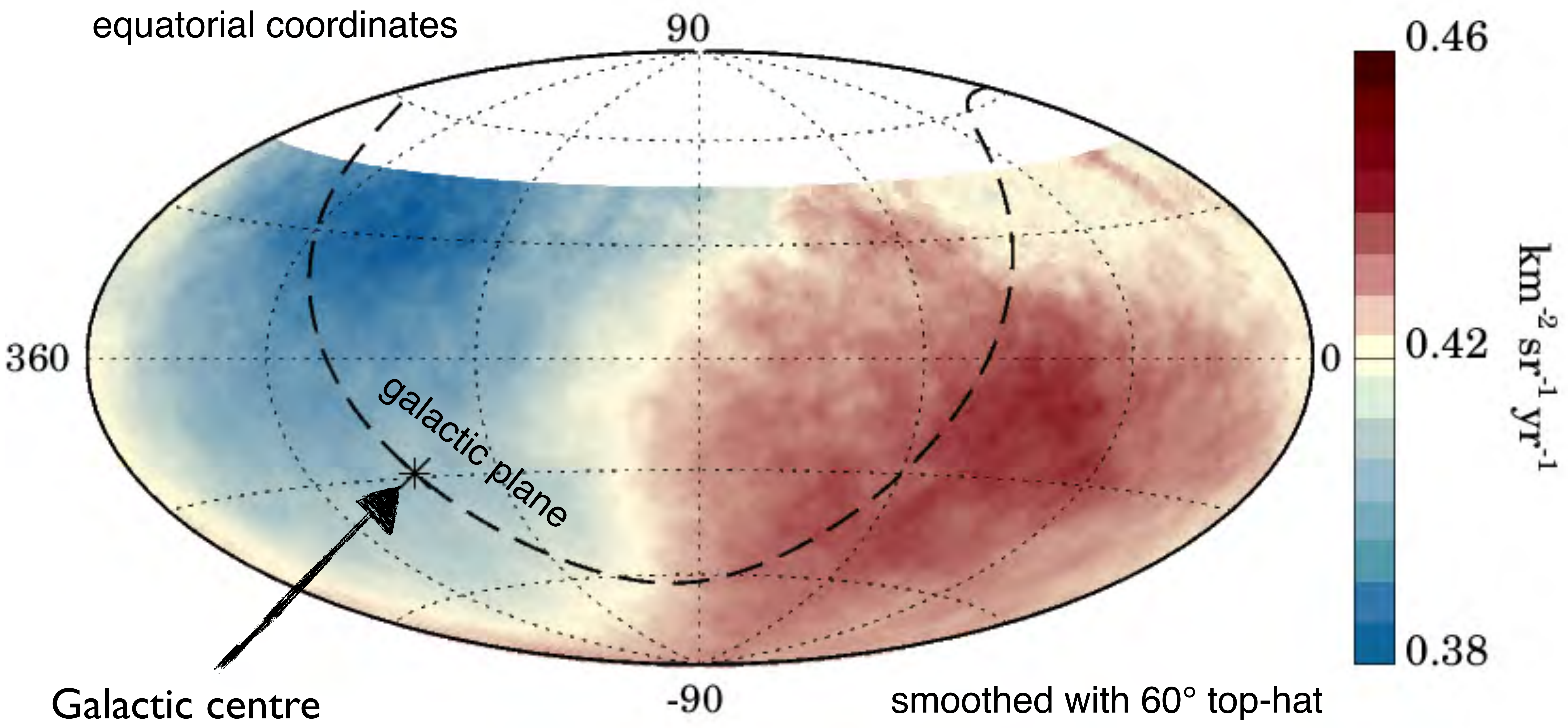




➠ combine harmonic analysis in RA with harmonic in North-South

Flux Map above 8 EeV

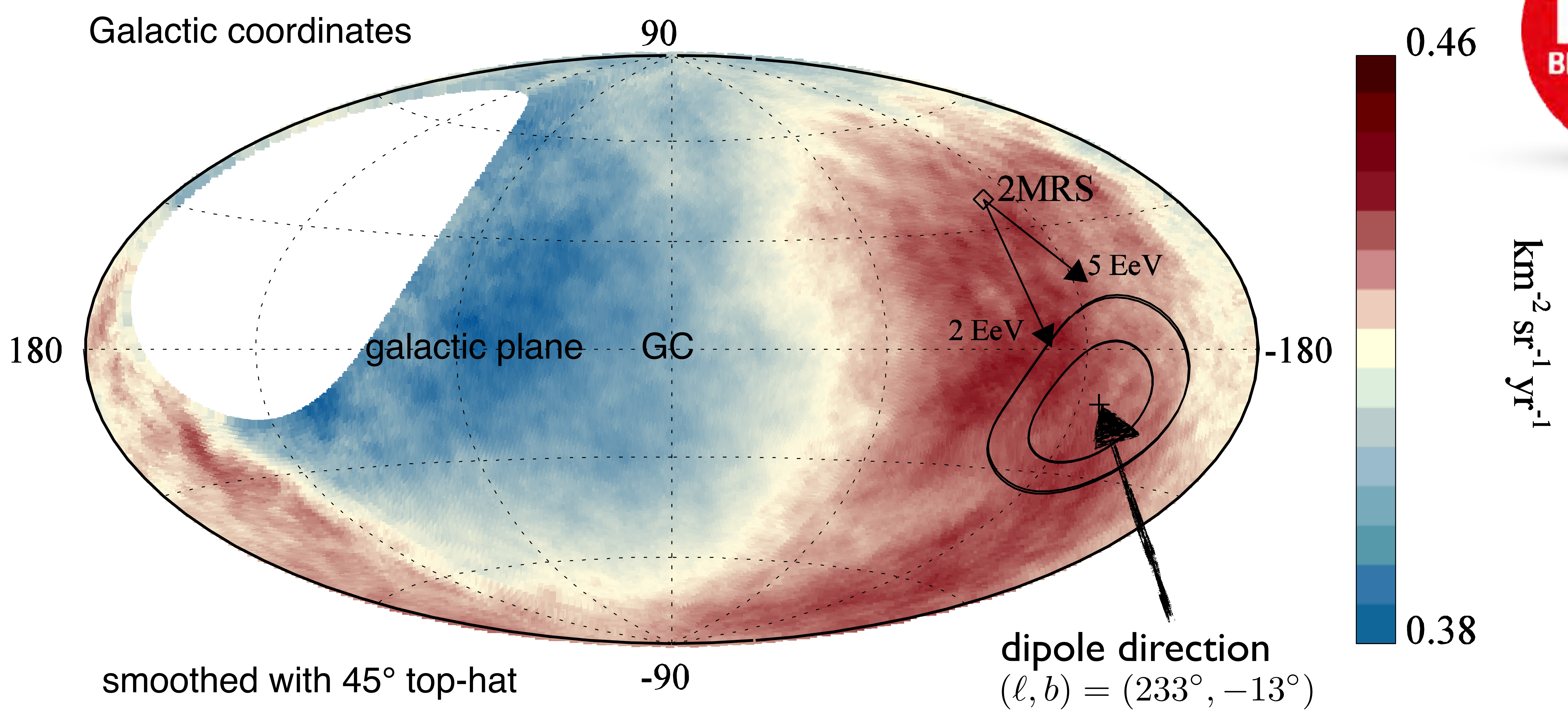
Auger Collaboration, Science 357 (2017) 1266



$$\mathcal{A} = 6.5^{+1.3}_{-0.9}\,\% \;\; ; \;\; \alpha_d = (100 \pm 10)^\circ \;\; ; \;\; \delta_d = (-24^{+12}_{-13})^\circ$$

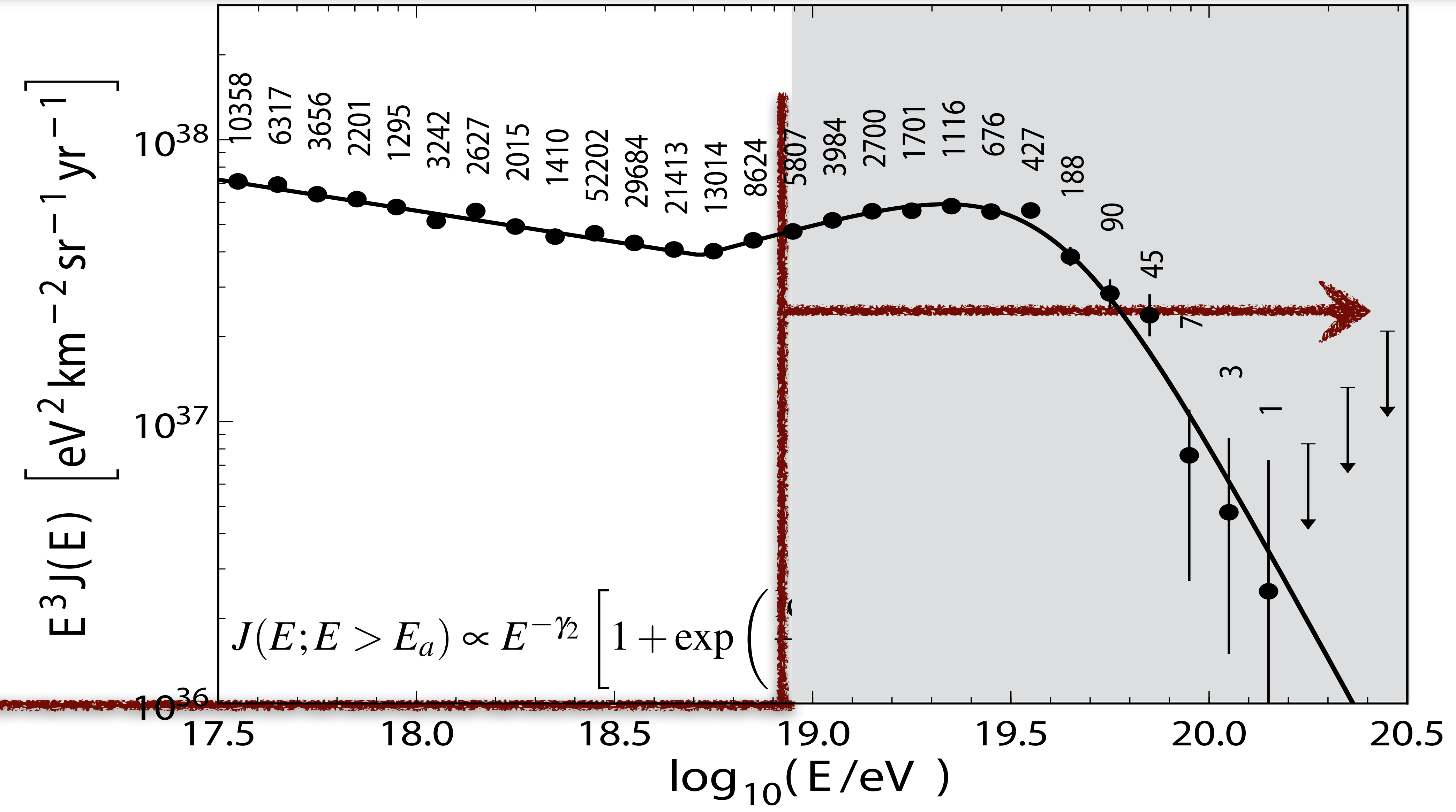
Flux Map above 8 EeV

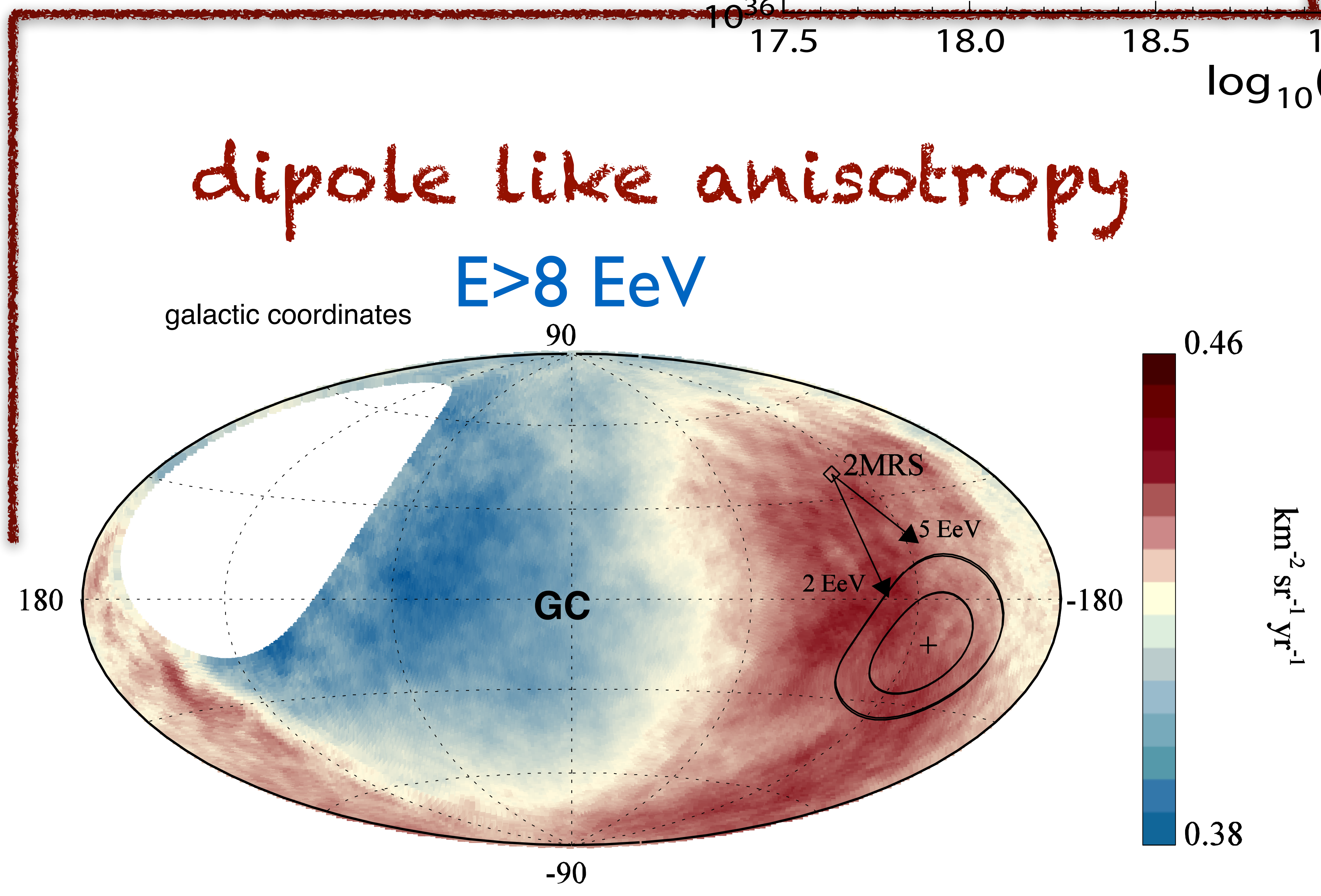
Auger Collaboration, Science 357 (2017) 1266

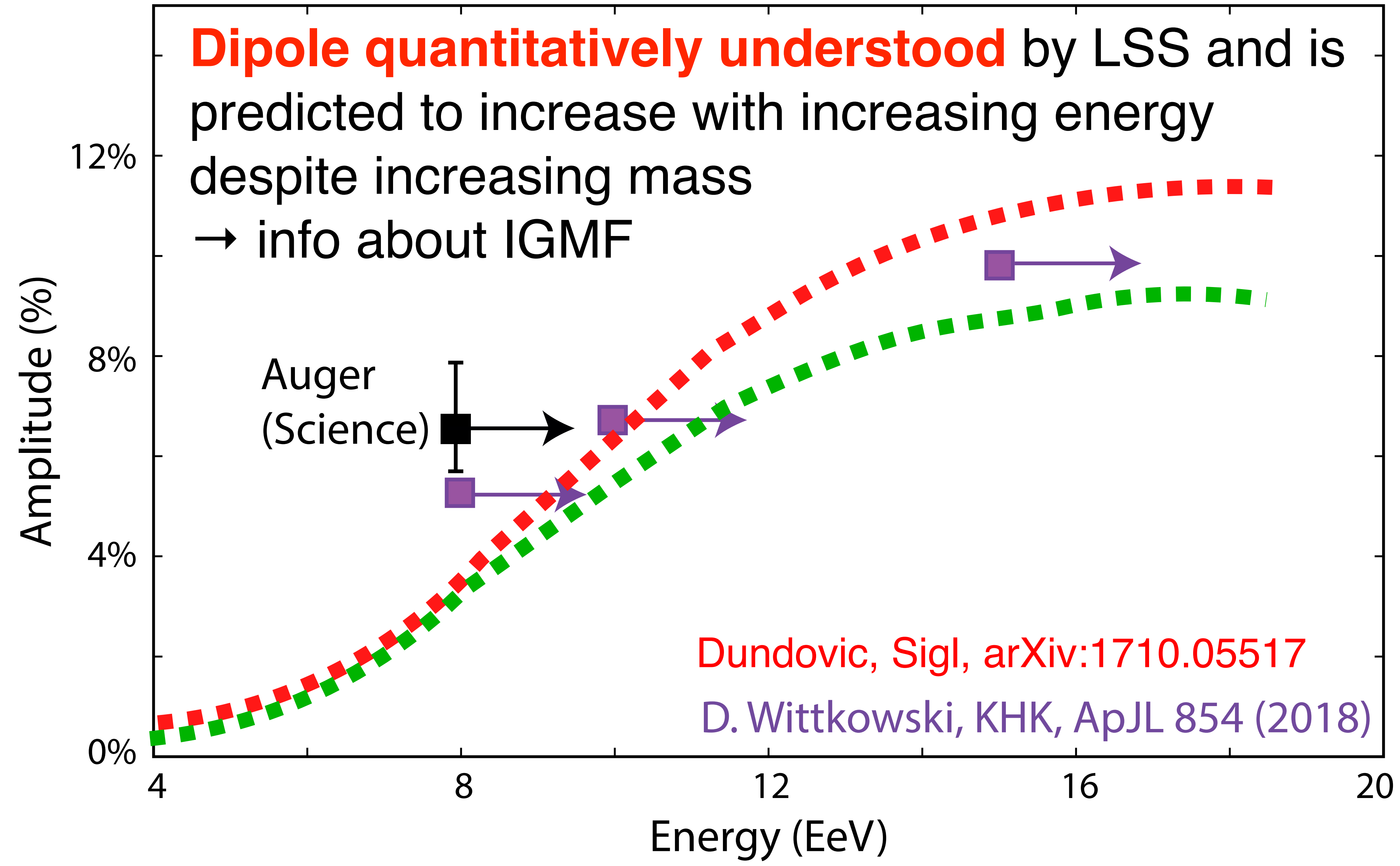


Extragalactic origin of UHECR confirmed

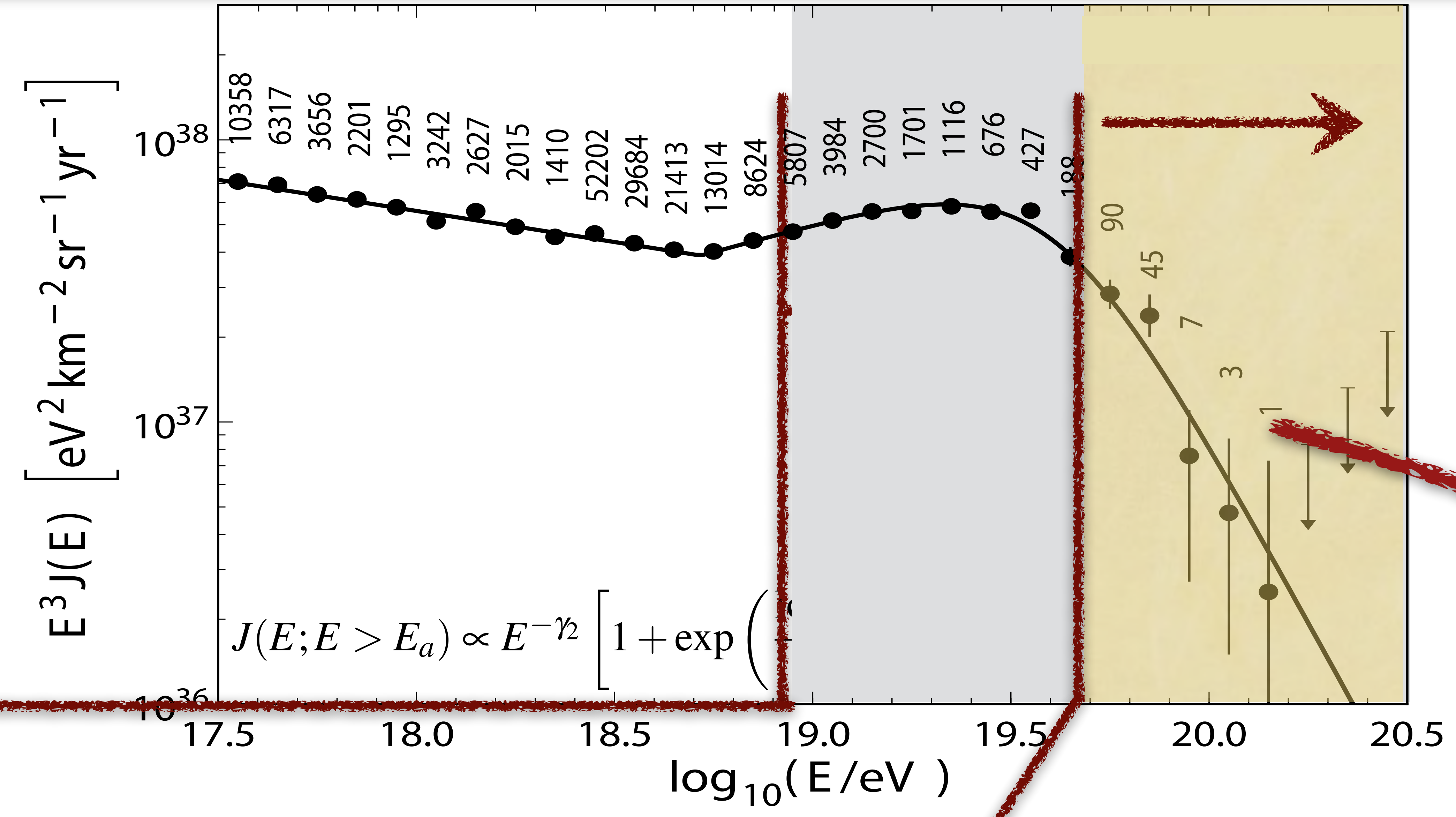
Detection of UHECR Anisotropies





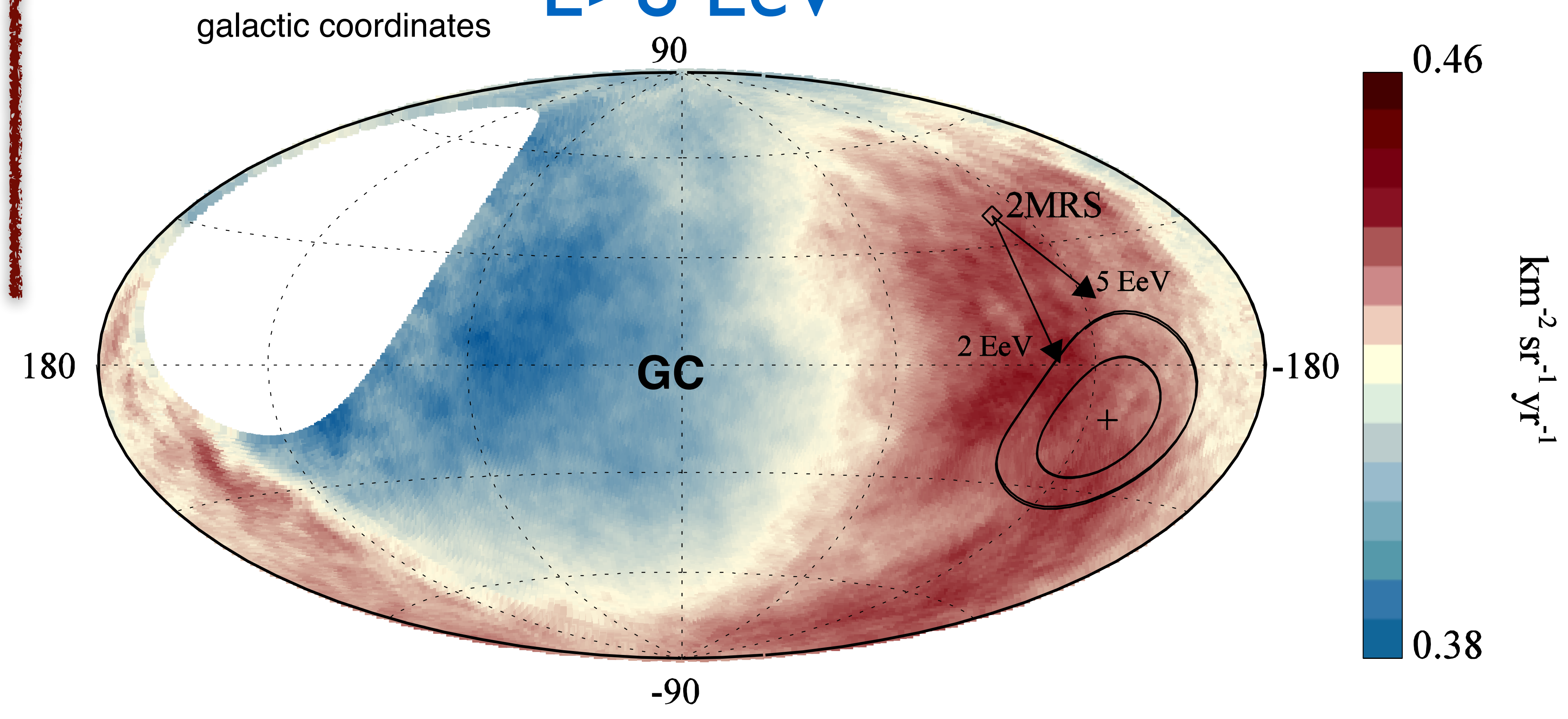


Detection of UHECR Anisotropies



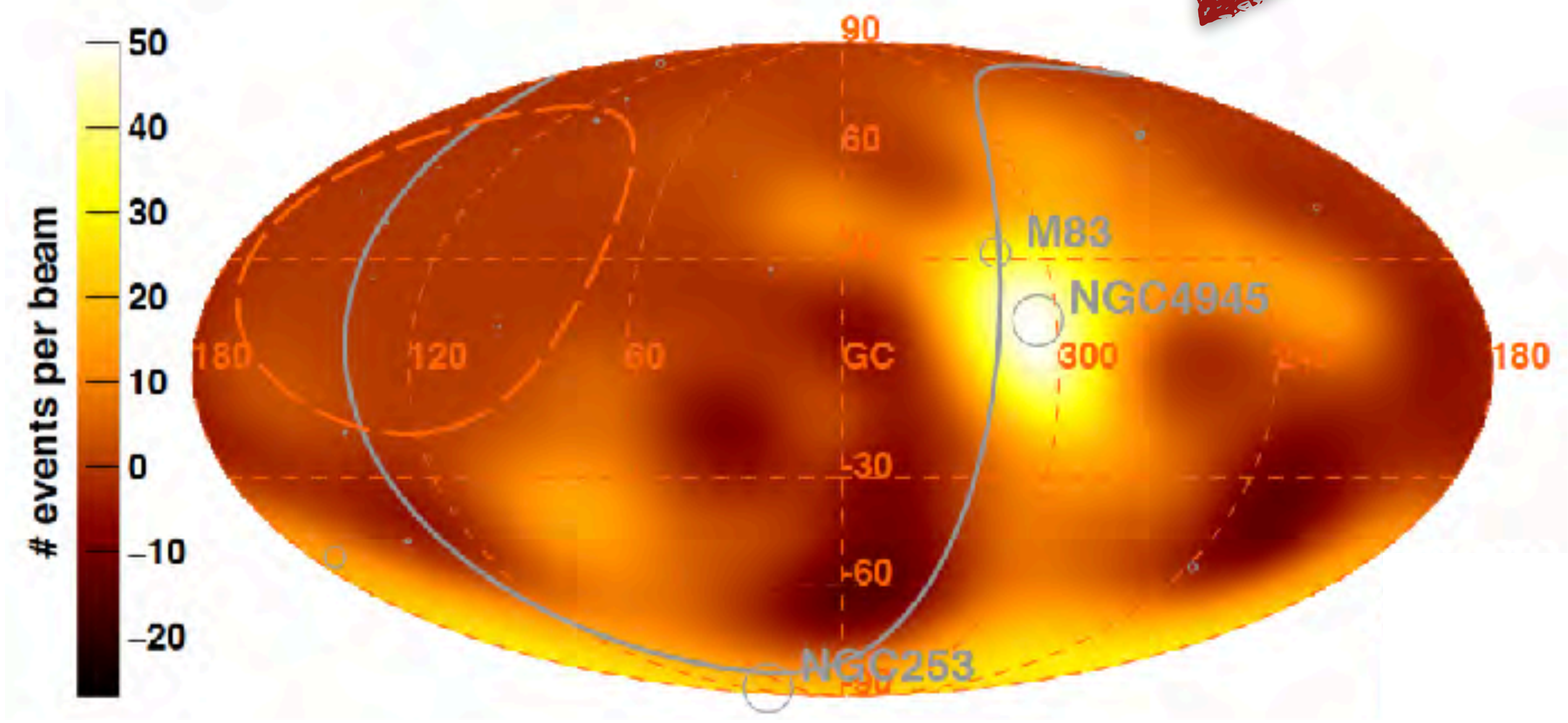
dipole like anisotropy

E>8 EeV



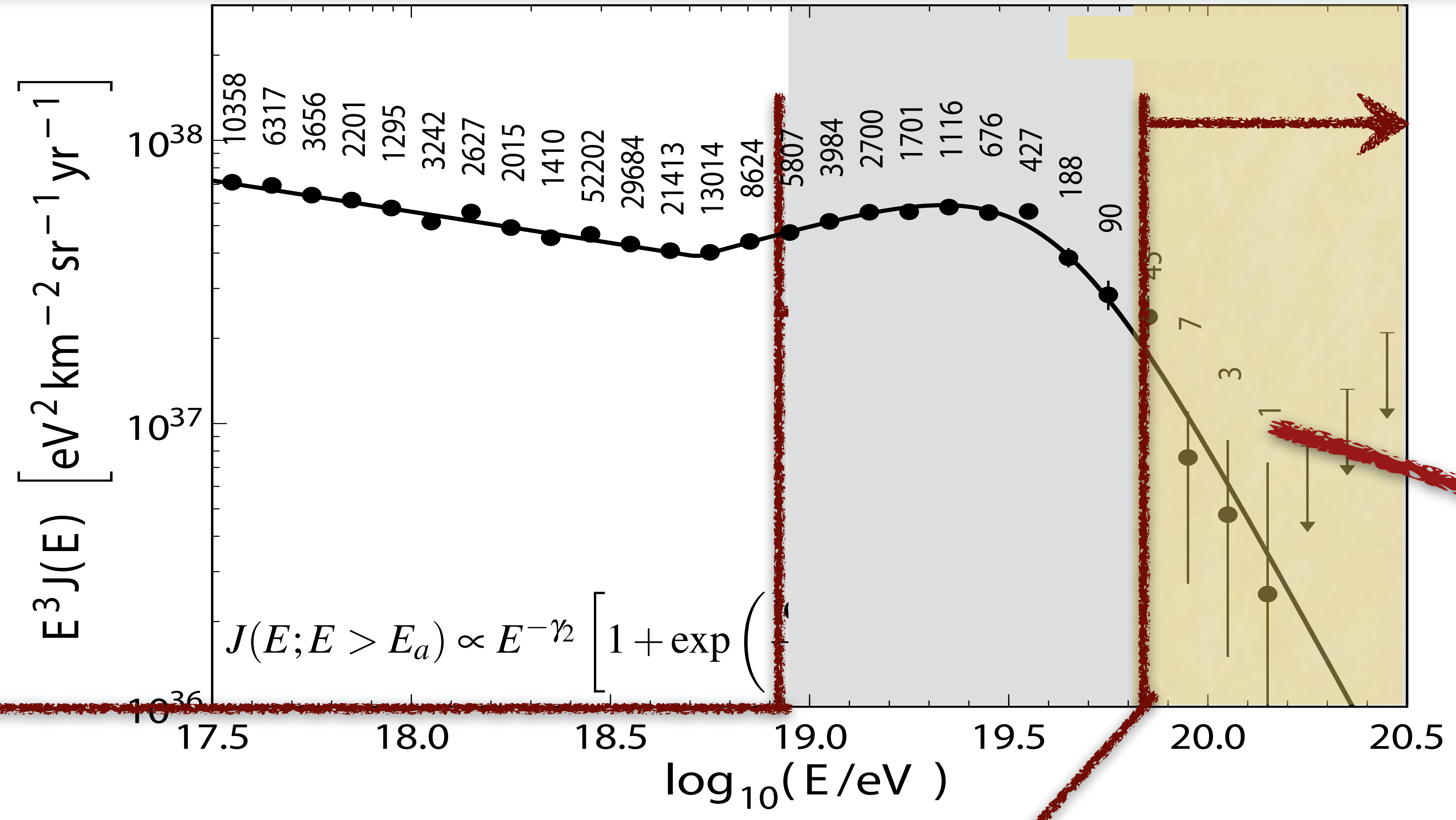
more structures

E>39 EeV



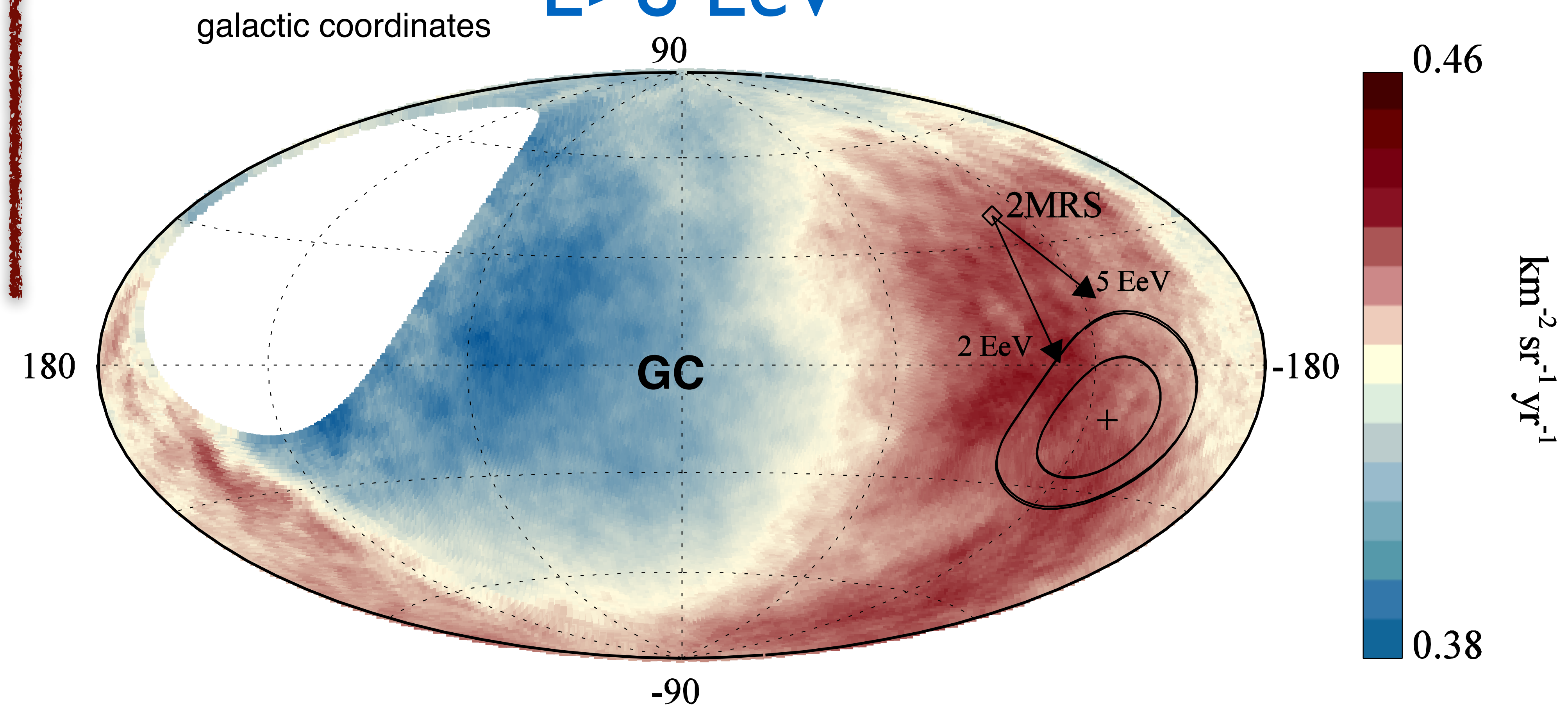
Auger: ApJL 853:L29 (2018)

Detection of UHECR Anisotropies



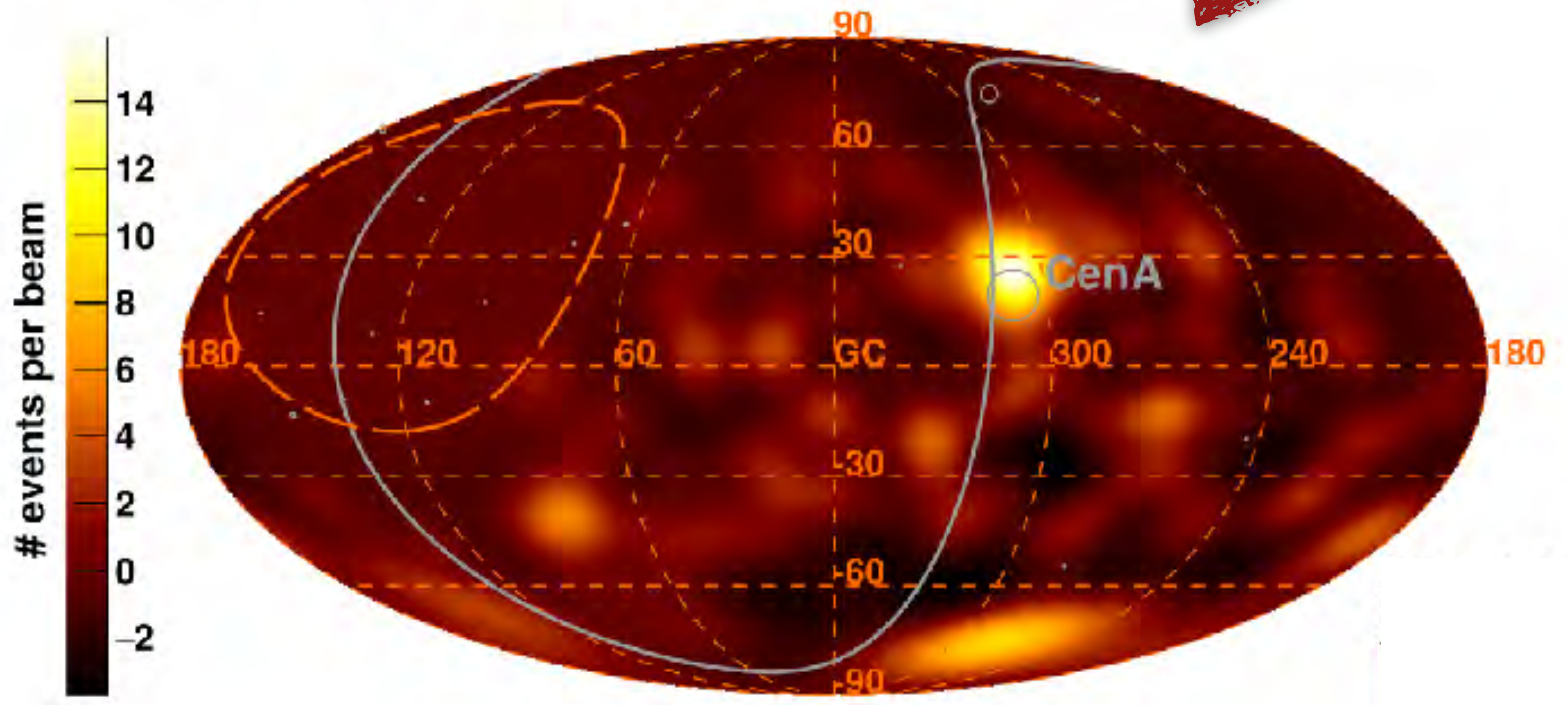
dipole like anisotropy

E>8 EeV



even more structures

E>60 EeV

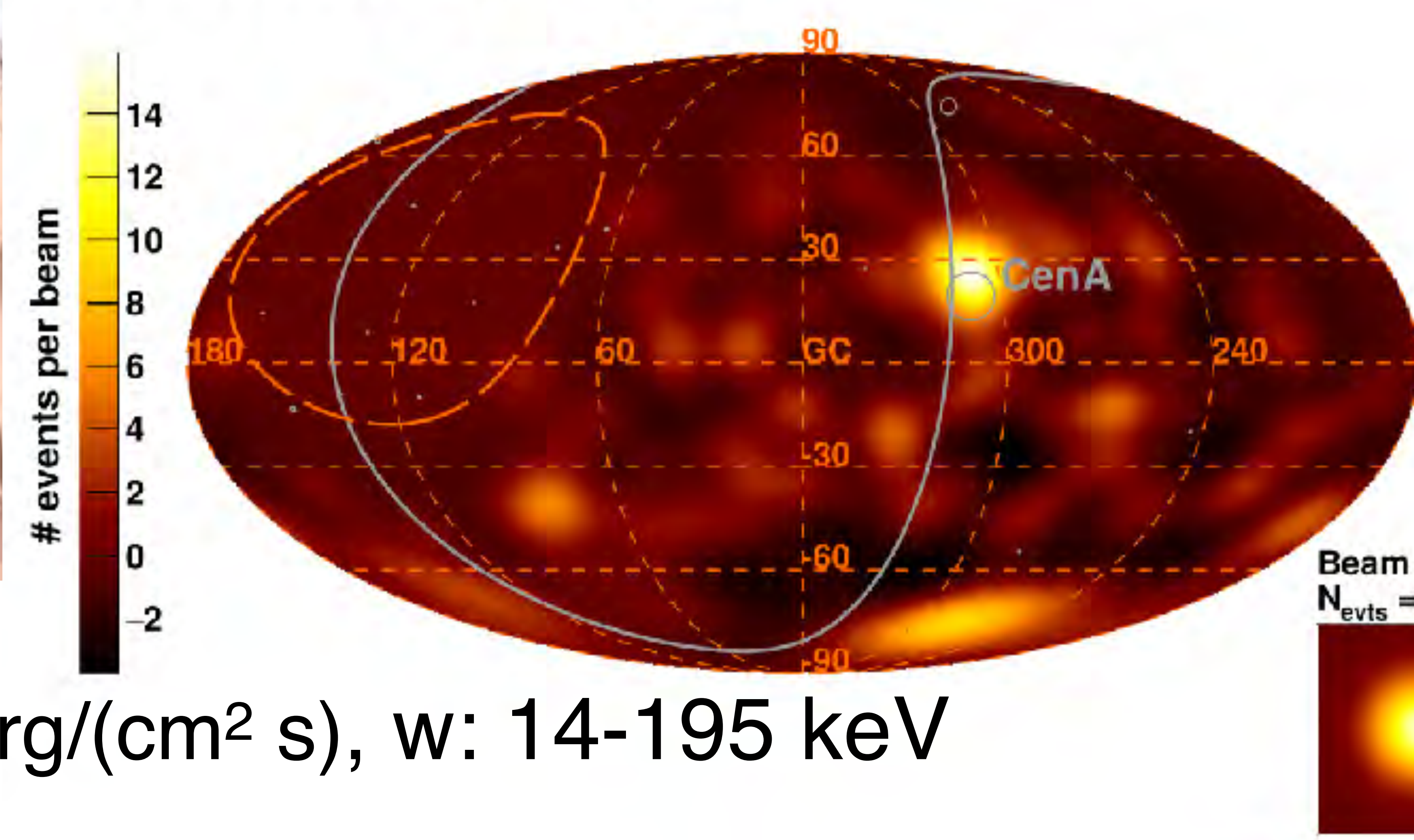


Auger: ApJL 853:L29 (2018)

The Usual UHECR Source Suspects

Swift-BAT 2MRS SGB γAGN ~~VCV~~

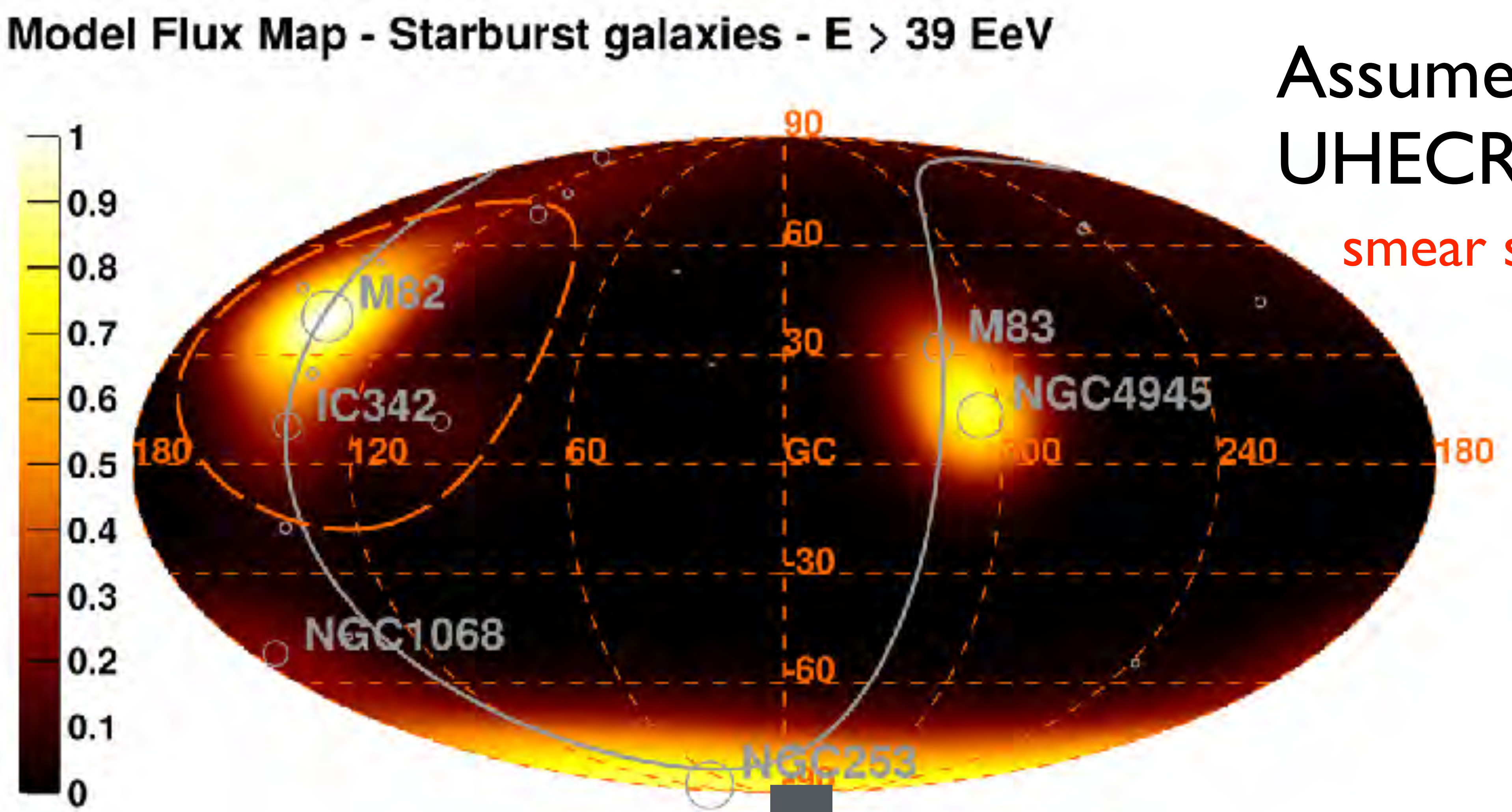
Adapted from M. Unger



- Swift-BAT X-ray-selected galaxies, D < 250 Mpc, $\Phi > 1{:}3\times10^{-11}$ erg/(cm² s), w: 14-195 keV
- 2MRS IR-selected galaxies, D > 1 Mpc, w: K-band
- SBG: 23 nearby starburst galaxies, $\Phi > 0.3$ Jy, w: radio at 1.4 GHz
- γAGN: 17 2FHL blazars and radio galaxies, D < 250 Mpc, w: γ-ray 50 GeV - 2 TeV

Understanding the UHECR Sky

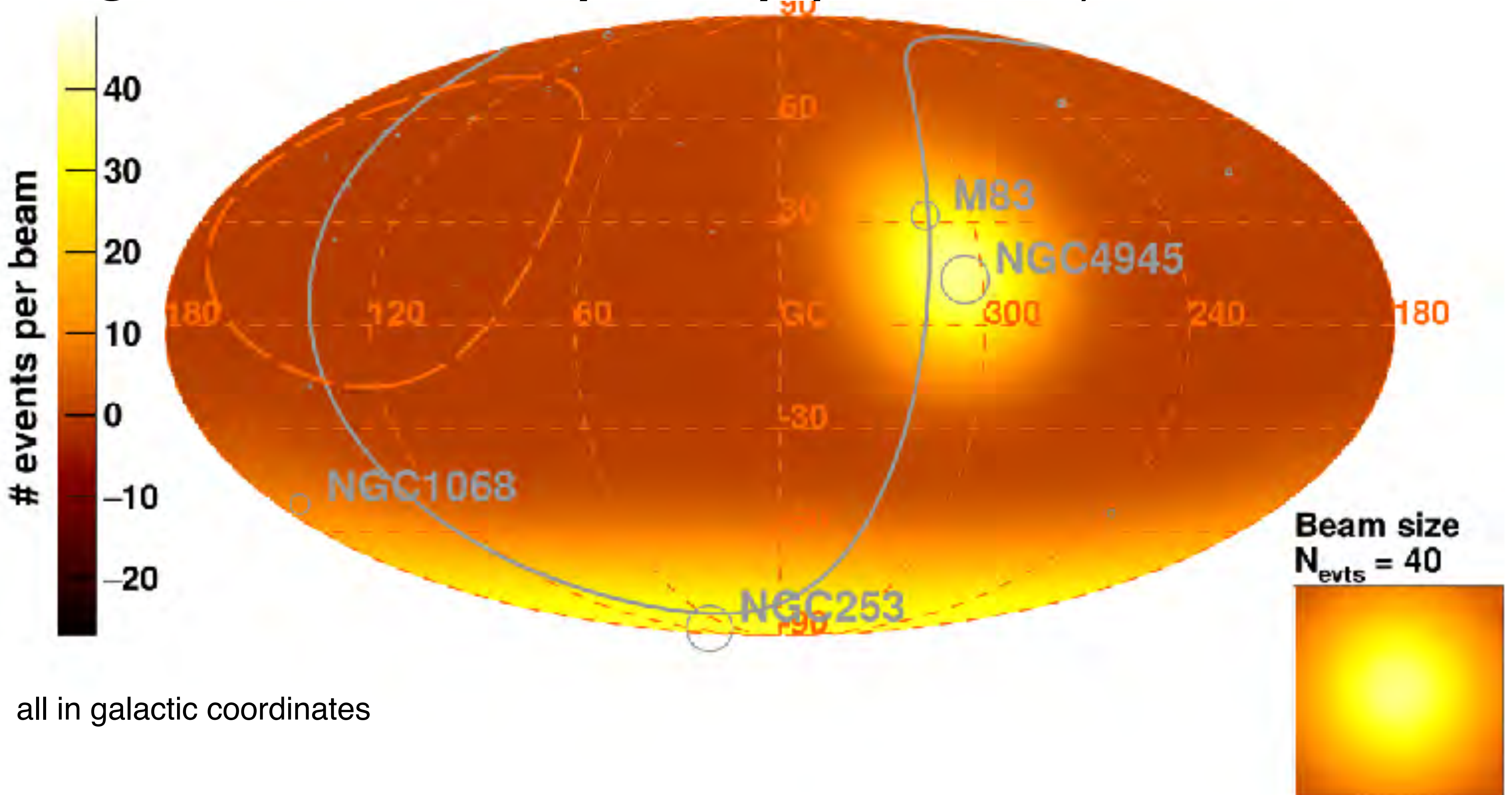
Starburst Galaxy Model



Assume, starburst galaxies produce UHECR with $L_{UHECR} \sim L_{\gamma}$ @ 1.4 GHz

smear sources to account for B-field deflections

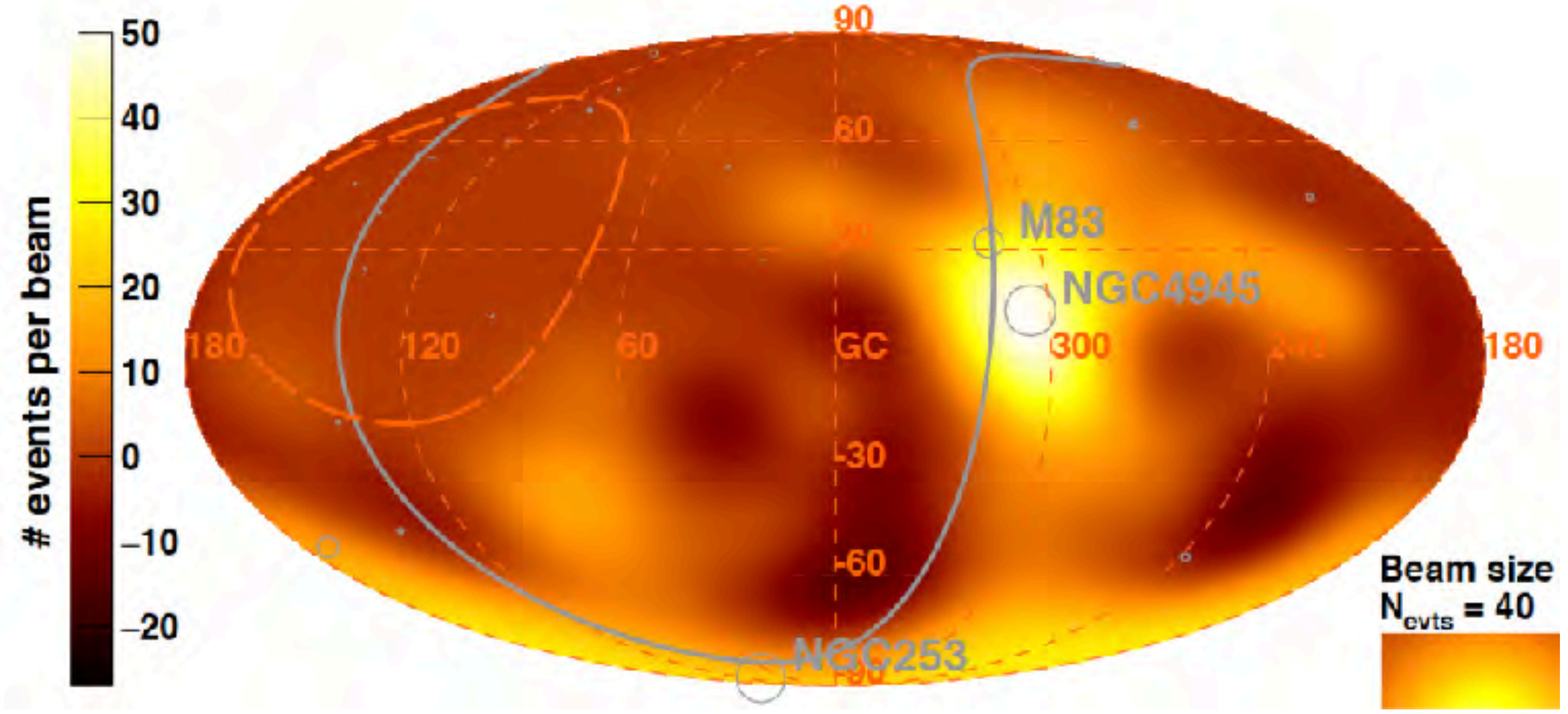
Add isotropic background (allow background sources a/o larger deflections by heavy primaries), such that model map…



all in galactic coordinates

…maximises degree of correlation with observed UHECR sky

Auger data map at E>39 EeV



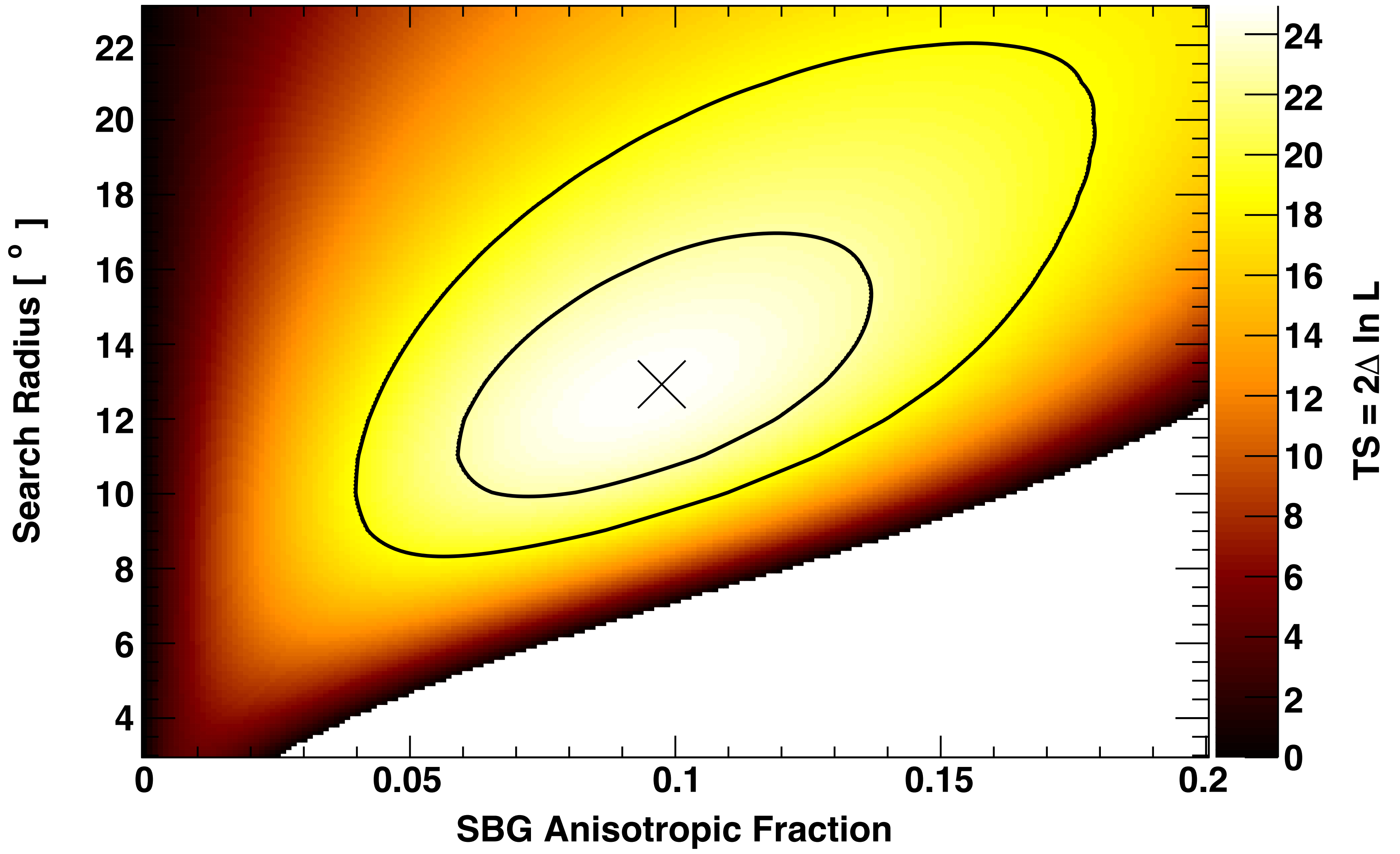
Sources assumed to emit UHECR spectrum and composition according to results from combined fit.
Propagation effects (attenuation) fully accounted for.

Auger: ApJL 853:L29 (2018)

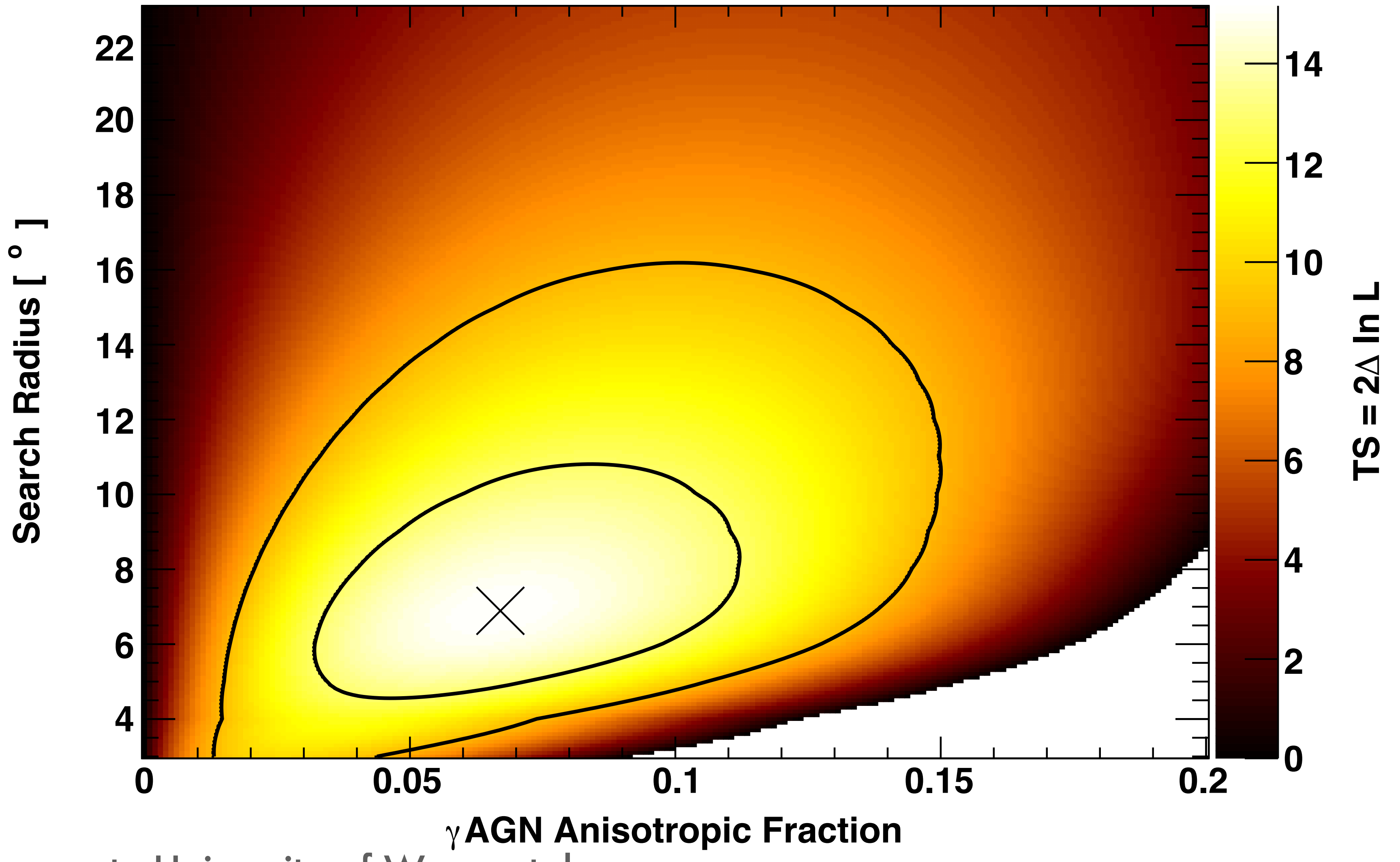
Test Statistics & 2D-Profiles

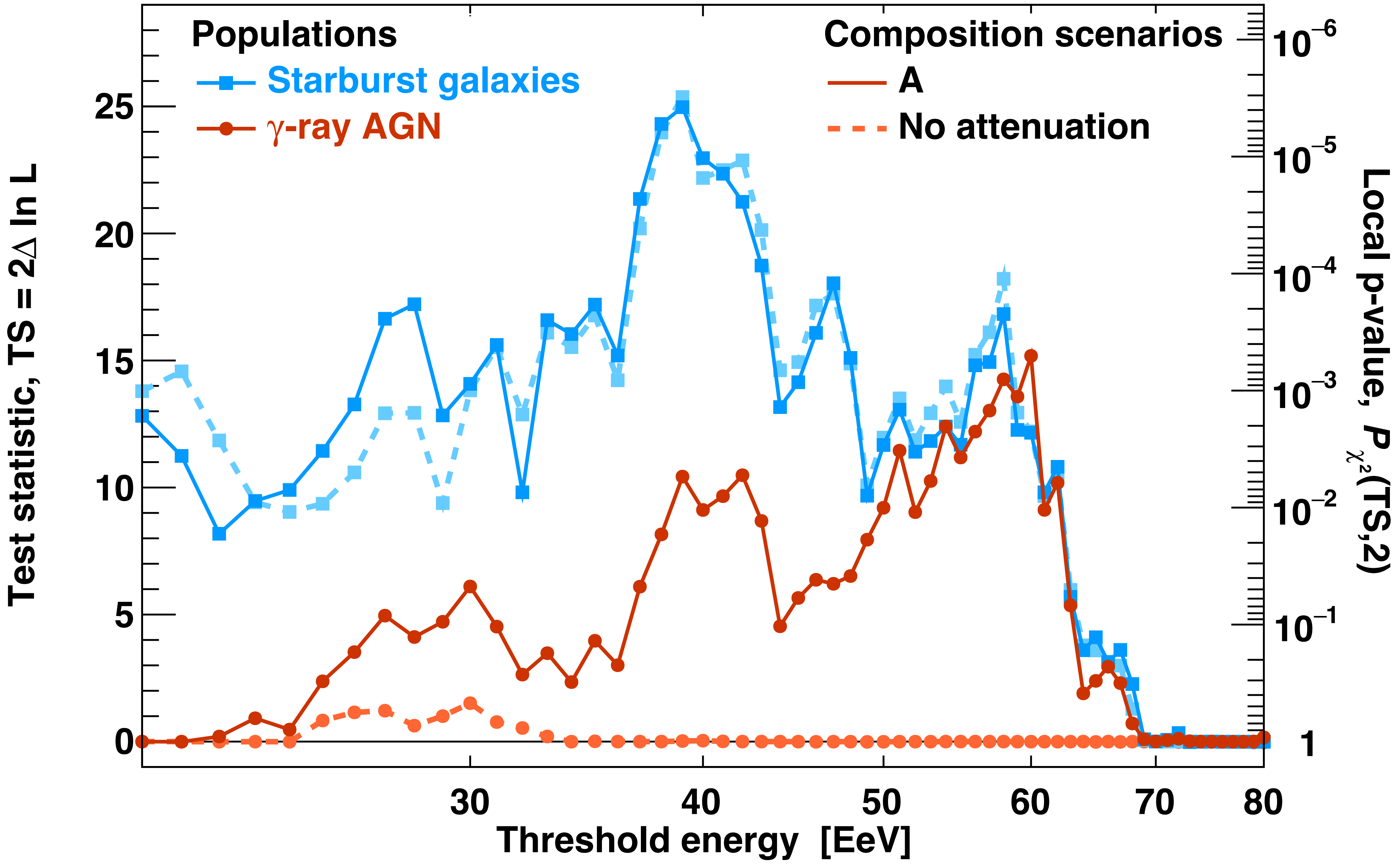
Two free parameters:
smearing angle, anisotropic fraction

Starburst galaxies - E > 39 EeV



Active galactic nuclei - E > 60 EeV





Result: SBG-model fits data better than isotropy at 4σ

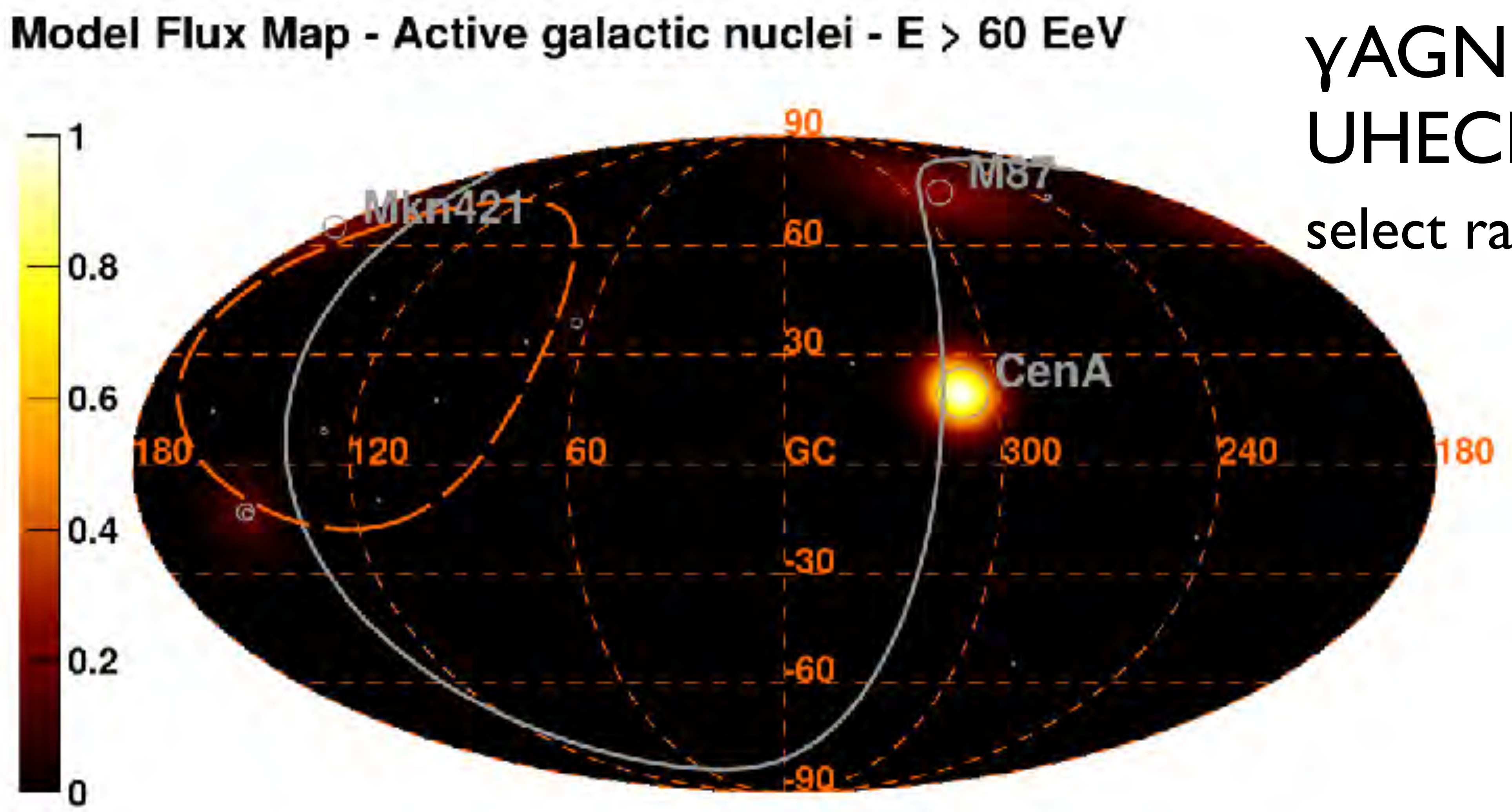
Result: γAGN-model fits data better than isotropy at 2.7σ

Auger: ApJL 853:L29 (2018)

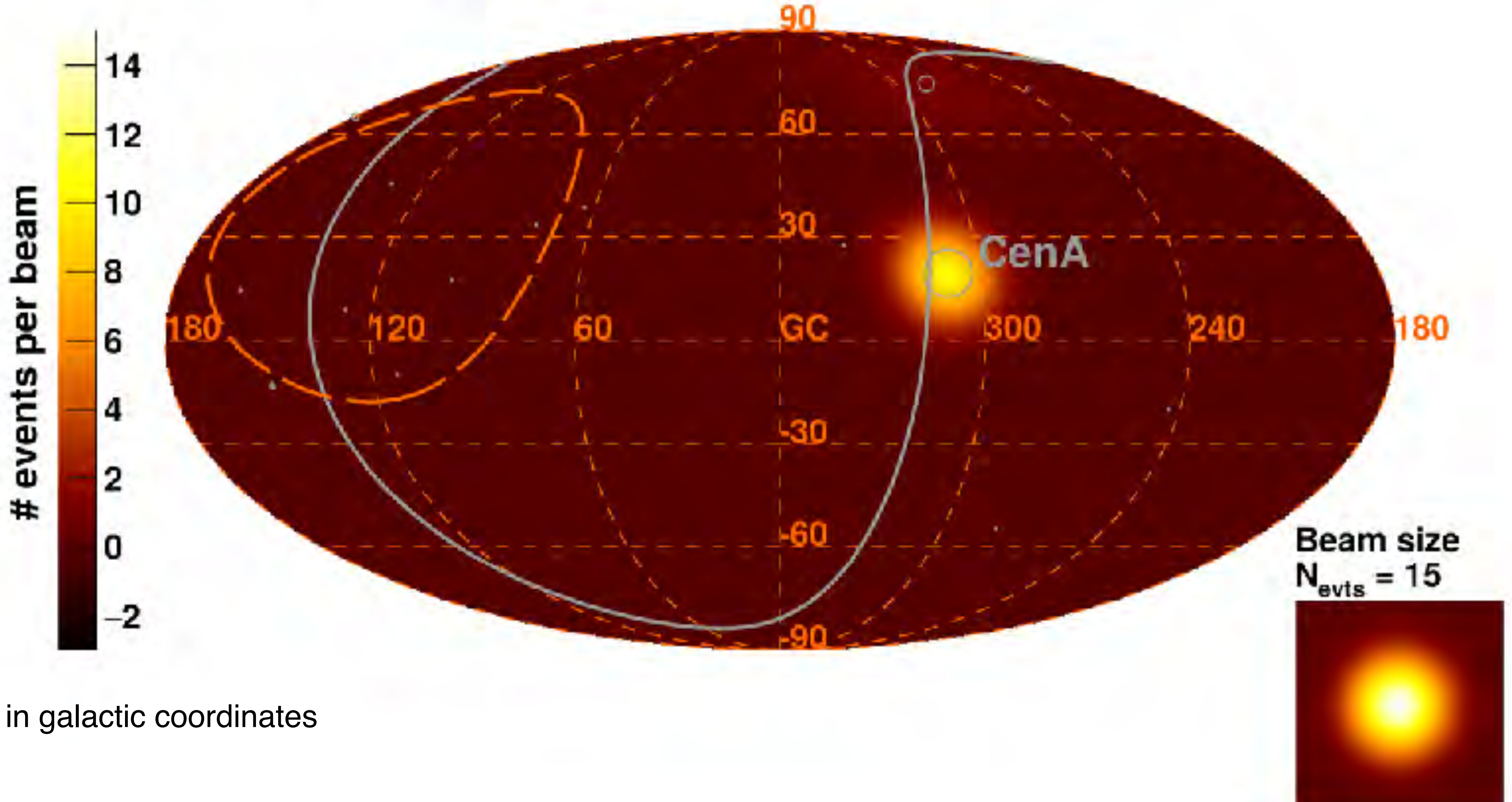
Understanding the UHECR Sky

γAGN Model

γAGN sky assuming
UHECR with $L_{UHECR} \sim L_\gamma$
select radio loud AGN within 250 Mpc

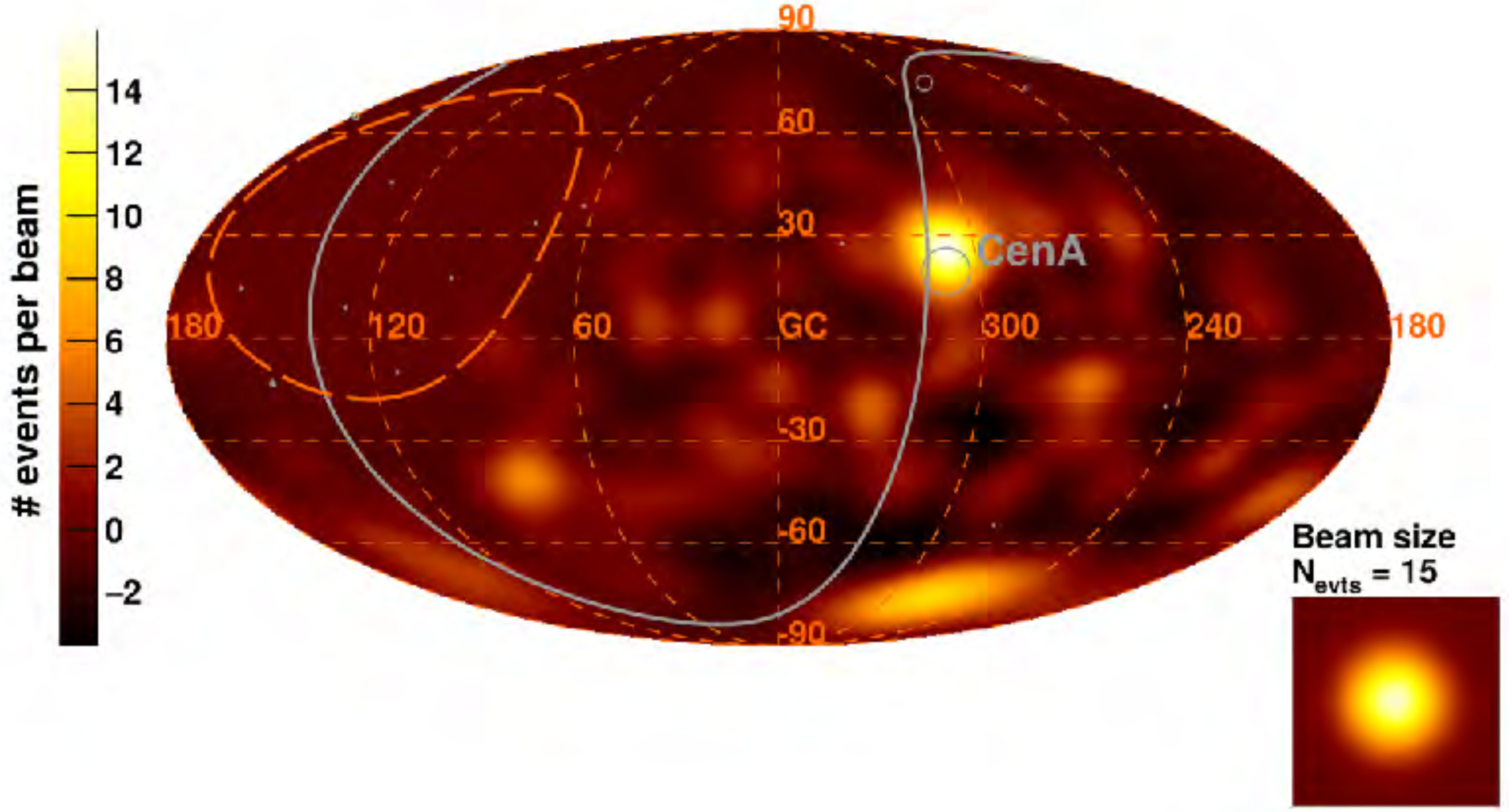


isotropic background added…



all in galactic coordinates

UHECR data map @ E> 60 EeV

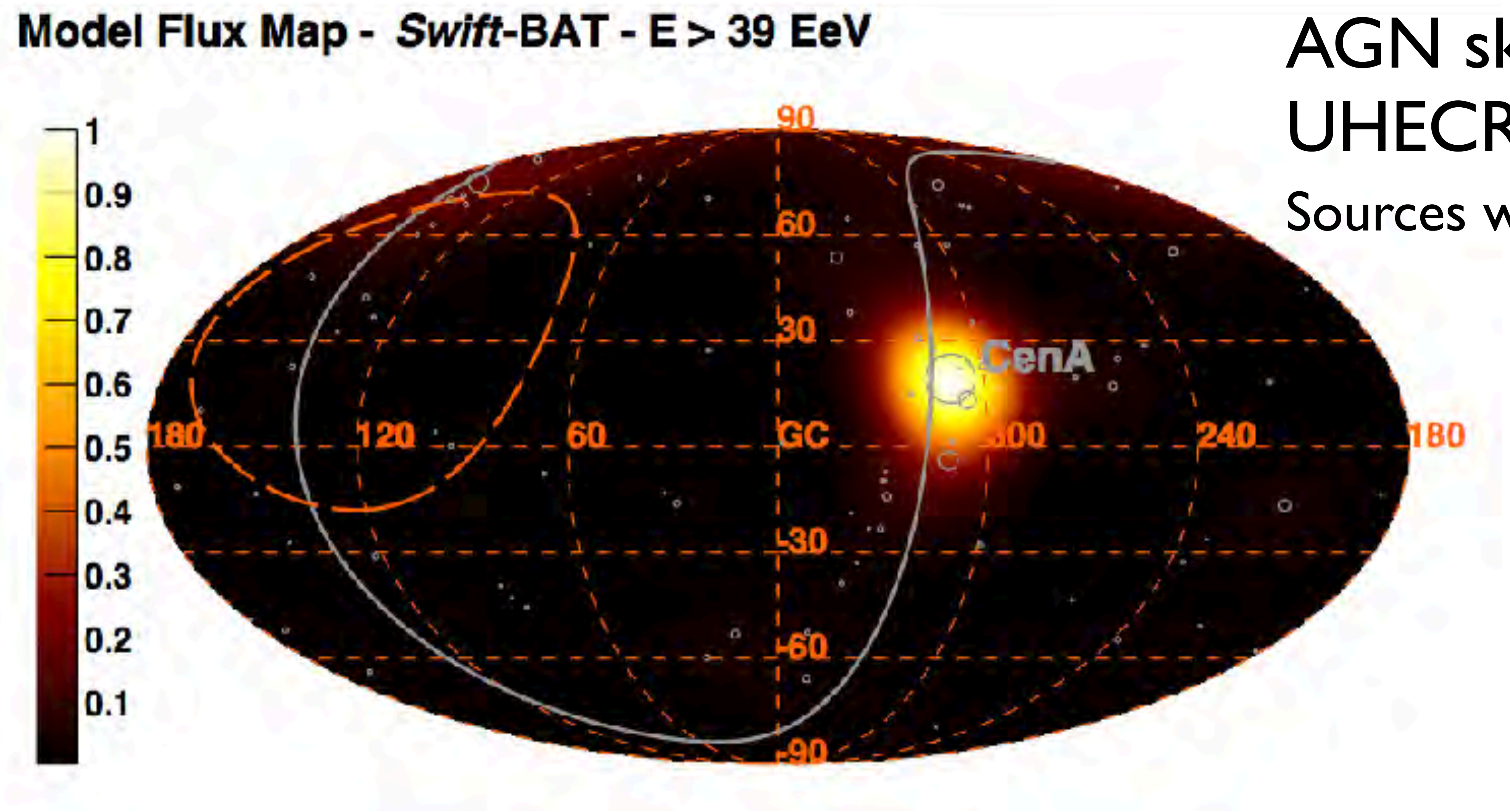


Result: γAGN-model fits data better than isotropy at 2.7σ

Auger: ApJL 853:L29 (2018)

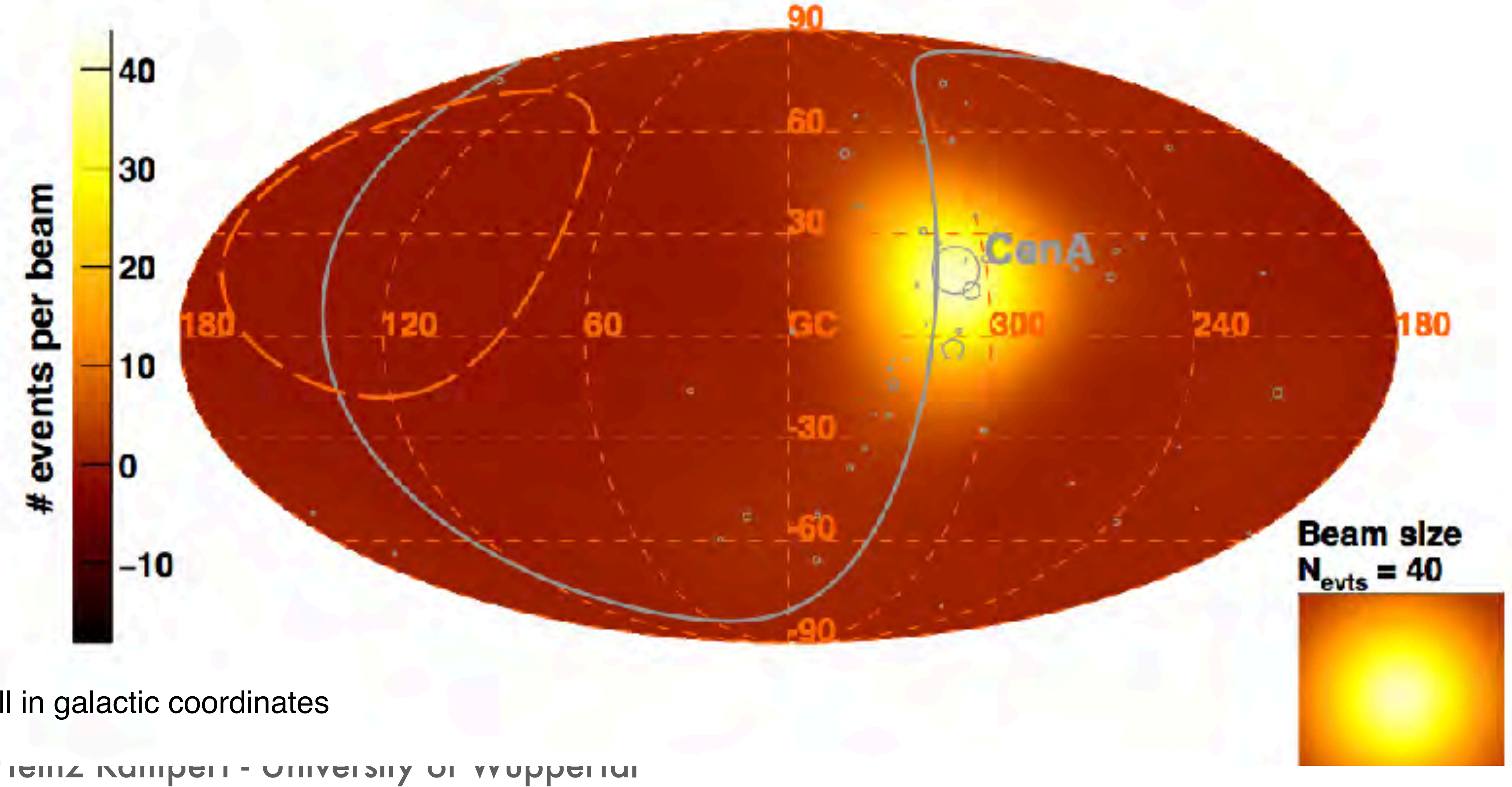
Understanding the UHECR Sky

Swift-BAT Model

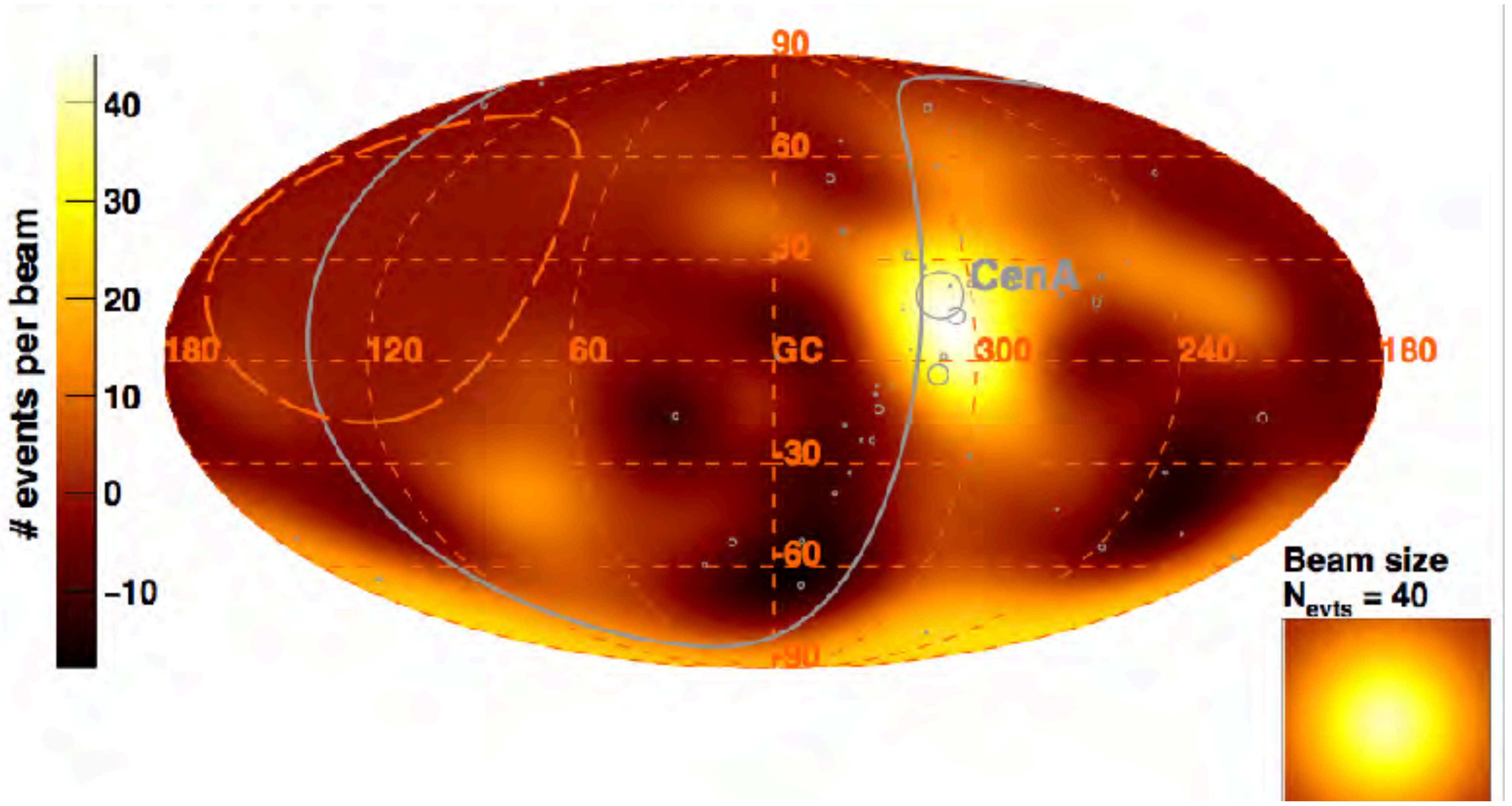


AGN sky assuming
UHECR with $L_{UHECR} \sim L_{\gamma}$
Sources with $F > 13.4 \times 10^{-12}$ erg/cm²/s within 250 Mpc

isotropic background added…



UHECR data map @ E> 39 EeV



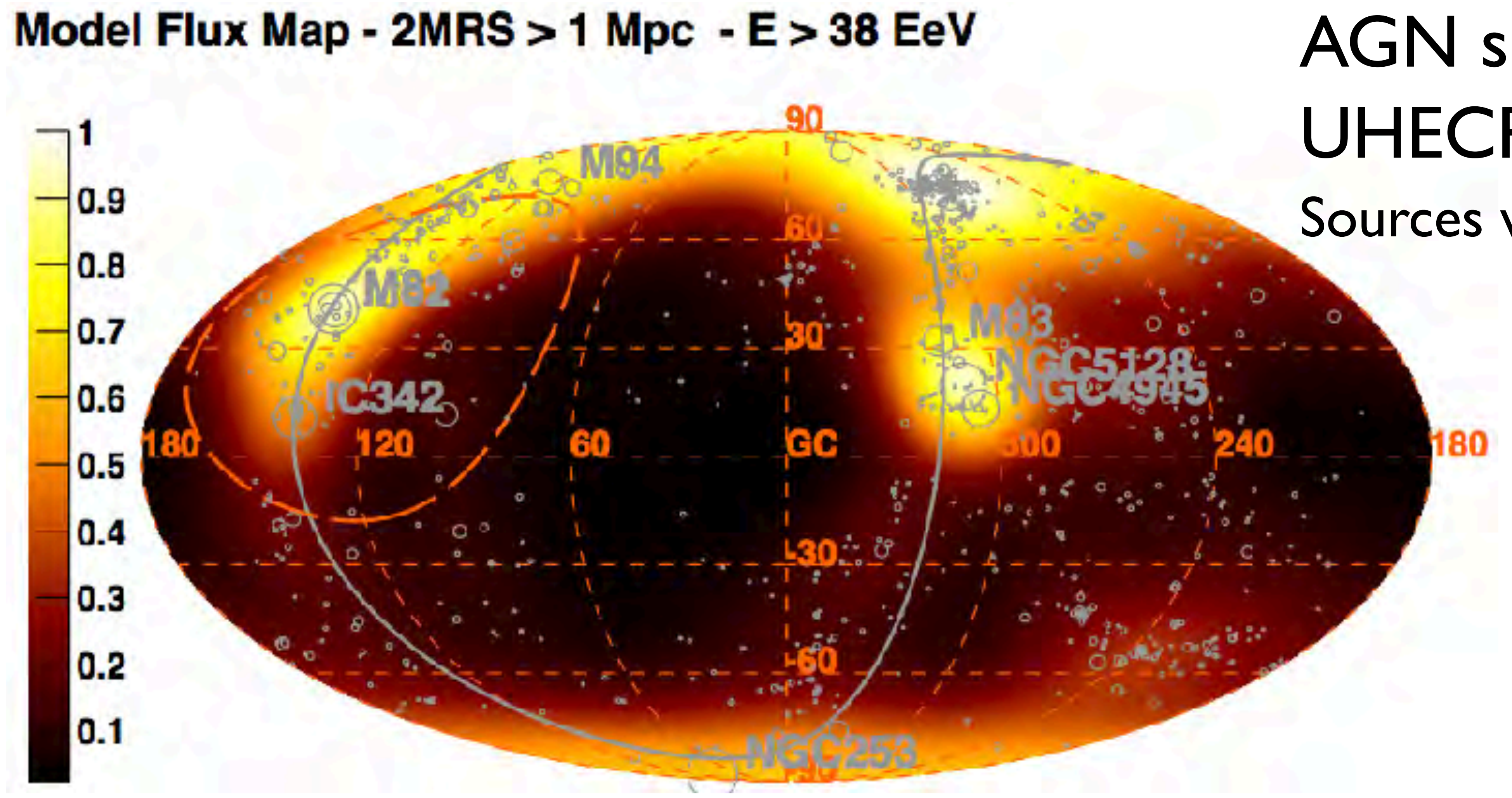
all in galactic coordinates

Result: Swif-BAT model fits data better than isotropy at 3.2σ

Auger: ApJL 853:L29 (2018)

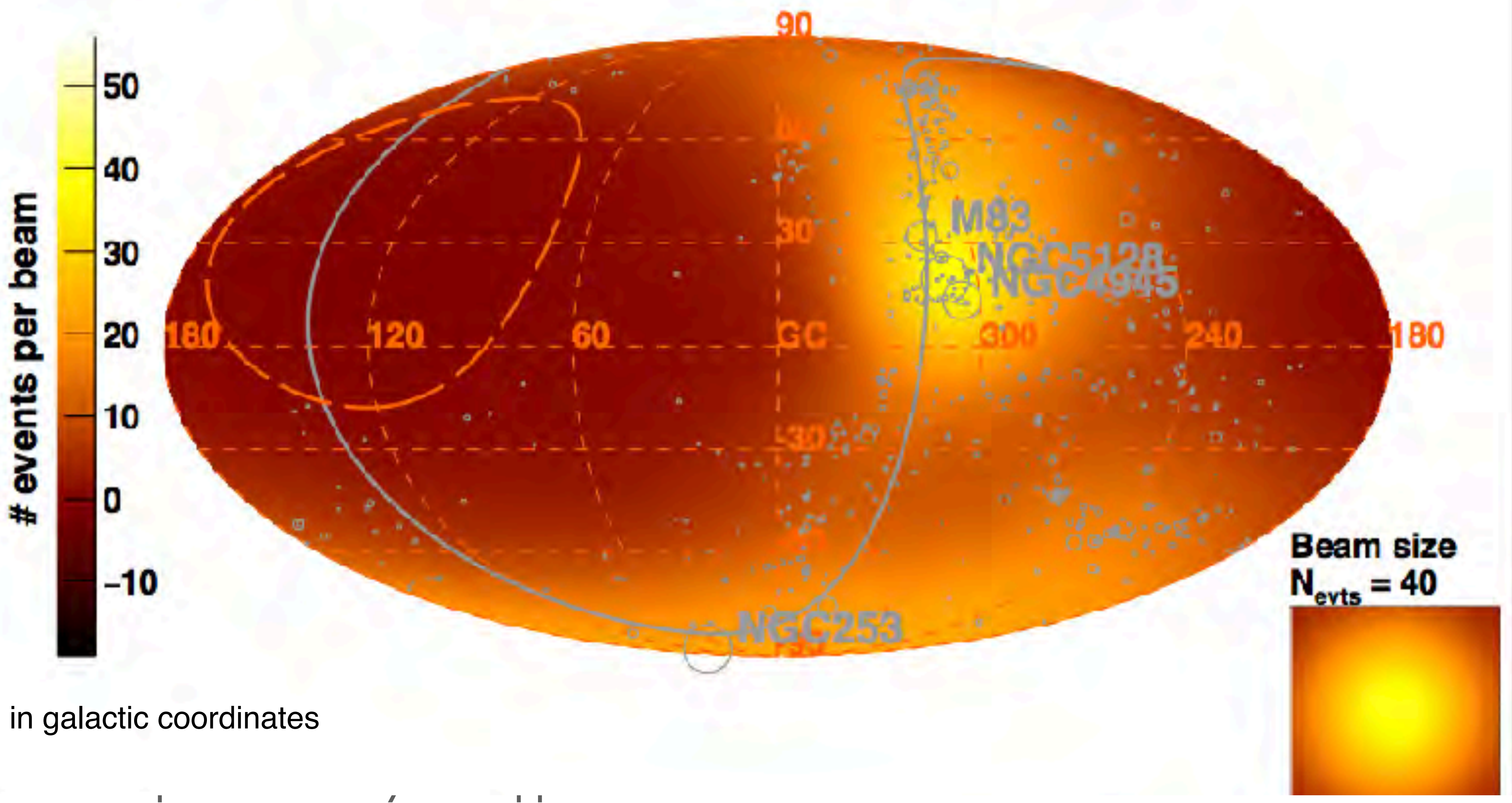
Understanding the UHECR Sky

2MRS Model

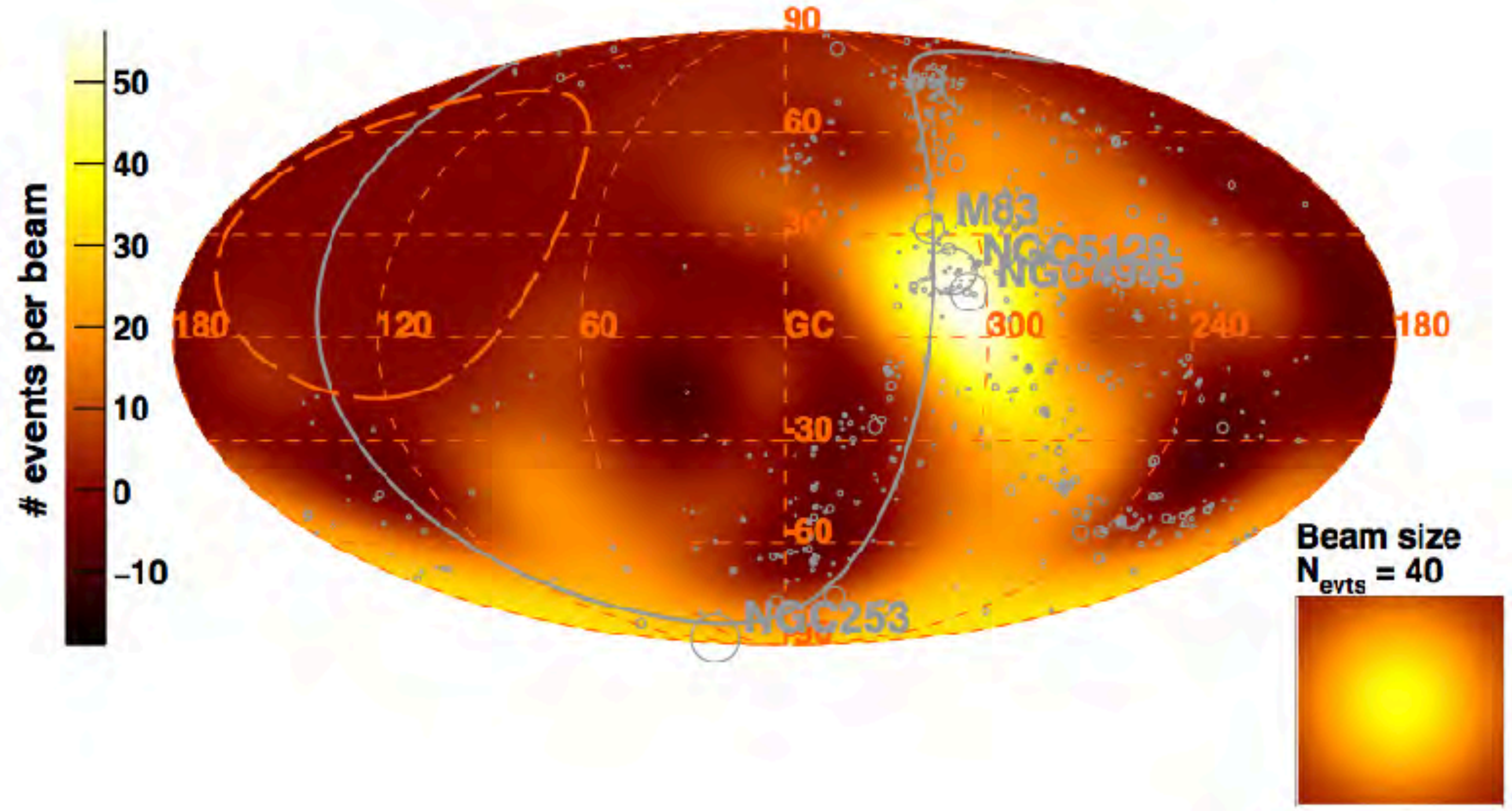


AGN sky assuming
UHECR with $L_{UHECR} \sim L_{K\text{-band}}$
Sources within 250 Mpc, exclude local group

isotropic background added…



UHECR data map @ E> 38 EeV



Result: 2MRS-model fits data better than isotropy at 2.7σ

Auger: ApJL 853:L29 (2018)

all in galactic coordinates

Understanding the UHECR Sky

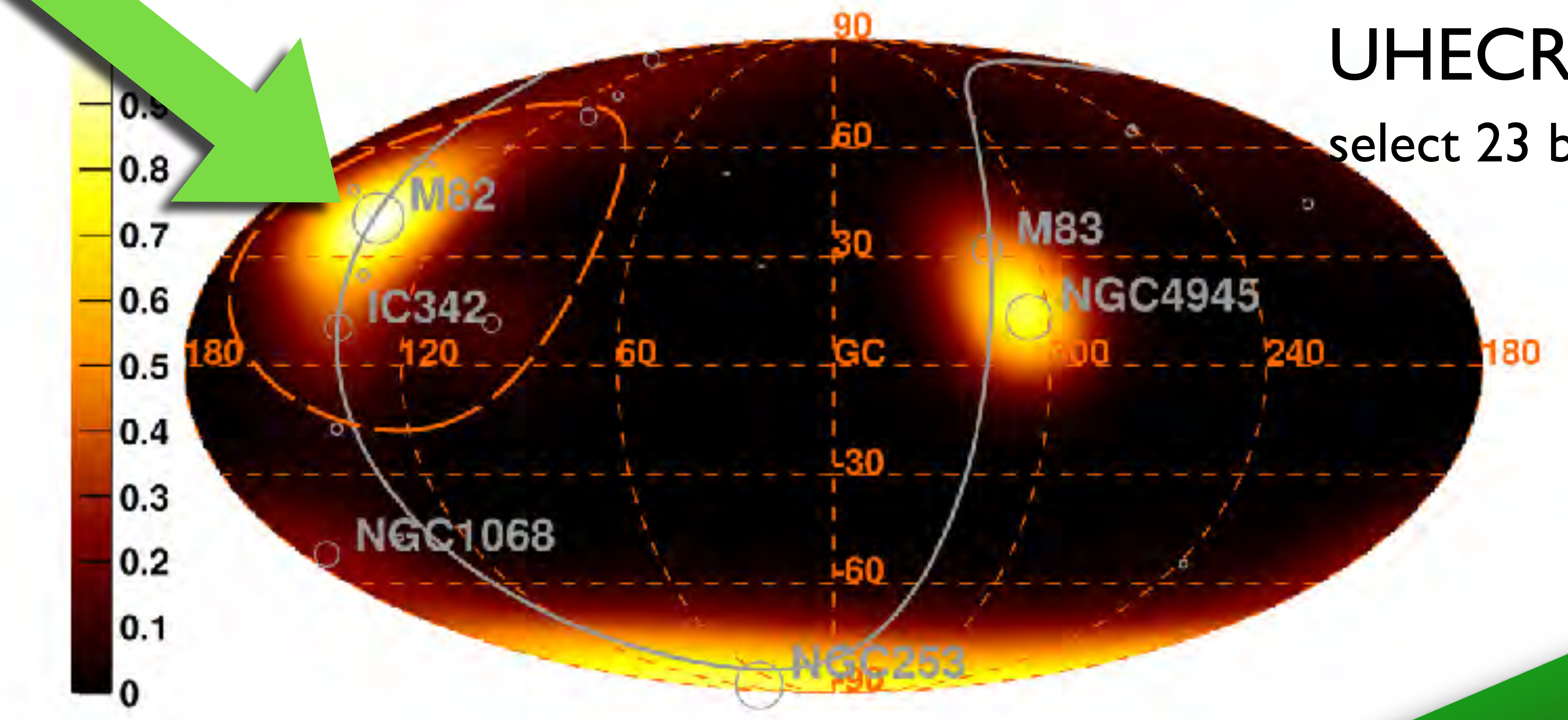
Starburst Galaxy Model

Assume, starburst galaxies produce UHECR with $L_{UHECR} \sim L_{\gamma\ @\ 1.4\ GHz}$

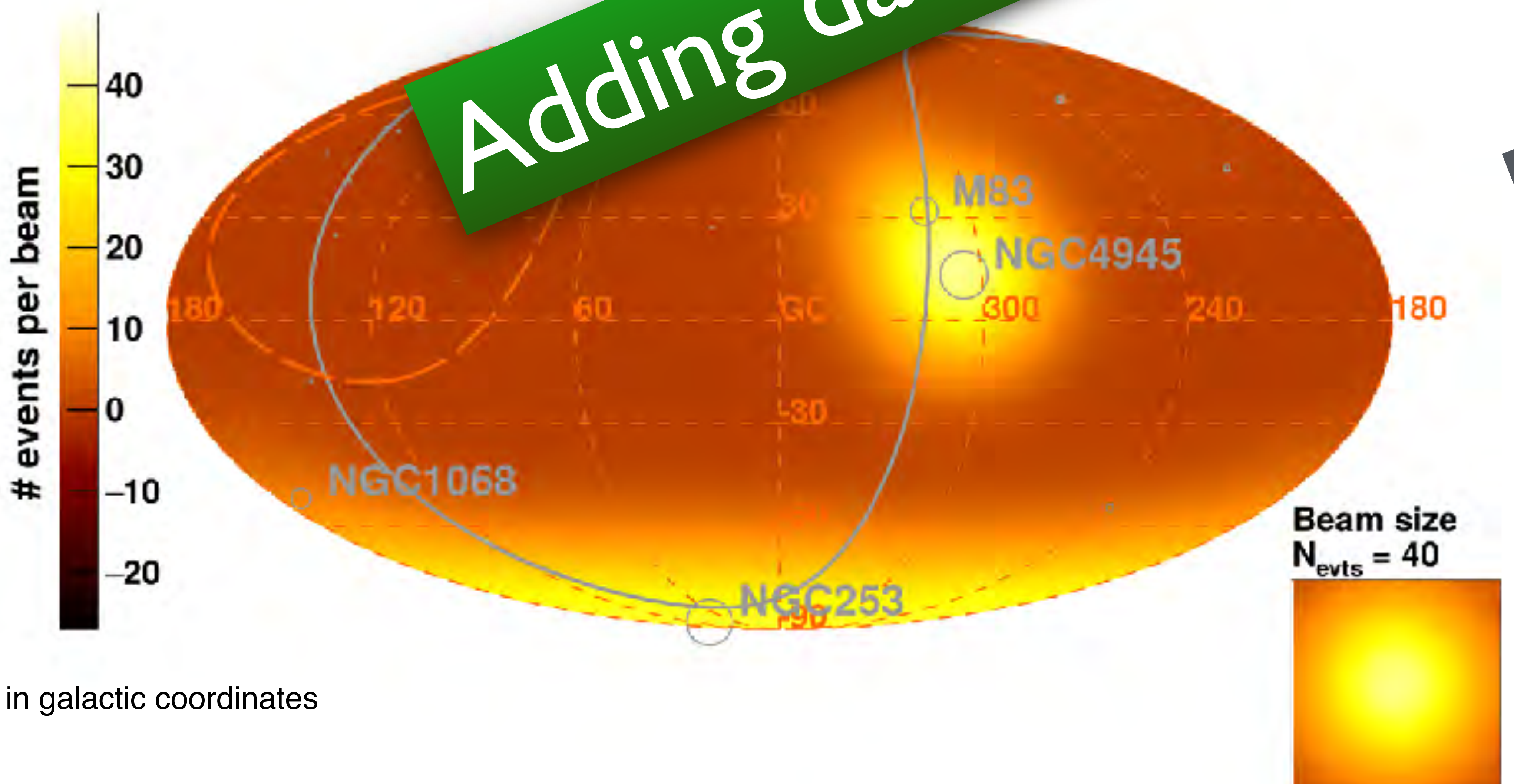
select 23 brightest SBG from Fermi-LAT catalogue and smear s

Adding data from Northern Sky (TA) next logical step

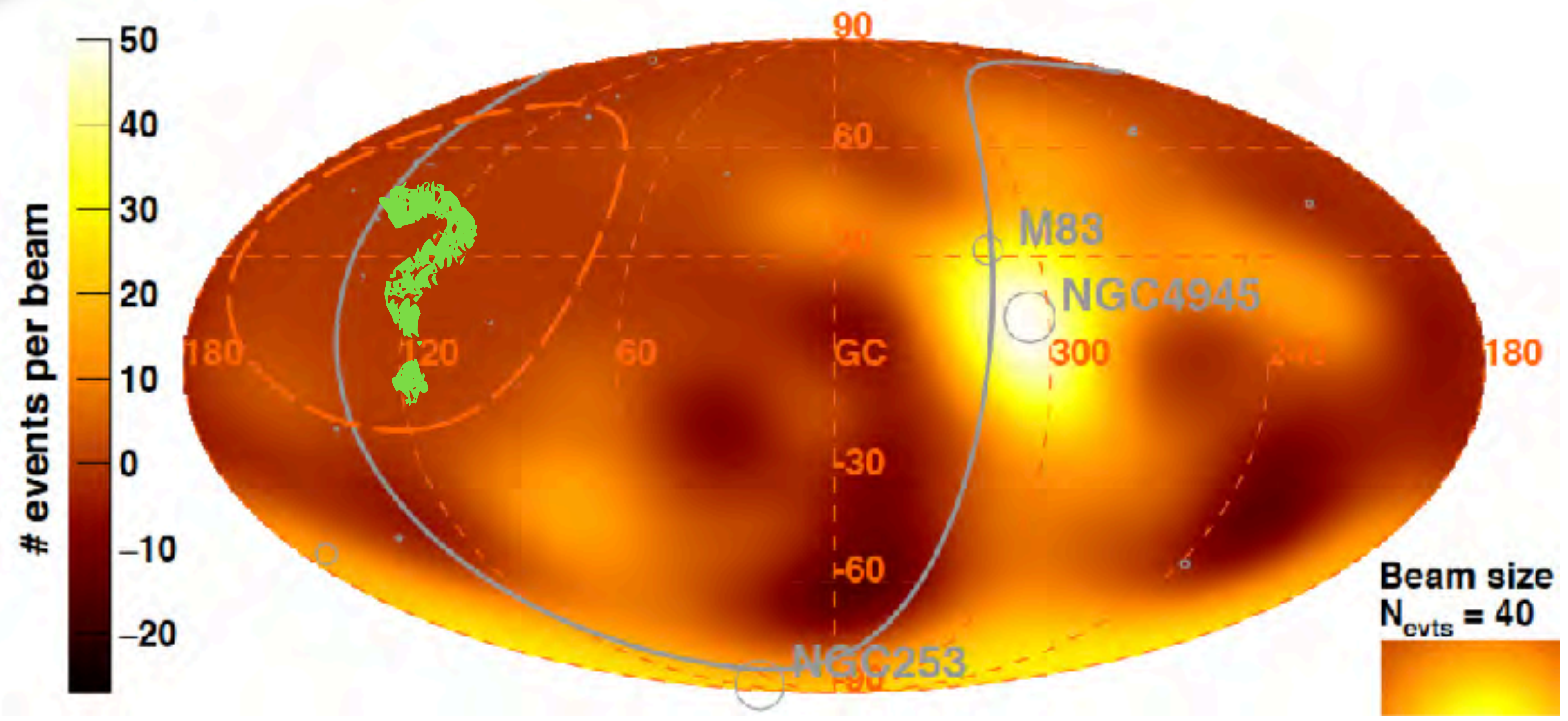
Model Flux Map - Starburst galaxies - E > 39 EeV



isotropic back... ...d...



UHECR data map @ E> 39 EeV

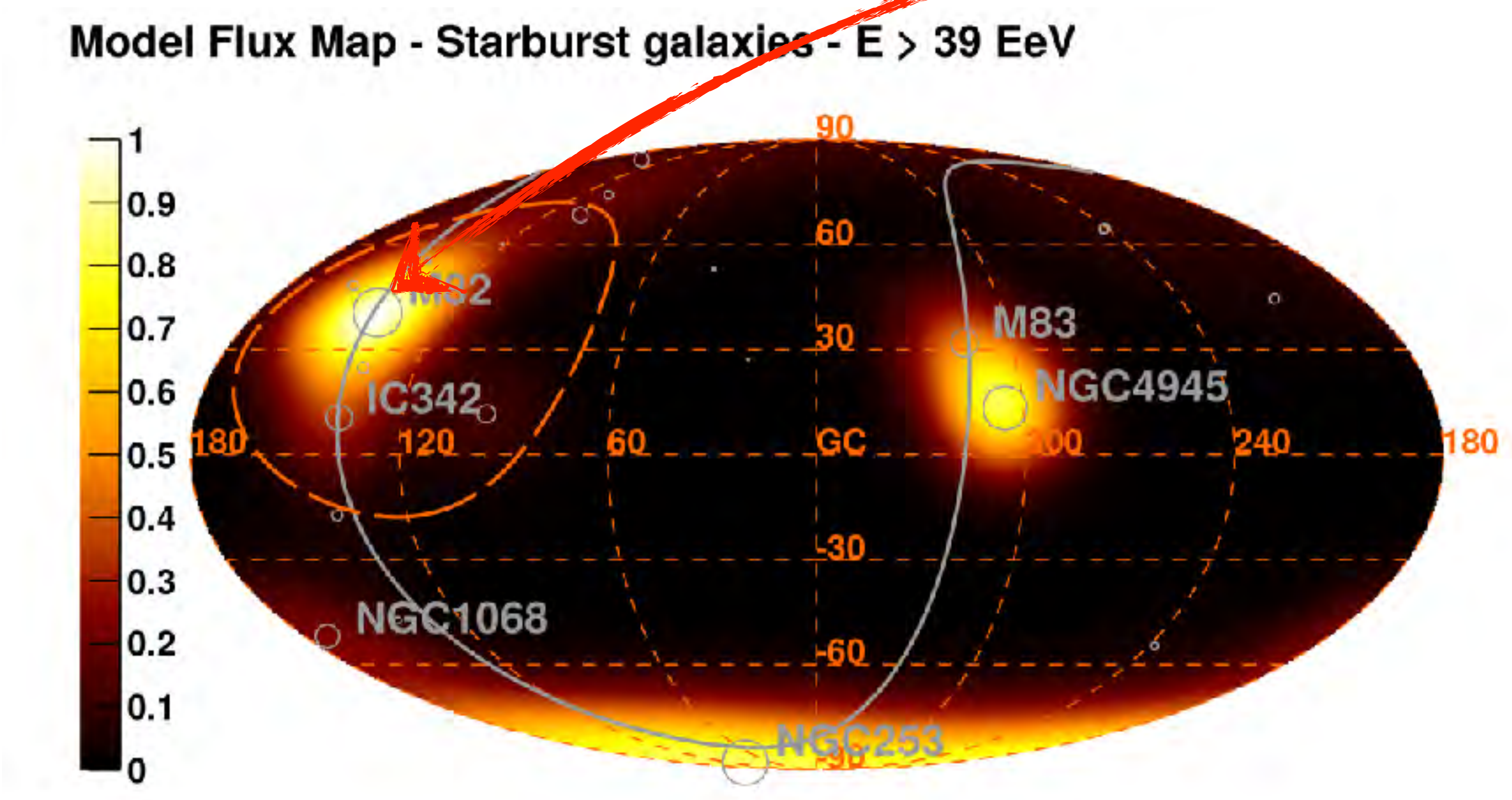


Result: 2MRS-model fits data better than isotropy at 4.0σ

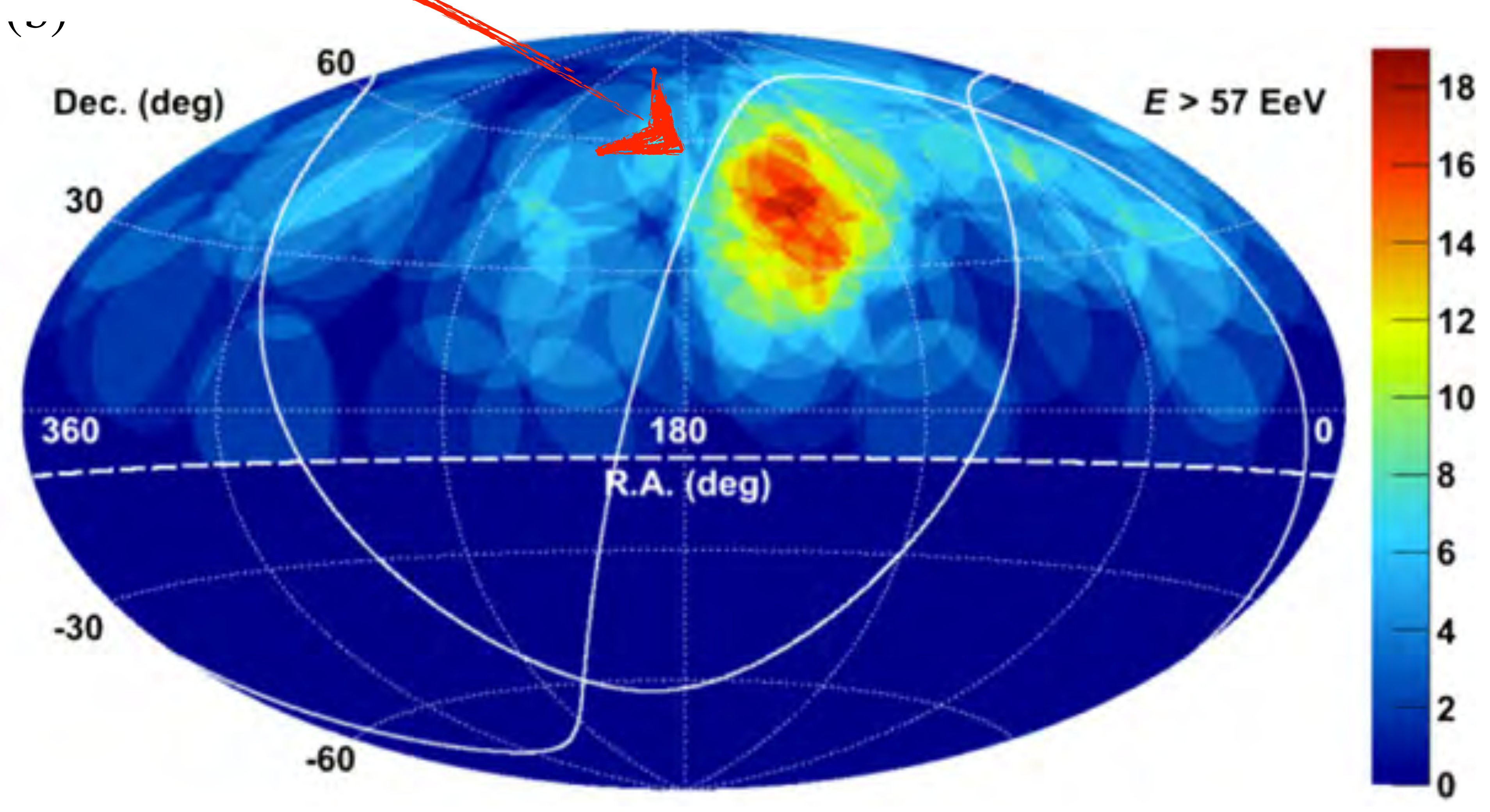
all in galactic coordinates

Auger: ApJL 853:L29 (2018)

TA Hot Spot (M82) may fit into the SBG (2MRS) picture



galactic coordinates



equatorial coordinates

TA-collaboration, ApJ 790:L21 (2014)

Summary of Results

Spectrum, composition, secondaries

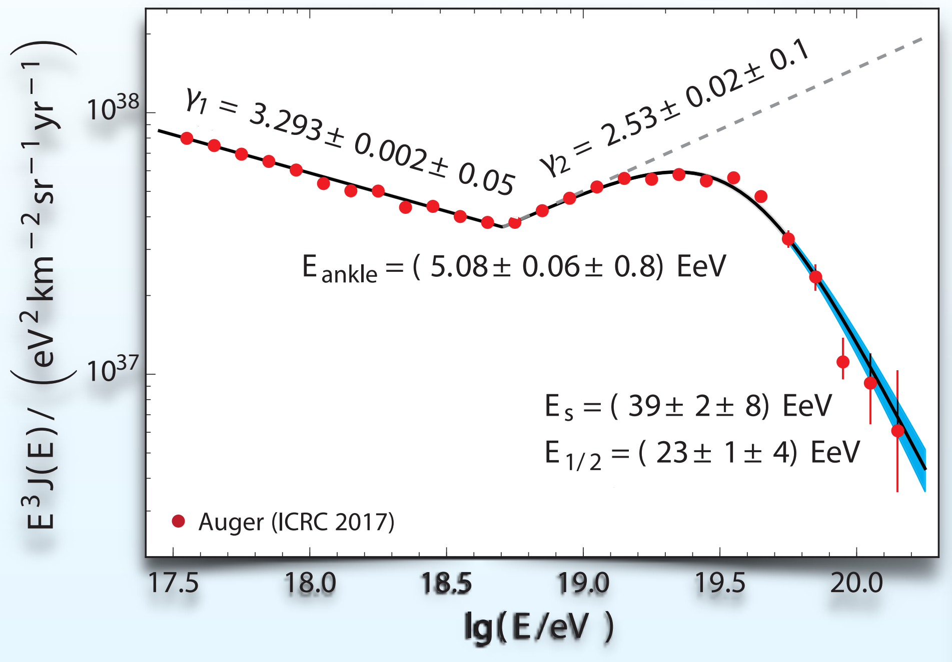
- Strong UHECR flux suppression observed
- Increasingly heavy mass composition observed:
  light at ankle, mixed at UHE
- Proton dominated sources constrained by ν & γ
- Data compatible with rigidity-dependent E_{max}

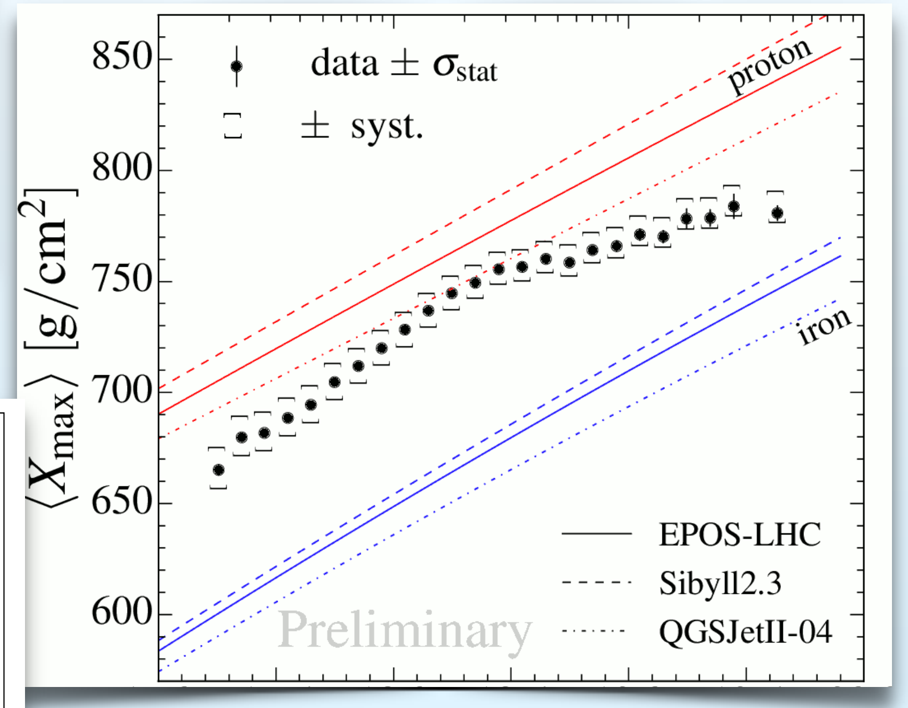
Hadronic Interactions

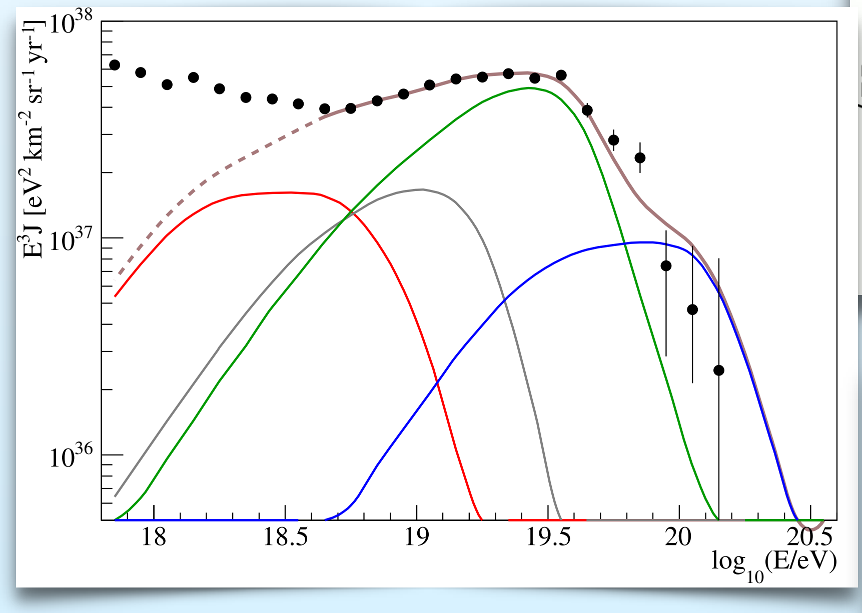
- standard UHE cross sections
- muon deficit in models

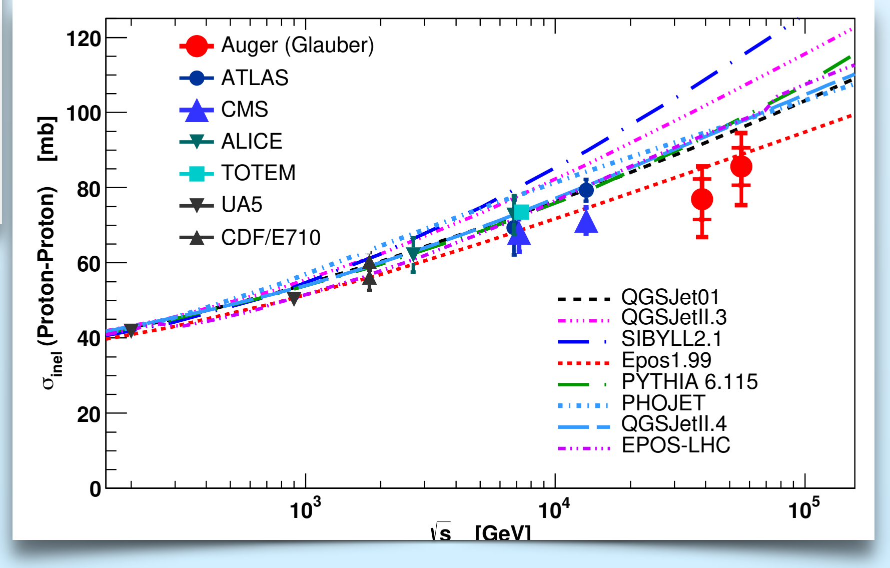
Anisotropies

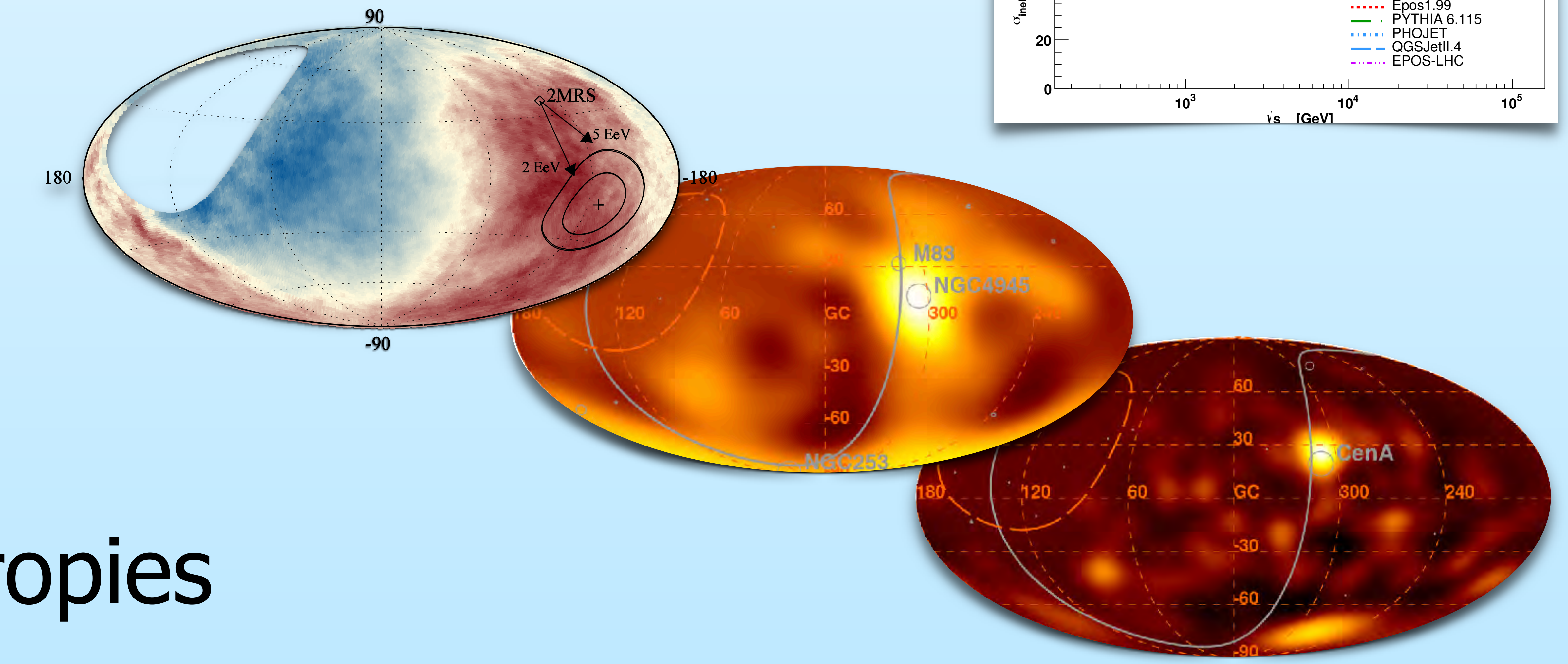
- Observation of dipole like anisotropy
- Indications for intermediate scale anisotropies







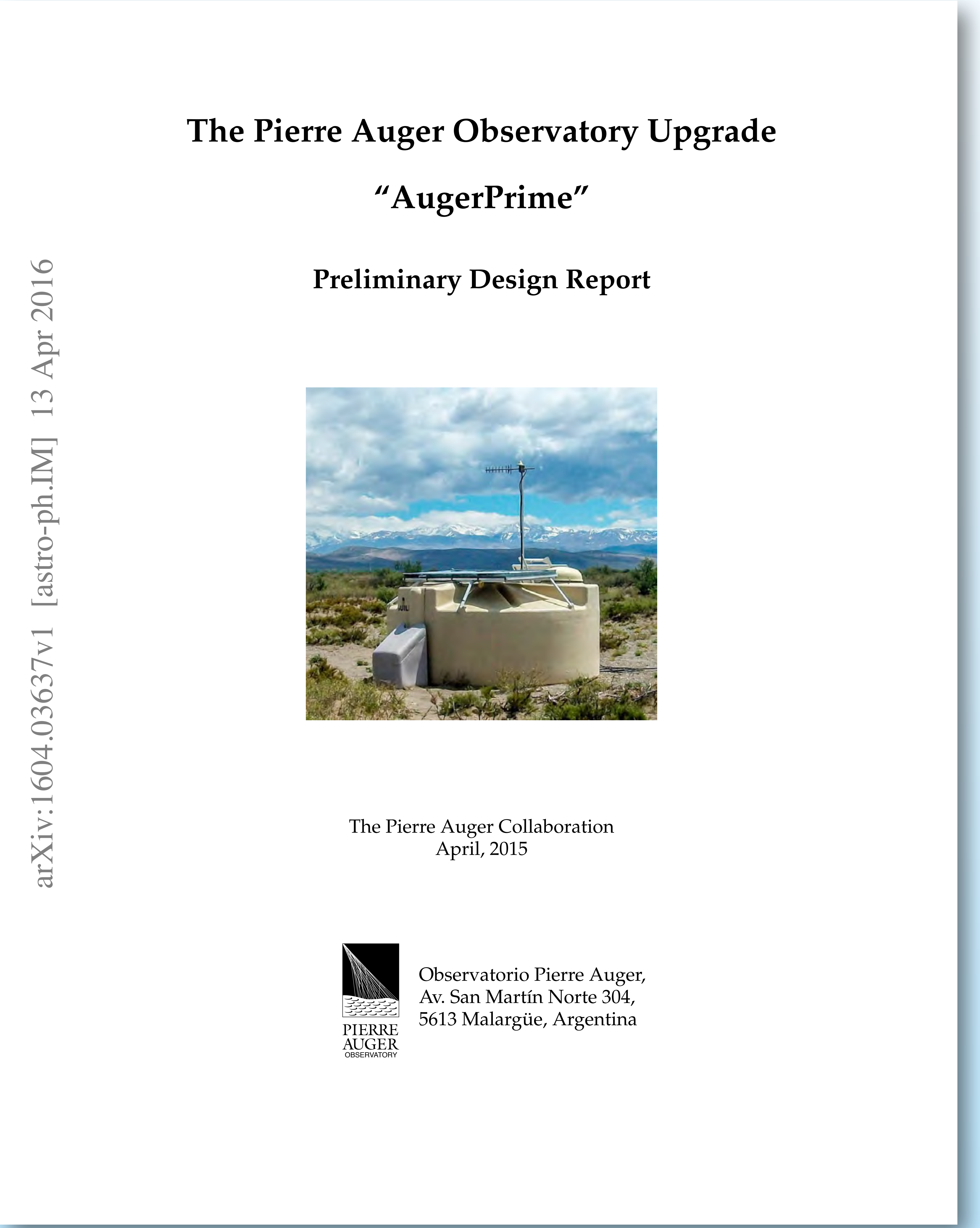




Open Key Questions

- Origin of the flux suppression ?
- Proton fraction at the highest energies ?
- Composition enhanced anisotropy
- Can we do proton astronomy ?
- Hadronic Physics above $\sqrt{s}$ = 140 TeV

→ **Need large exposure with composition sensitivity !**



The Pierre Auger Observatory Upgrade

"AugerPrime"

Preliminary Design Report

The Pierre Auger Collaboration
April, 2015

PIERRE AUGER OBSERVATORY

Observatorio Pierre Auger,
Av. San Martín Norte 304,
5613 Malargüe, Argentina

AugerPrime

Measure primary mass with 10 times better statistics

- 3.8 m^2 scintillators (SSD) on each 1500 m array stations improve e/μ discr.
- upgrade of station electronics
- additional small PMT to increase dynamic range
- buried muon counters in 750 m array (AMIGA)
- increased FD uptime

Scintillators on top of each Water Cherenkov Tank

(non invasive, fast to install, robust technology, relatively inexpensive)

Time line

- engineering array 09/2016
- construction has started
- deployment to start end 2018
- data taking into 2025

- ☑ do composition enhanced anisotropy
- ☑ p-astronomy: proof of principle
- ☑ particle physics beyond LHC
- ☑ ....

Extremely exciting times
with
great prospects for the near future!

Thank you !